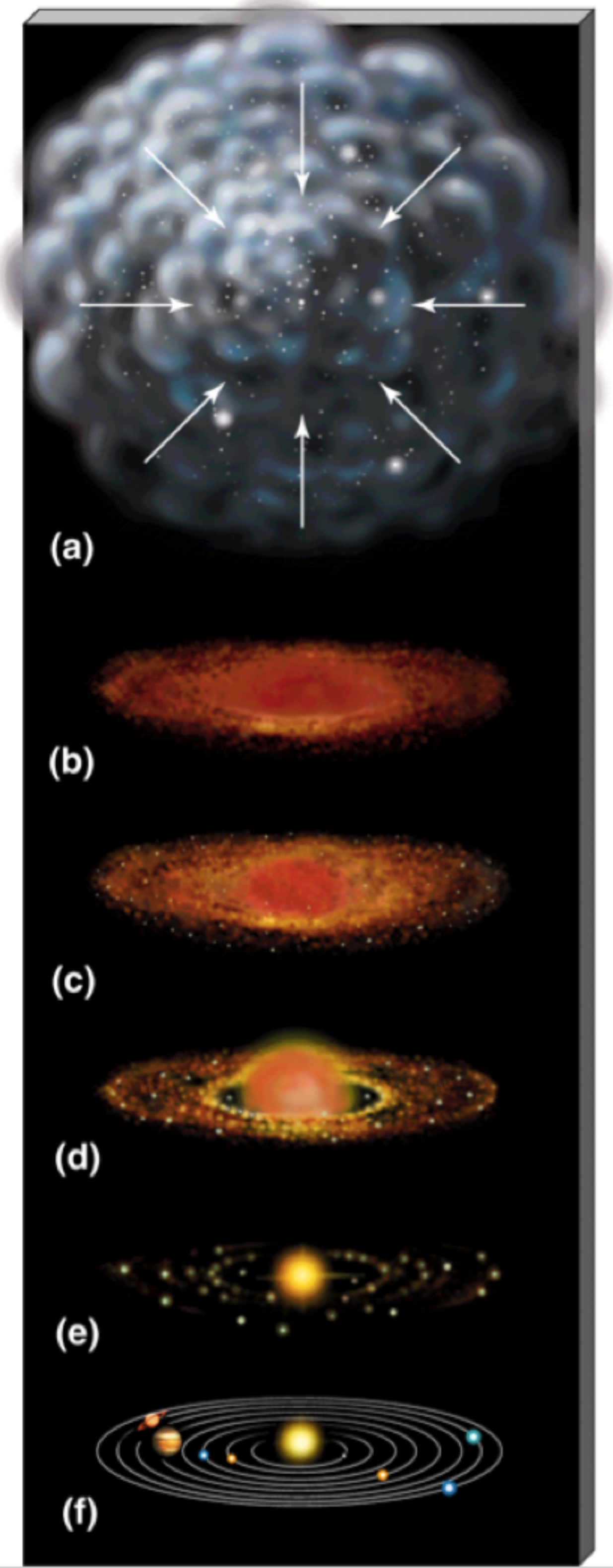
Gas in debris disks:
A new way to produce patterns?

Wlad Lyra
Sagan Fellow
Caltech-JPL

Marc Kuchner
NASA Goddard Space Flight Center

Nagoya University – March 2014



Collapse of gas cloud

A disk life story

Formation of proto-star

Gas-rich phase (< 10 Myr)

T-Tauri Disks

Accretion and Planet Formation

Dust settling

Thinning phase (~10 Myr)

Transitional Disks

Planetesimal formation

Planet retention

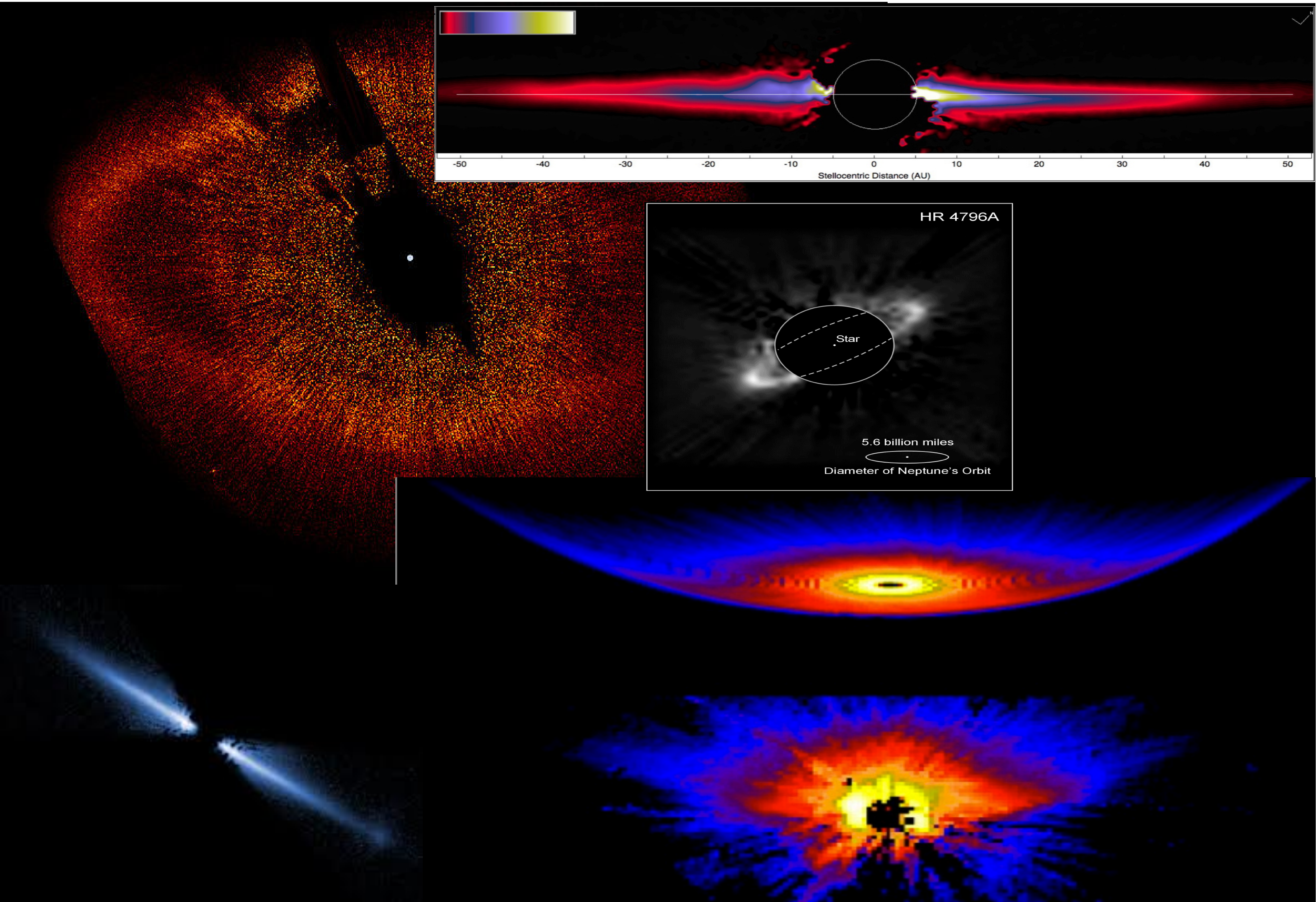
Gas dispersal

Gas-poor phase (>10 Myr)

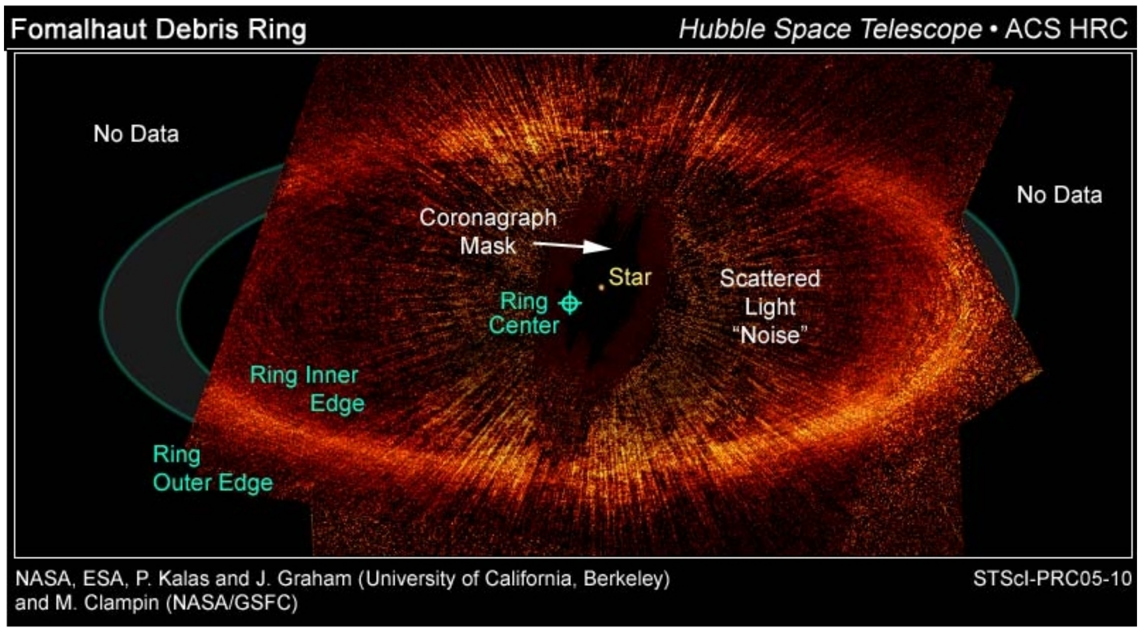
Debris Disks

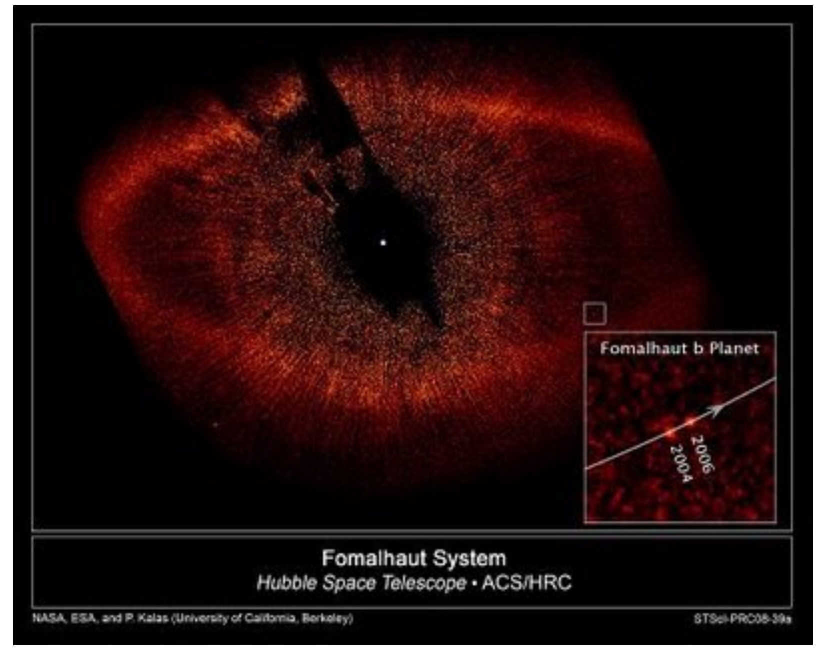
Stabilization of architecture and Planet Detection

Debris disks – The gas-poor phase



Sharp and eccentric rings in debris disks: Signposts of planets





Rings!
Narrow: 10AU wide, at 100 AU
Sharp: inner edge falls off abruptly
Eccentric: star and ring center do not coincide

Detection of a source quickly heralded as a planet Fomalhaut b

LETTER

doi:10.1038/nature12281

Formation of sharp eccentric rings in debris disks with gas but without planets

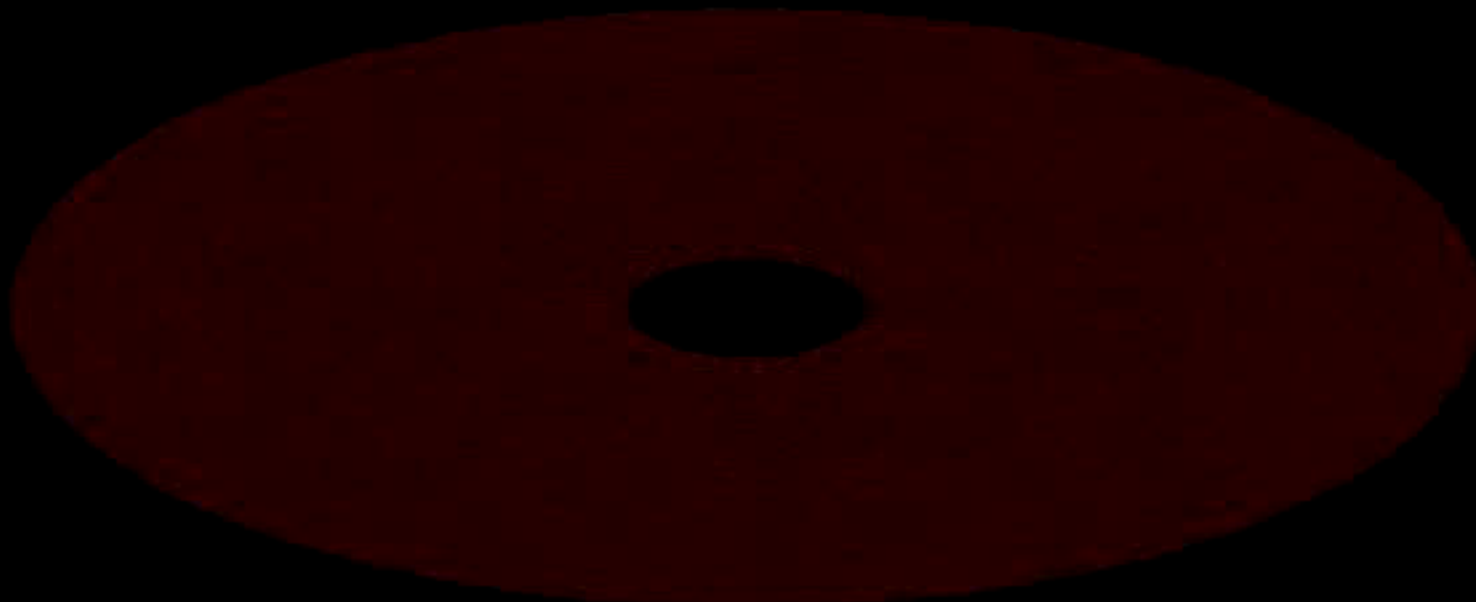
W. Lyra[1,2,3] & M. Kuchner[4]

'Debris disks' around young stars (analogues of the Kuiper Belt in our Solar System) show a variety of non-trivial structures attributed to planetary perturbations and used to constrain the properties of those planets[1–3]. However, these analyses have largely ignored the fact that some debris disks are found to contain small quantities of gas[4–9], a component that all such disks should contain at some level[10,11]. Several debris disks have been measured with a dust-to-gas ratio of about unity[4–9], at which the effect of hydrodynamics on the structure of the disk cannot be ignored[12,13]. Here we report linear and nonlinear modelling that shows that dust–gas interactions can produce some of the key patterns attributed to planets. We find a robust clumping instability that organizes the dust into narrow, eccentric rings, similar to the Fomalhaut debris disk[14]. The conclusion that such disks might contain planets is not necessarily required to explain these systems.

Disks around young stars seem to pass through an evolutionary phase when the disk is optically thin and the dust-to-gas ratio ε ranges from 0.1 to 10. The nearby stars β Pictoris[5,6,15–17], HD32297 (ref. 7), 49 Ceti (ref. 4) and HD 21997 (ref. 9) all host dust disks resembling ordinary debris disks and also have stable circumstellar gas detected in molecular CO, Na I or other metal lines; the inferred mass of gas ranges from lunar masses to a few Earth masses (Supplementary Information). The gas in these disks is thought to be produced by planetesimals or dust grains

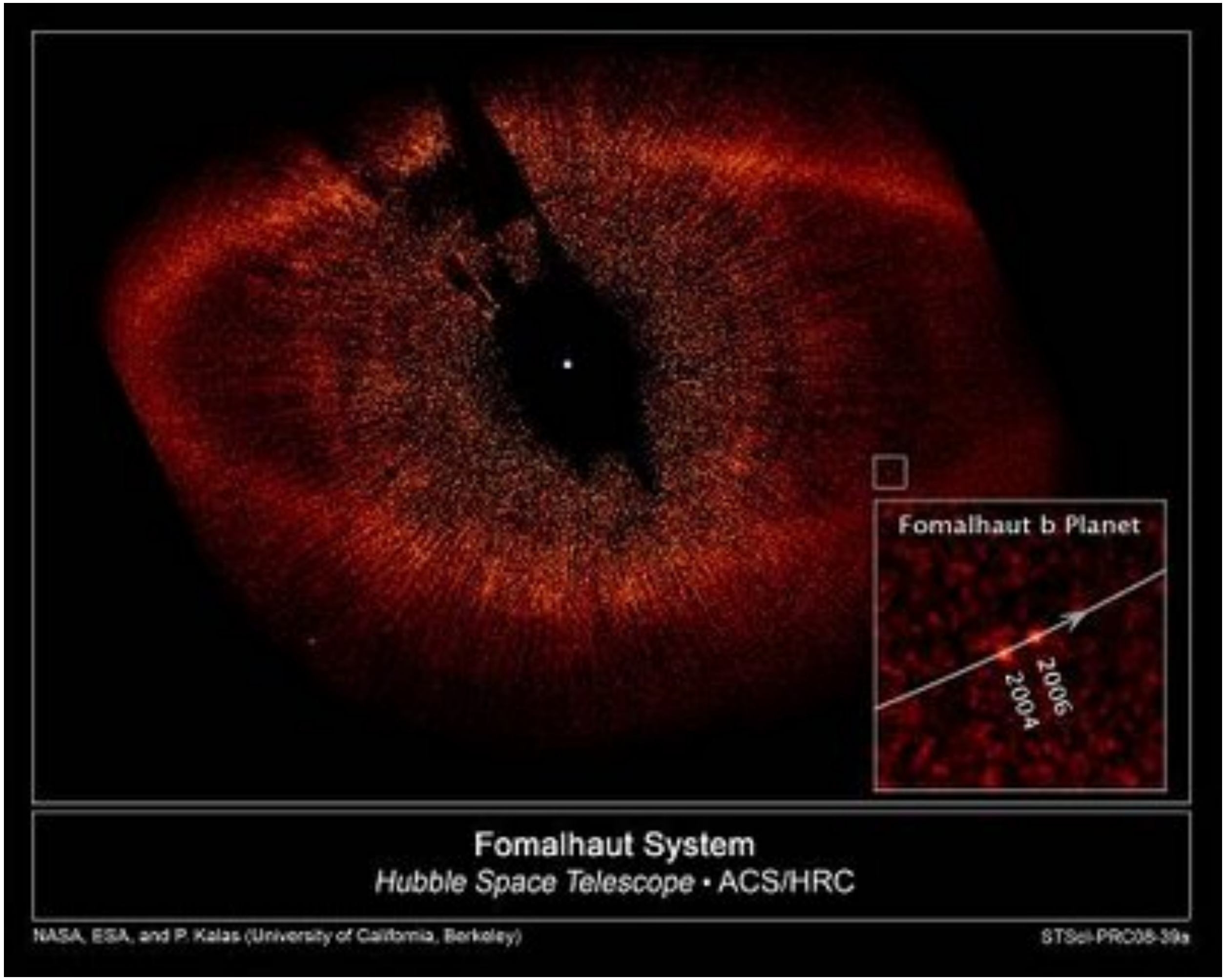
We present simulations of the fully compressible problem, solving for the continuity, Navier–Stokes and energy equations for the gas, and the momentum equation for the dust. Gas and dust interact dynamically through a drag force, and thermally through photoelectric heating. These are parametrized by a dynamical coupling time τ_f and a thermal coupling time τ_T (Supplementary Information). The simulations are performed with the Pencil Code[21–24], which solves the hydrodynamics on a grid. Two numerical models are presented: a three-dimensional box embedded in the disk that co-rotates with the flow at a fixed distance from the star; and a two-dimensional global model of the disk in the inertial frame. In the former the dust is treated as a fluid, with a separate continuity equation. In the latter the dust is represented by discrete particles with position and velocities that are independent of the grid.

We perform a stability analysis of the linearized system of equations that should help interpret the results of the simulations (Supplementary Information). We plot in Fig. 1a–c the three solutions that show linear growth, as functions of ε and $n = kH$, where k is the radial wavenumber and H is the gas scale height ($H = c_s/\sqrt{\gamma}\Omega_K$, where c_s is the sound speed, Ω_K the Keplerian rotation frequency and γ the adiabatic index). The friction time τ_f is assumed to be equal to $1/\Omega_K$. The left and middle panels show the growth and damping rates. The right panels show the oscillation frequencies. There is no linear insta-



Lyra & Kuchner (2013, Nature, 499, 184)

Fom b



Moves like a planet, reflects light.

but...

it's

TOO BIG

to be a planet
(several Jupiter radii)

Some of the Fom b controversy

Janson et al. 2012

Variability by 0.7-0.8 mag in F606W band

Astrometric orbit not apsidally aligned with the ring

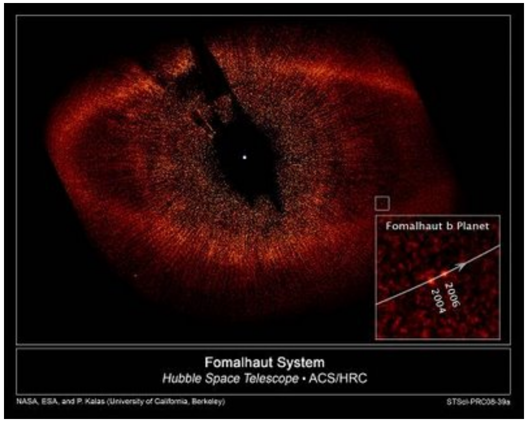
No infrared emission

Currie et al. 2012

No variability found within 0.15 mag in the same band

Consistent with apsidal alignment

Thermal emission from 0.5 MJ would not be detectable.



Kalas et al. 2013 … or …
"The Fom b guessing game"

9.3. Belt Collision Scenarios

In the coplanar scenario, Fomalhaut *b* is on a collision course with the main belt. Fomalhaut *b* will begin entering the inner edge of the dust belt around 2032 C.E., at which point the emergent phenomena would elucidate the physical nature of Fomalhaut *b*. For example, if Fomalhaut *b*'s optical light is due to a dust cloud, it may appear to episodically brighten and

9.3.2. Planet with a Satellite System

The circumplanetary dust disk hypothesis presented by Kalas et al. (2008) received a measure of plausibility with the discovery of Saturn's Phoebe ring at >200 R_p (Verbiscer et al. 2009). The basic physical mechanism is that the surface of a small (radius ~100 km), distant ($a = 215\ R_p$) planetary moon is bombarded by interplanetary meteoroids, launching ejecta that spirals toward

10. SUMMARY: IS IT A PLANET?

Our finding of a likely periastron passage near 30 AU radius now confers to Fomalhaut *b* a direct physical connection to the region where planetesimals grow to planets because the dynamical timescales are shorter and the primordial disk is denser closer to the star. On the other hand, compared to the present-day dynamics of the solar system, the orbit of

9.3.1. Planetesimal with a Dust Cloud

Here we assume that Fomalhaut *b* is a low-mass planetesimal that is optically bright because of reflected light from a fresh dust cloud surrounding it. For example, it could be a planetes-imal that was recently disrupted by forces associated with its

9.3.5. Recent Giant Impact as the Origin of the Main Belt

Extending the impact theme even further, is it possible that Fomalhaut *b* collided with a hypothetical second planet, Fomalhaut *c*, and the main belt is now the remnant debris of Fomalhaut *c*? Giant impacts that can produce transient circumstellar dust rings have recently been invoked to explain

Some of the Fom b controversy

Janson et al. 2012

Variability by 0.7-0.8 mag in F606W band

Astrometric orbit not apsidally aligned with the ring

No infrared emission

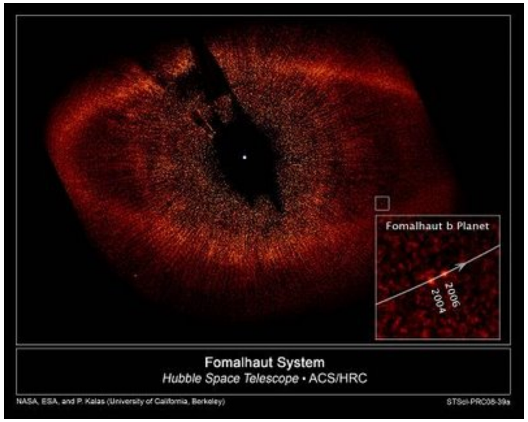
Currie et al. 2012

No variability found within 0.15 mag in the same band

Consistent with apsidal alignment

Thermal emission from 0.5 MJ would not be detectable.

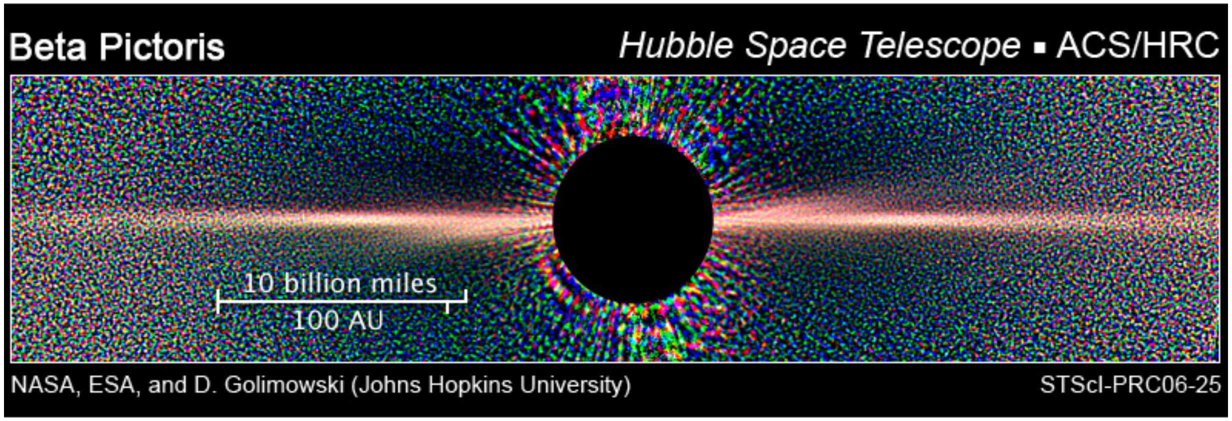
Observed optical emission requires reflection by ***something*** of several Jupiter radii



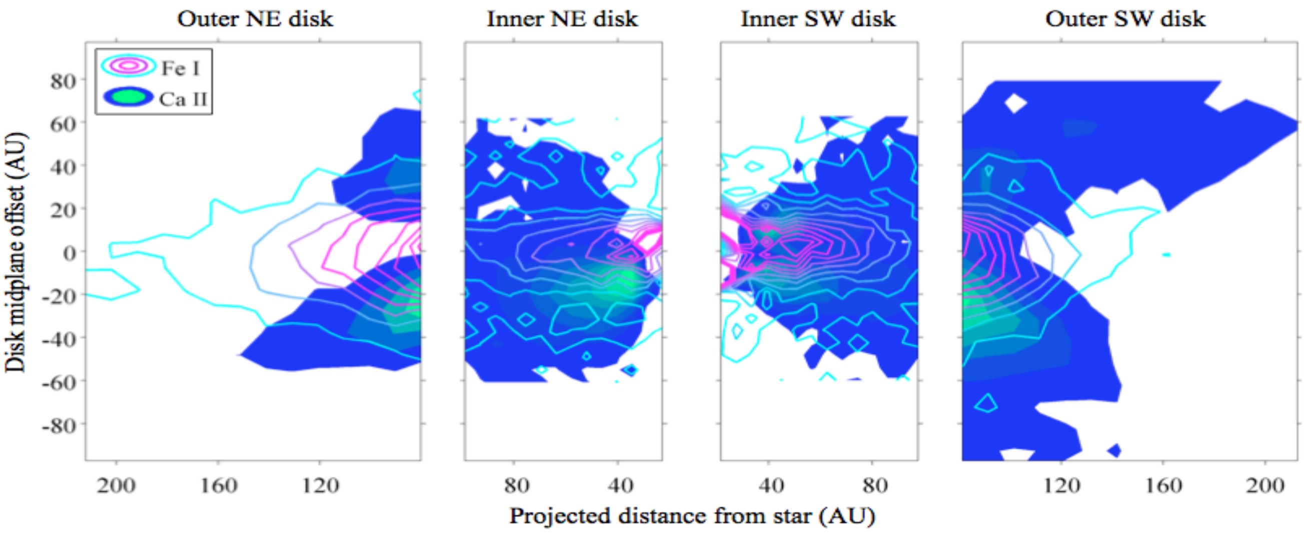
Are there alternative explanations?

Debris disks are not completely gas-free

Dust



Gas



VLT imaging by
Nilsson et al. (2012)

Gas in debris disks

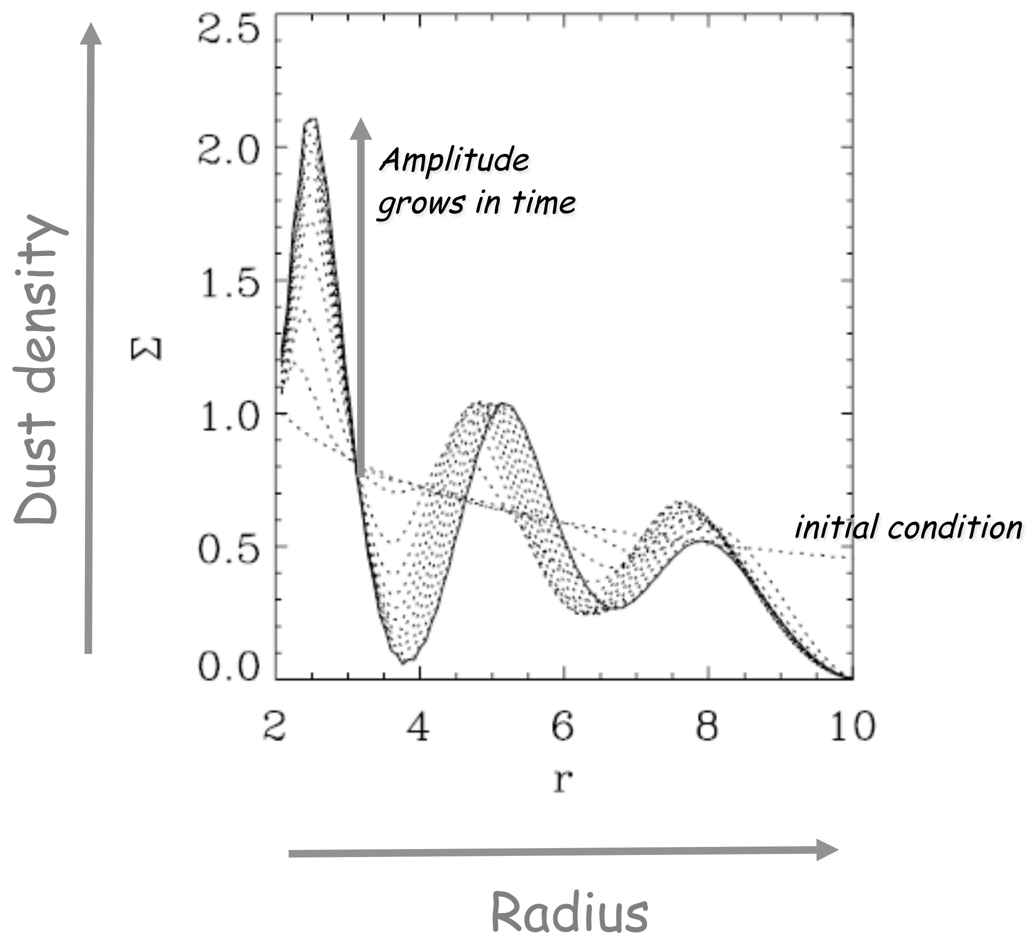
Detections

β Pictoris	many species	Lagrange et al. (1998), ...
51 Ophiuchi	many species	Roberge et al. (2002)
σ Herculis	C II, N II	Chen & Jura (2003)
HD 32297	Na I, CII	Redfield (2007), Donaldson et al. (2012)
HD 135344	H2, CO	Thi et al. (2001), Pontoppidan et al. (2008)
49 Ceti	H2, CO	Dent et al. (2005), Roberge et al. (2012)
AU Mic	H2	France et al. (2007)
HD172555	SiO	Lisse et al. (2009)

Source of gas: Outgassing processes

Infalling comets	Beust & Valiron (2007)
Grain sublimation	e.g. Rafikov (2012)
Grain-Grain collisions	Czechowski & Mann (2007)
Photo-stimulated desorption	Chen et al. (2007)
Planet-Planet collisions	Van den Ancker (2001), Lisse (2008)
Primordial?	

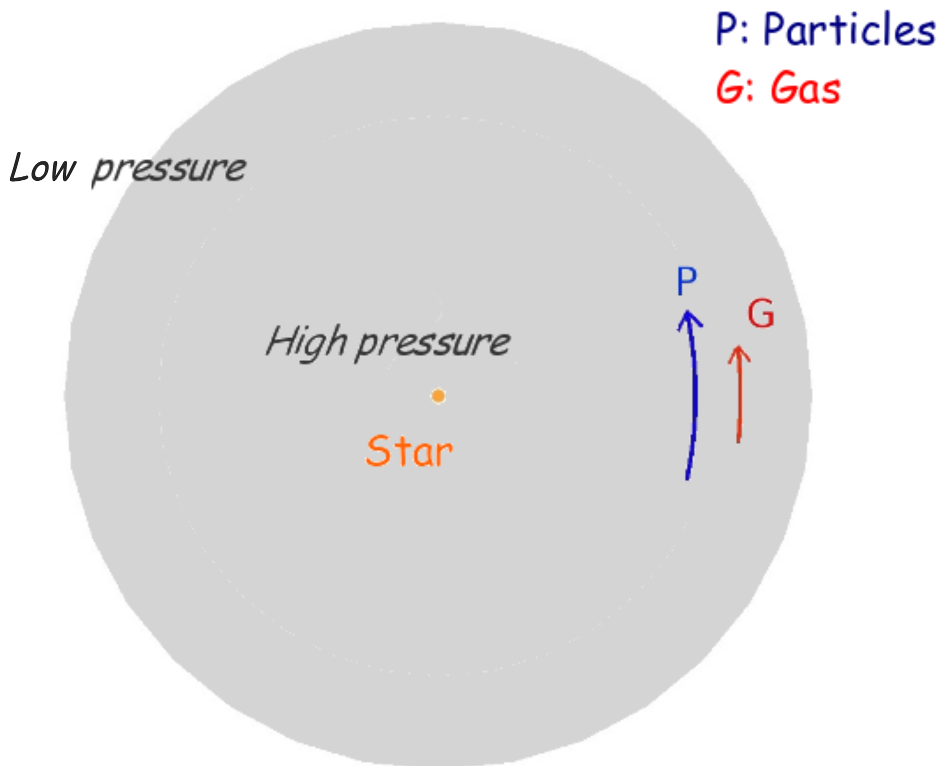
Dust and gas together leads to instability...



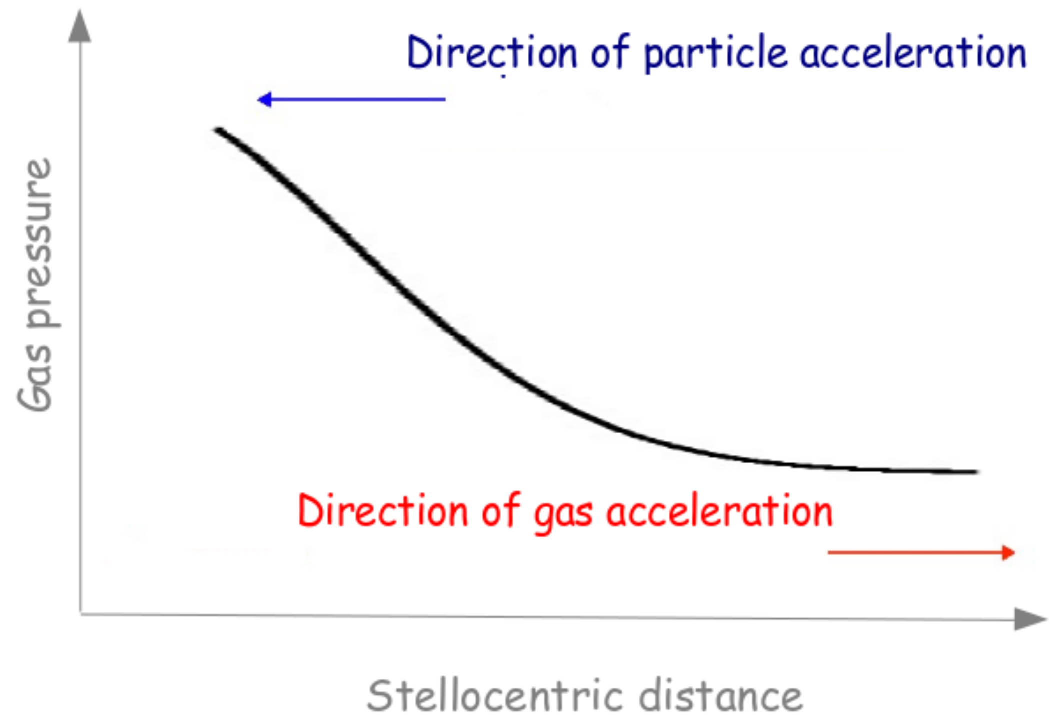
Klahr & Lin (2005)

Suggested that an instability causes dust in debris disks to clump together.

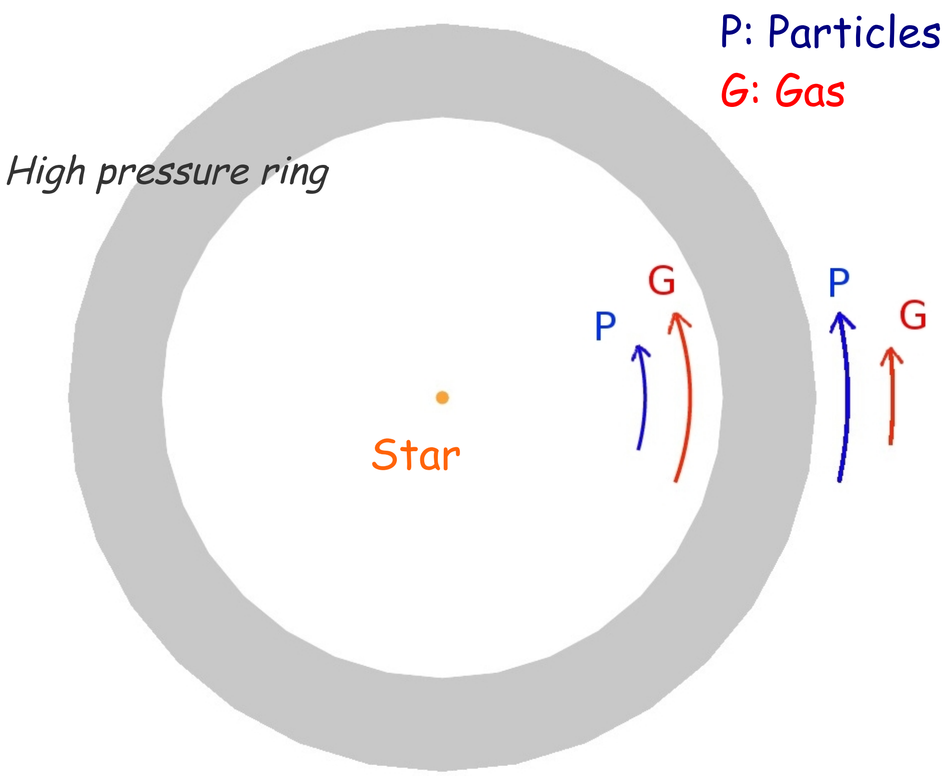
Particle drift



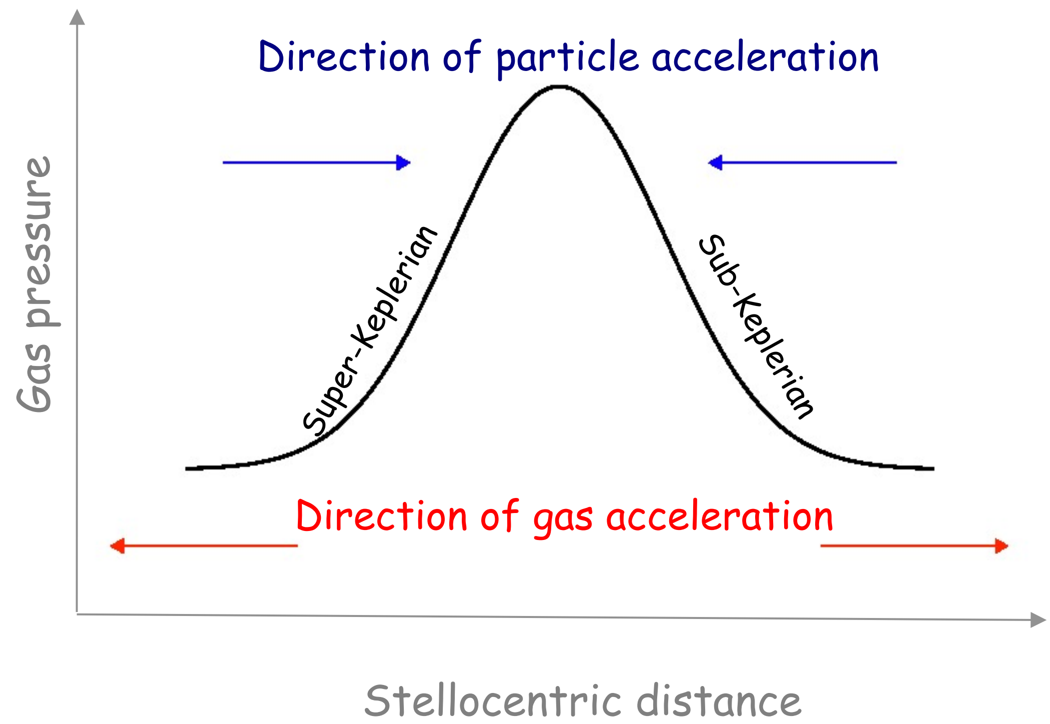
Adapted from Whipple (1972)



Pressure Trap

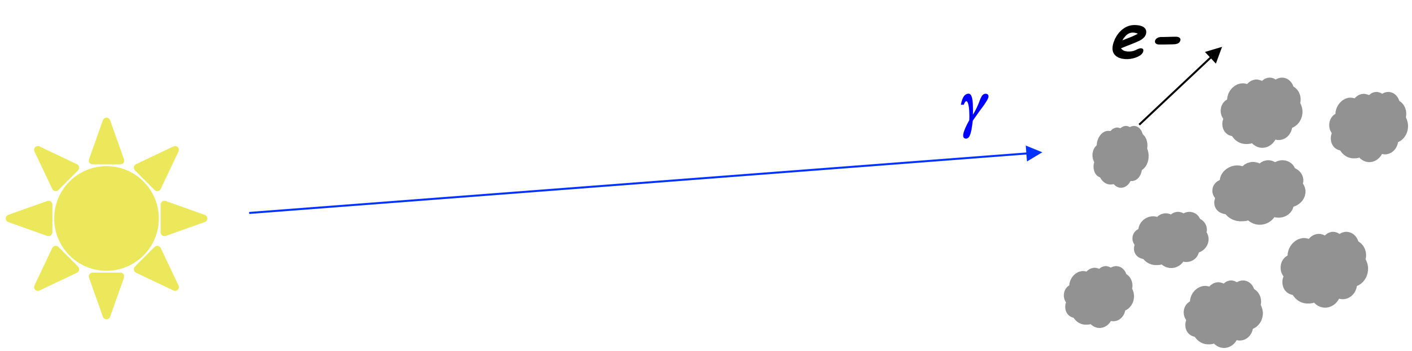


Adapted from Whipple (1972)



Photoelectric heating

In optically thin debris disks,
the **dust** is the **main heating agent** for the gas.



Dust intercepts starlight directly,
emits electron, that heats the gas.

Gas is photoelectrically heated by the dust

Runaway process: instability

Dust heats gas
Heated gas = high pressure region
High pressure concentrates dust

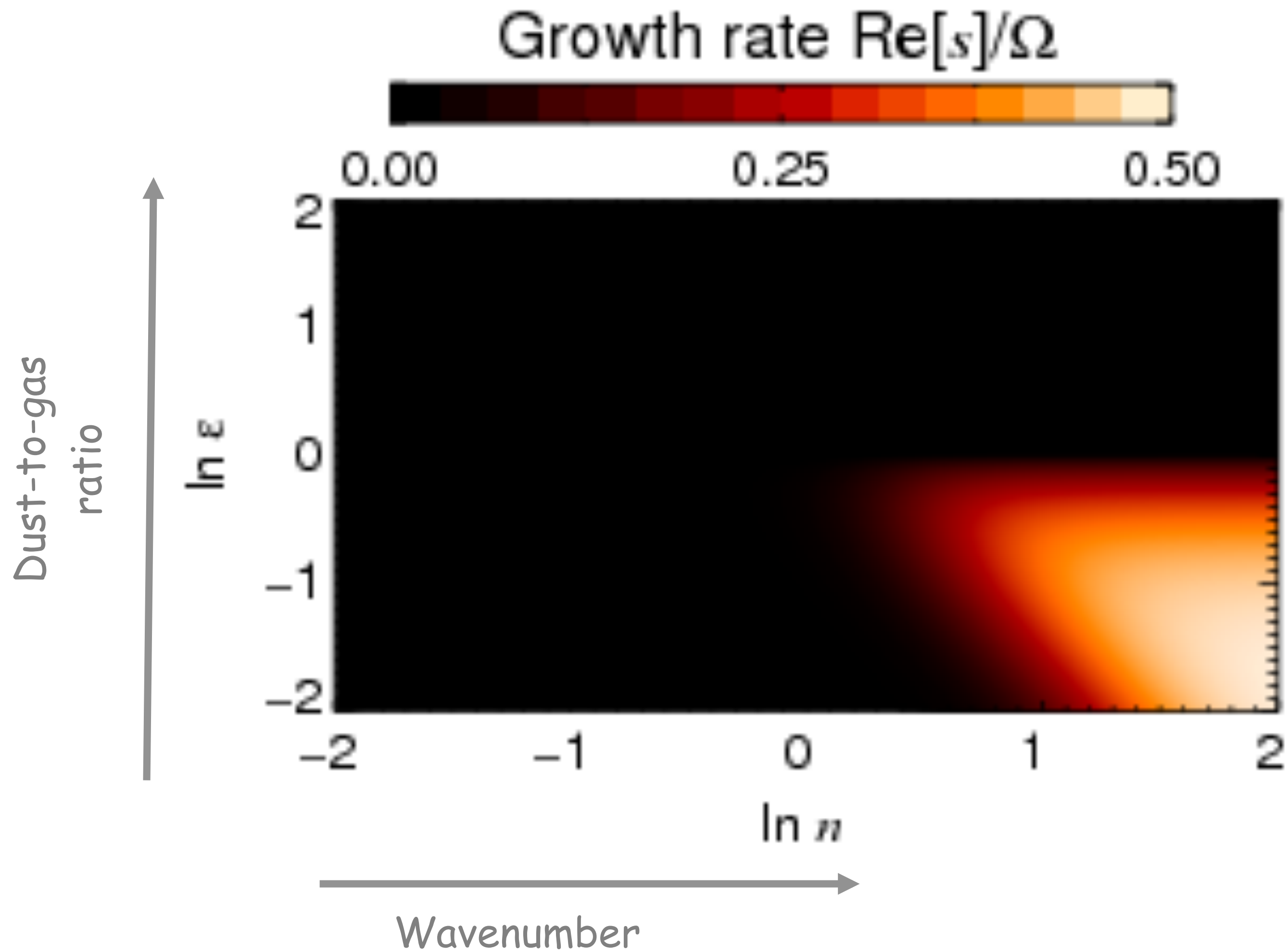
Runaway process: instability

Dust heats gas

Heated gas = high pressure region

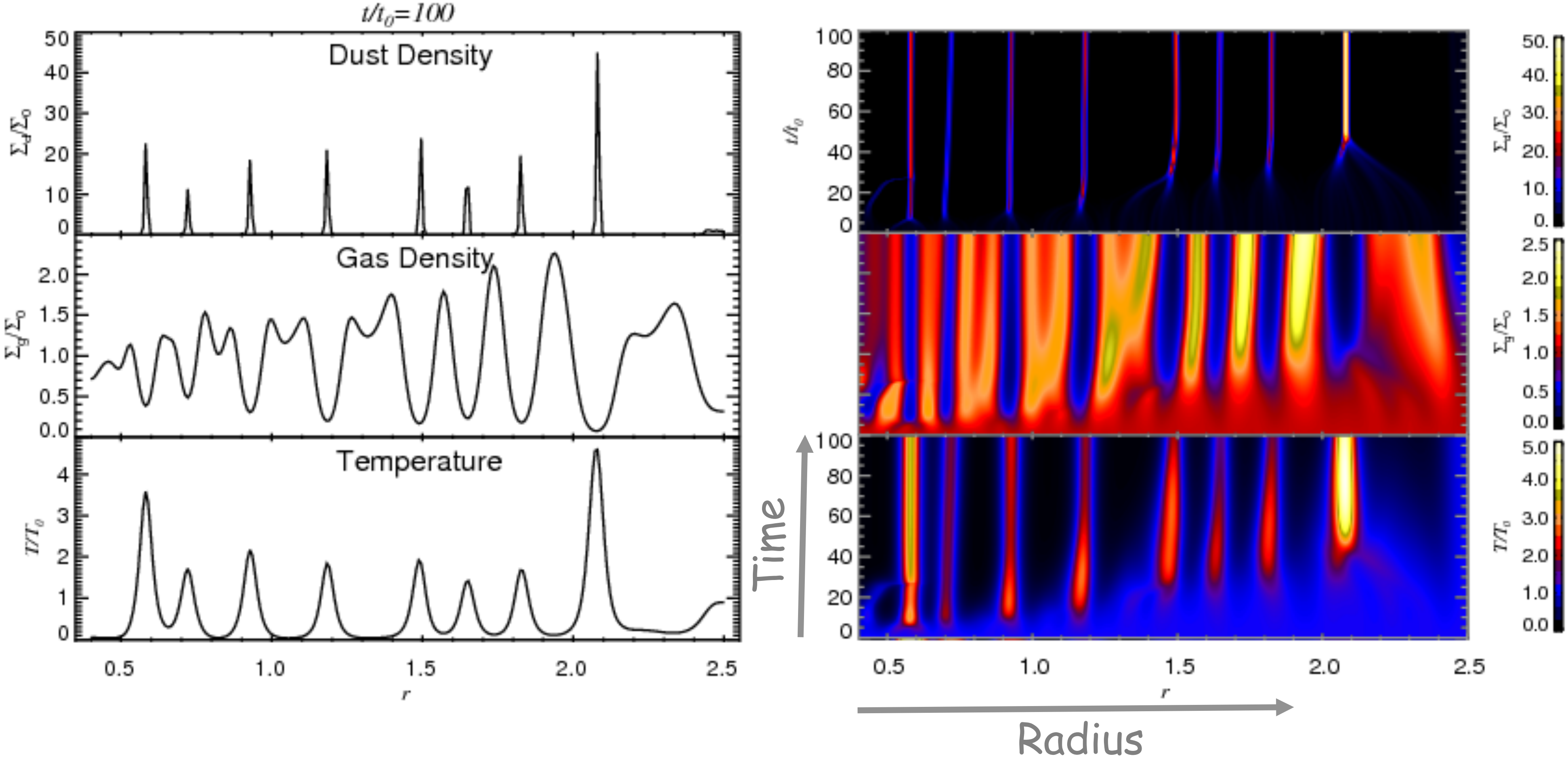
High pressure concentrates dust

Linear Analysis



Photoelectric Instability

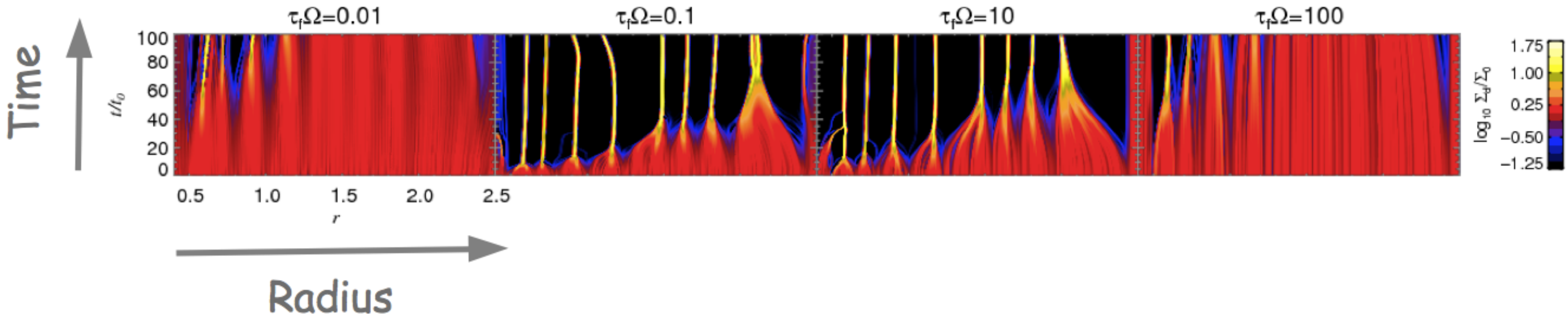
Dust heats gas
Heated gas = high pressure region
High pressure concentrates dust



Narrow hot dust rings
Cold gas collects between rings

Robustness

Growth over 4 orders of magnitude in dust-gas coupling time (friction time)



Photoelectric instability - 3D stratified local box

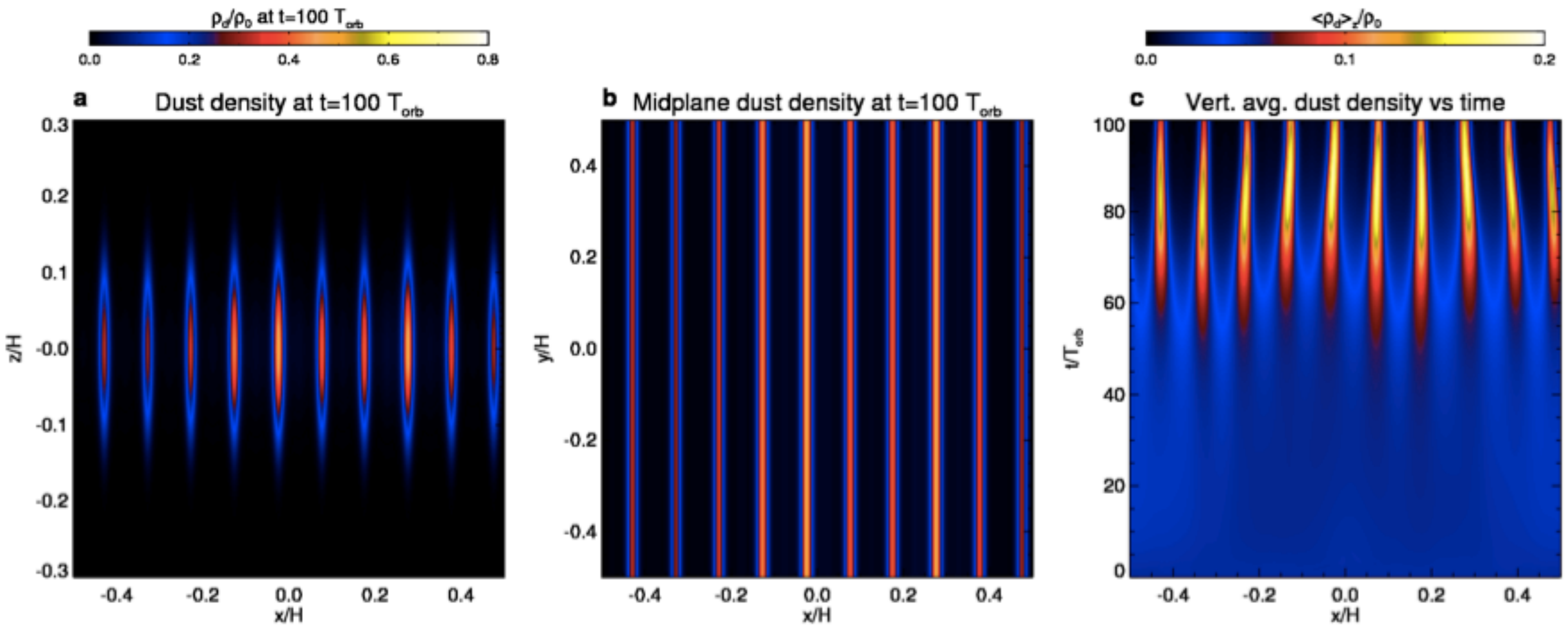
Photoelectric Instability

Dust heats gas

Heated gas = high pressure region

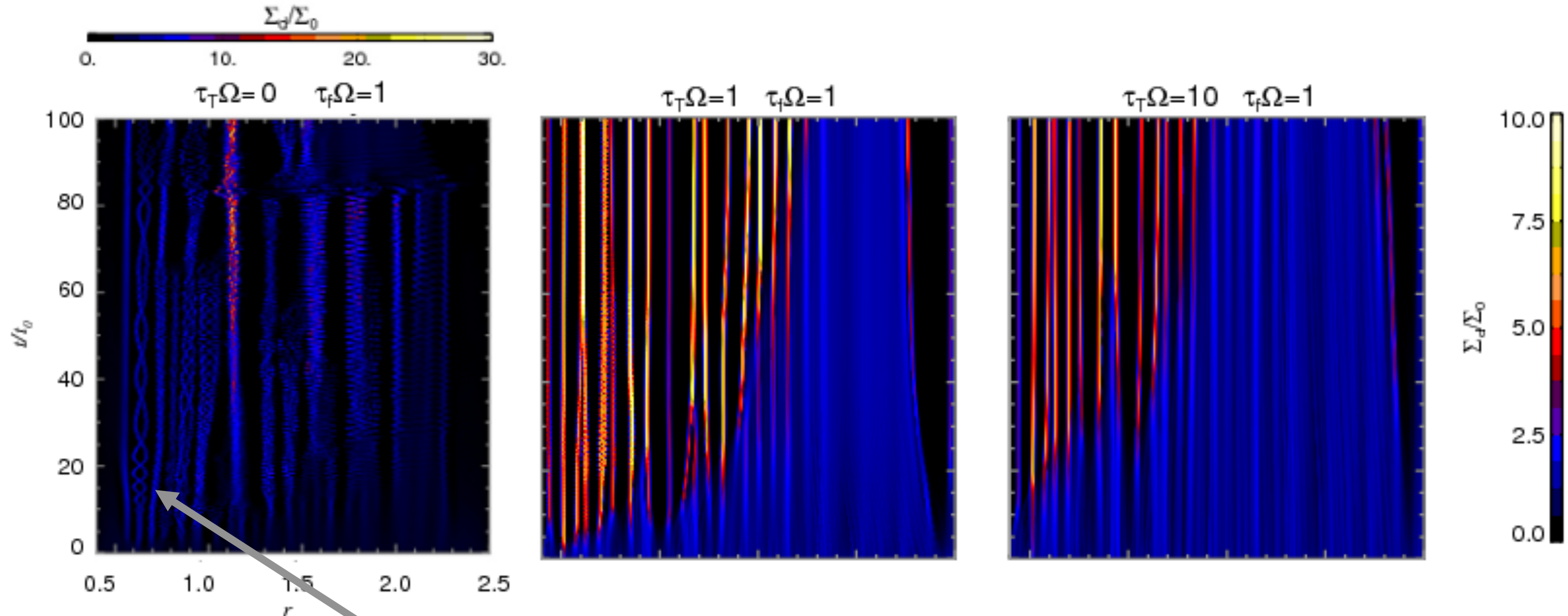
High pressure concentrates dust

3D Stratified



Oscillations

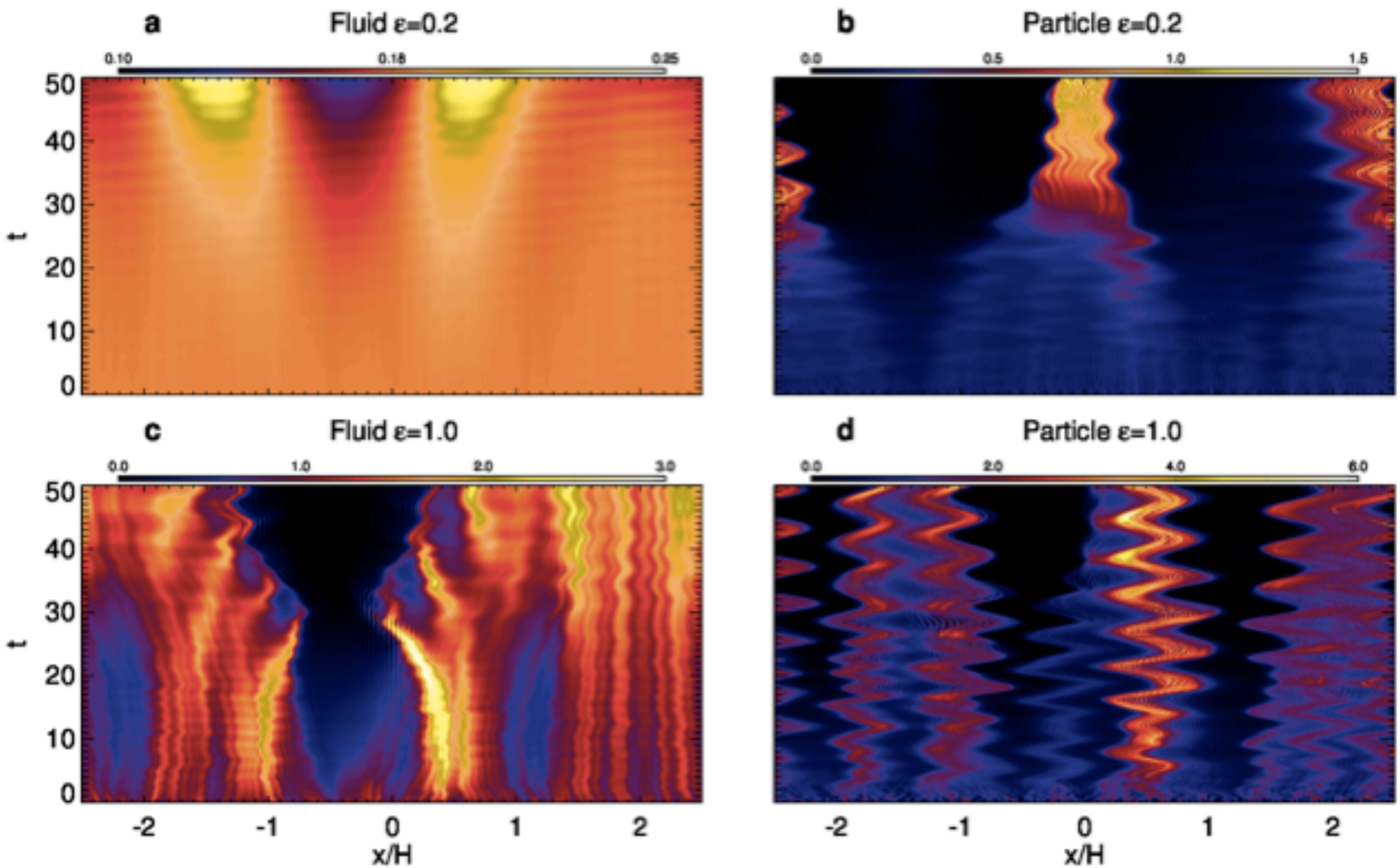
Thermal coupling time



Oscillations appear
with decreasing thermal time.

Oscillations

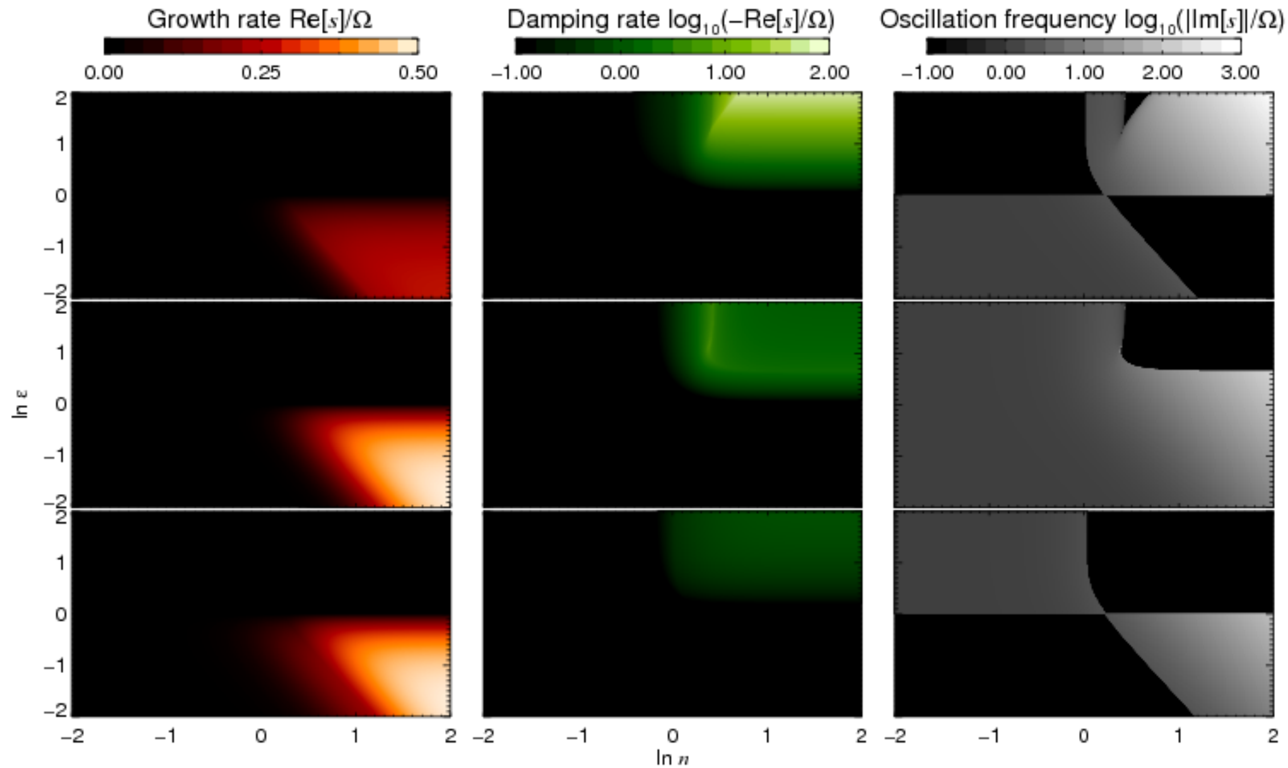
Low Reynolds number *High Reynolds number*



Epicyclic oscillations

clear at high Reynolds numbers!

Solutions



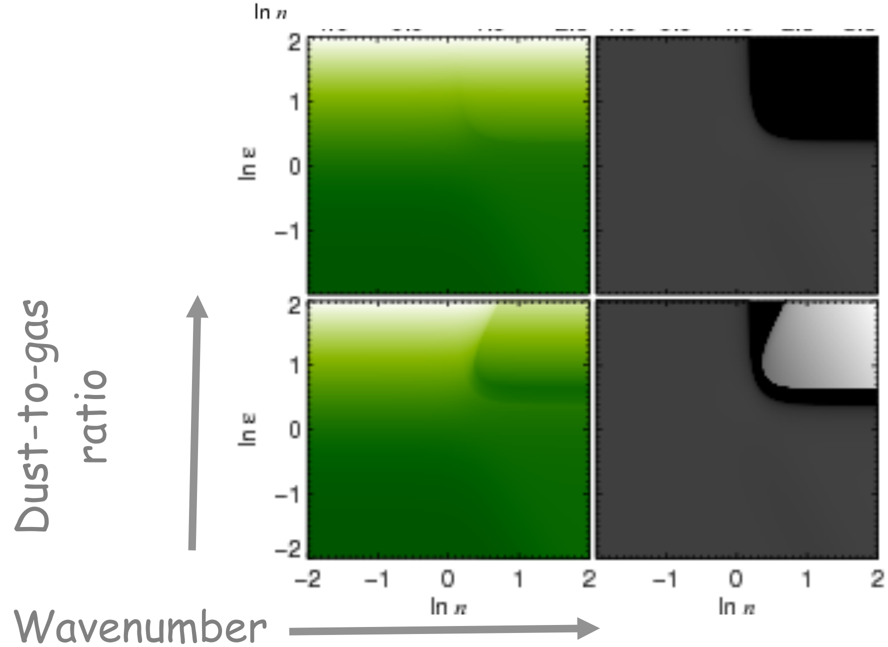
The dispersion relation is a 5[th] order polynomium, so there are five roots!

Dispersion relation

$$A\omega^5 + B\omega^4 + C\omega^3 + D\omega^2 + E\omega + F = 0$$

$$A = 1$$
$$B = 2\epsilon + 2$$
$$C = \epsilon^2 + \epsilon(n^2+2) + 3$$
$$D = \epsilon^2 n^2 + \epsilon(3n^2+2) + 2$$
$$E = \epsilon^2(2n^2+1) + \epsilon(3n^2+2) + 2$$
$$F = \epsilon^2 n^2 - \epsilon n^2$$

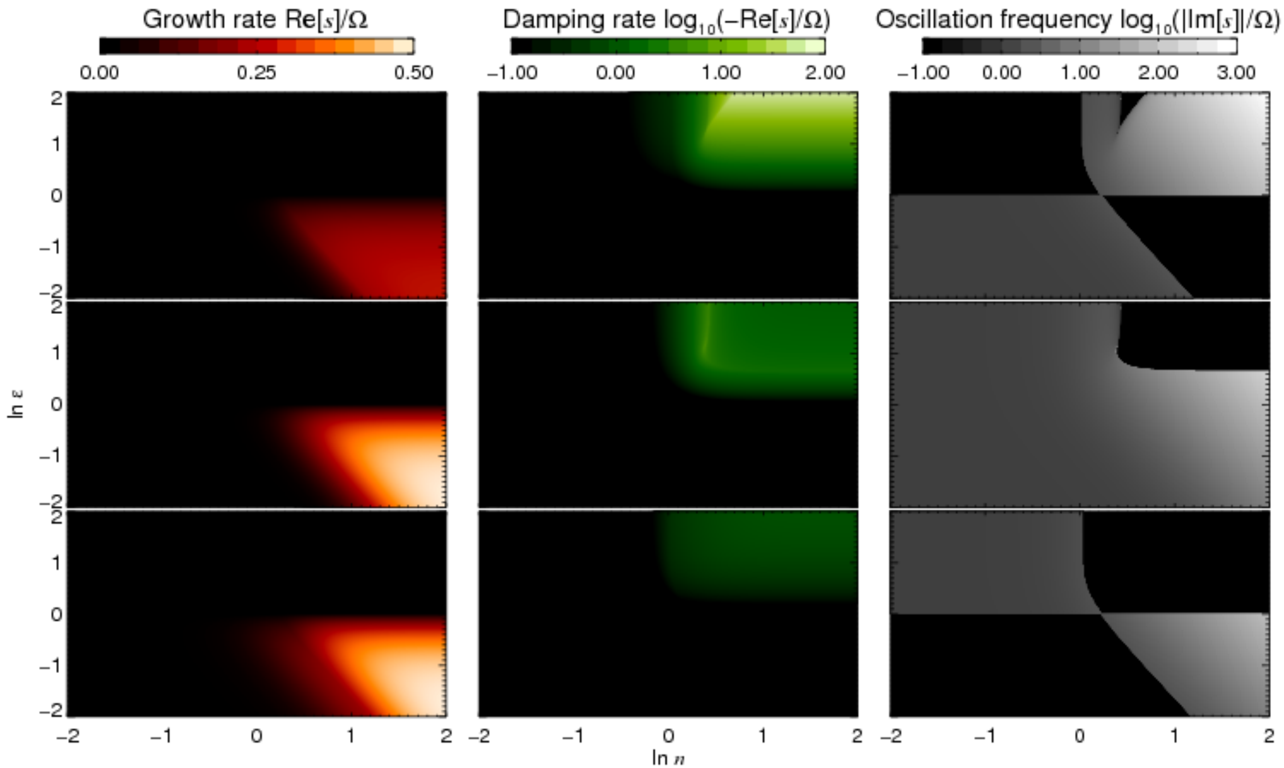
$\epsilon = \Sigma_d / \Sigma_g \qquad n = kH \qquad \omega = s/\Omega$



Solutions

Instability

Overstability



Damped and free Oscillations

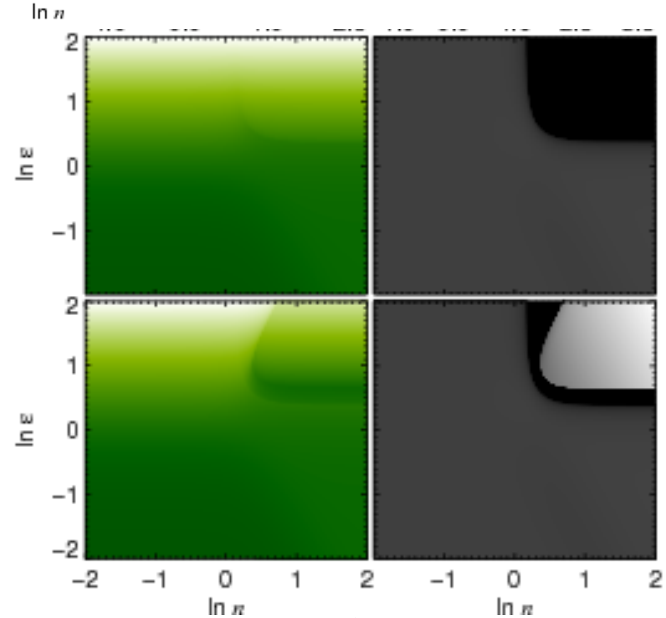
Dispersion relation

$$A\omega^5 + B\omega^4 + C\omega^3 + D\omega^2 + E\omega + F = 0$$

$$A = 1$$
$$B = 2\epsilon + 2$$
$$C = \epsilon^2 + \epsilon(n^2+2) + 3$$
$$D = \epsilon^2 n^2 + \epsilon(3n^2+2) + 2$$
$$E = \epsilon^2(2n^2+1) + \epsilon(3n^2+2) + 2$$
$$F = \epsilon^2 n^2 - \epsilon n^2$$

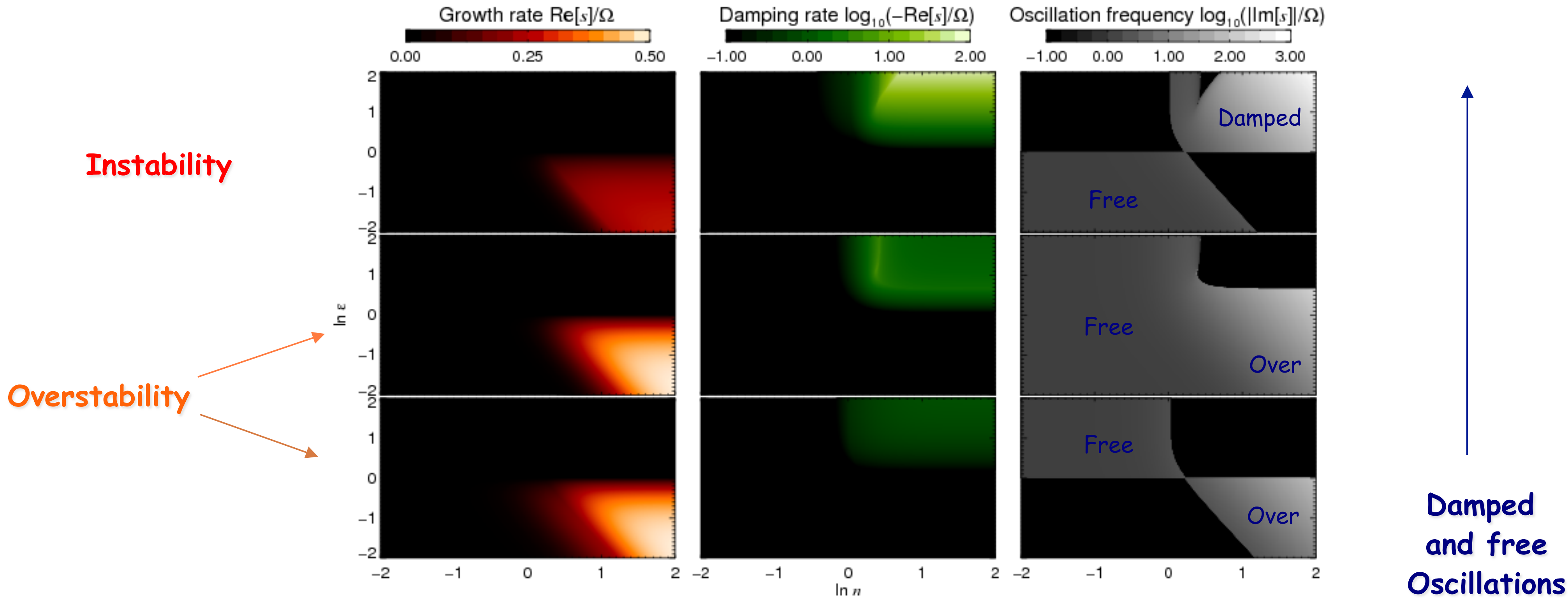
$\epsilon = \Sigma_d / \Sigma_g$ $\quad n = kH \quad$ $\omega = s/\Omega$

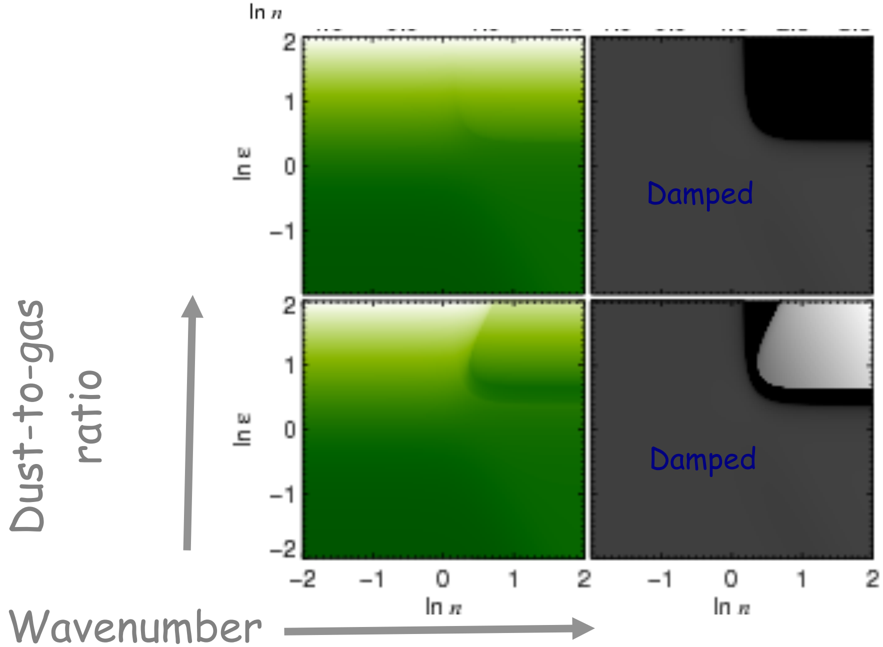
Dust-to-gas ratio



Wavenumber

Solutions





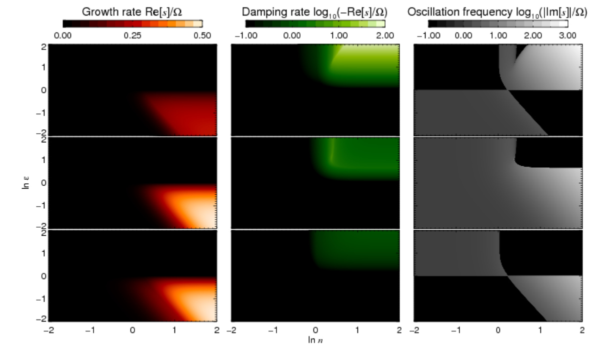
Dispersion relation

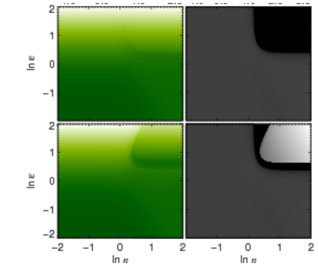
$$A\omega^5 + B\omega^4 + C\omega^3 + D\omega^2 + E\omega + F = 0$$

$$A = 1$$
$$B = 2\epsilon + 2$$
$$C = \epsilon^2 + \epsilon(n^2+2) + 3$$
$$D = \epsilon^2 n^2 + \epsilon(3n^2+2) + 2$$
$$E = \epsilon^2(2n^2+1) + \epsilon(3n^2+2) + 2$$
$$F = \epsilon^2 n^2 - \epsilon n^2$$

$\epsilon = \Sigma_d / \Sigma_g$ $\quad n = kH$ $\quad \omega = s/\Omega$

Solutions

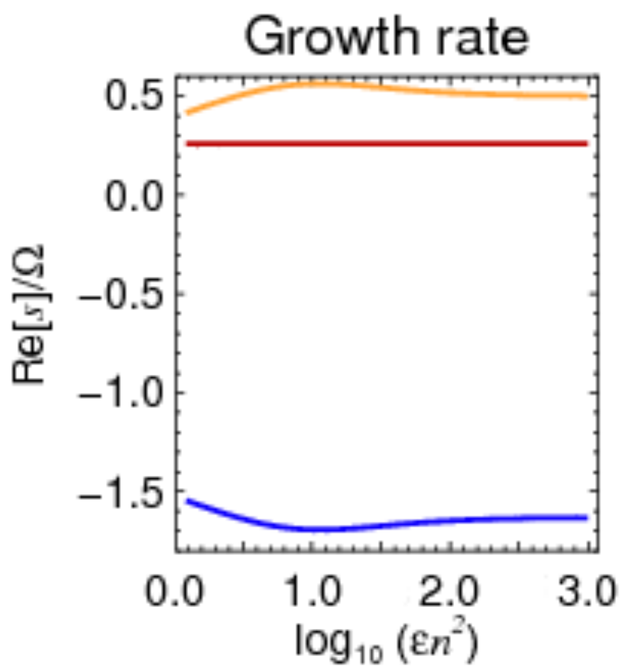


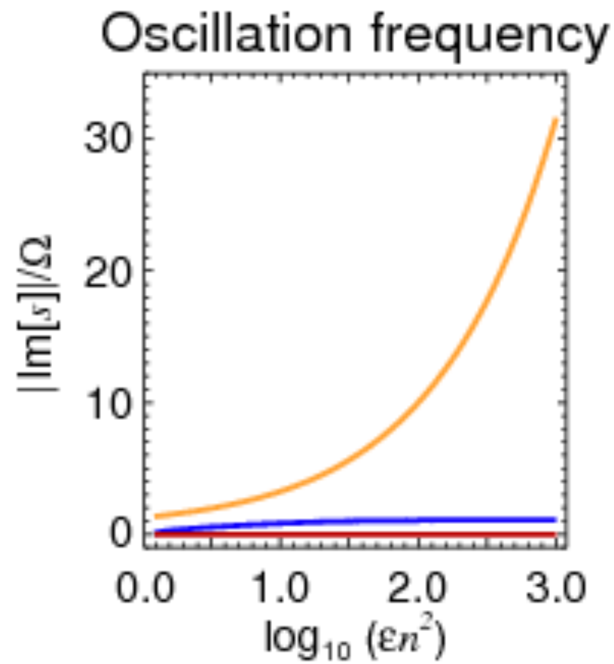


Overstability

Instability

Oscillations





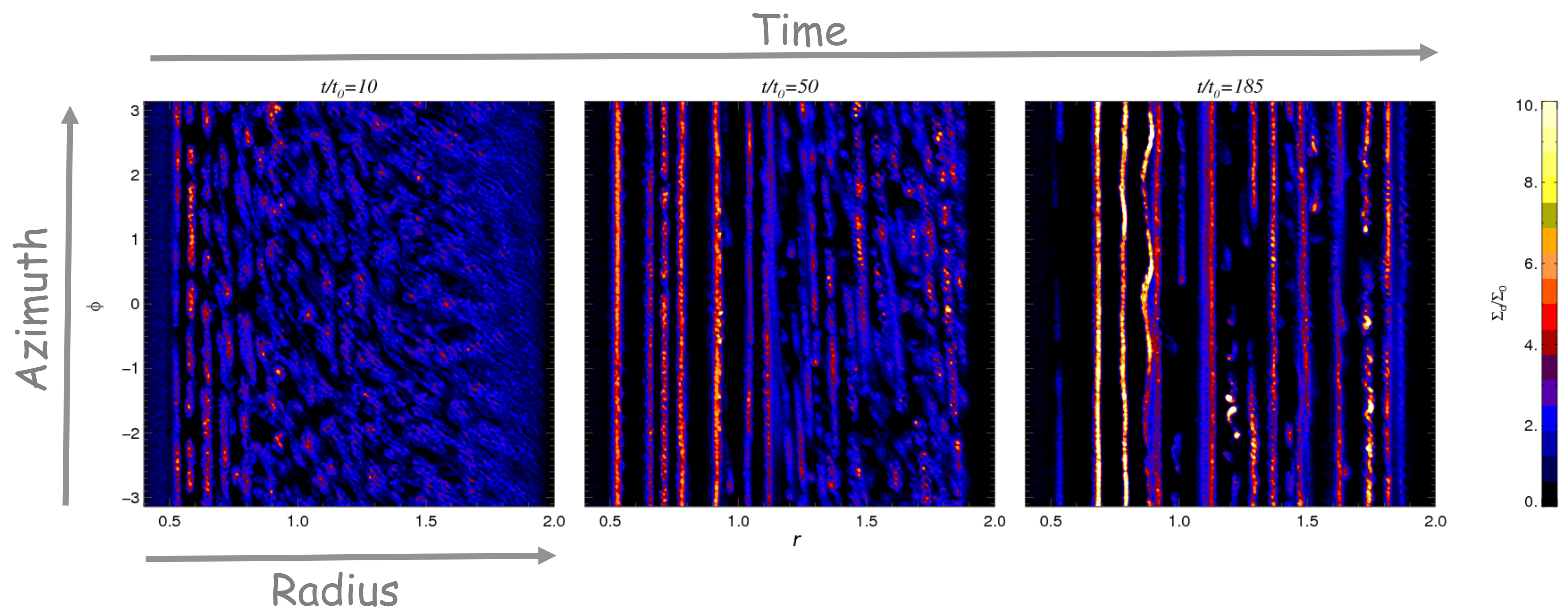
Max growth rate: Omega/2.

Million-fold amplification in five orbits!

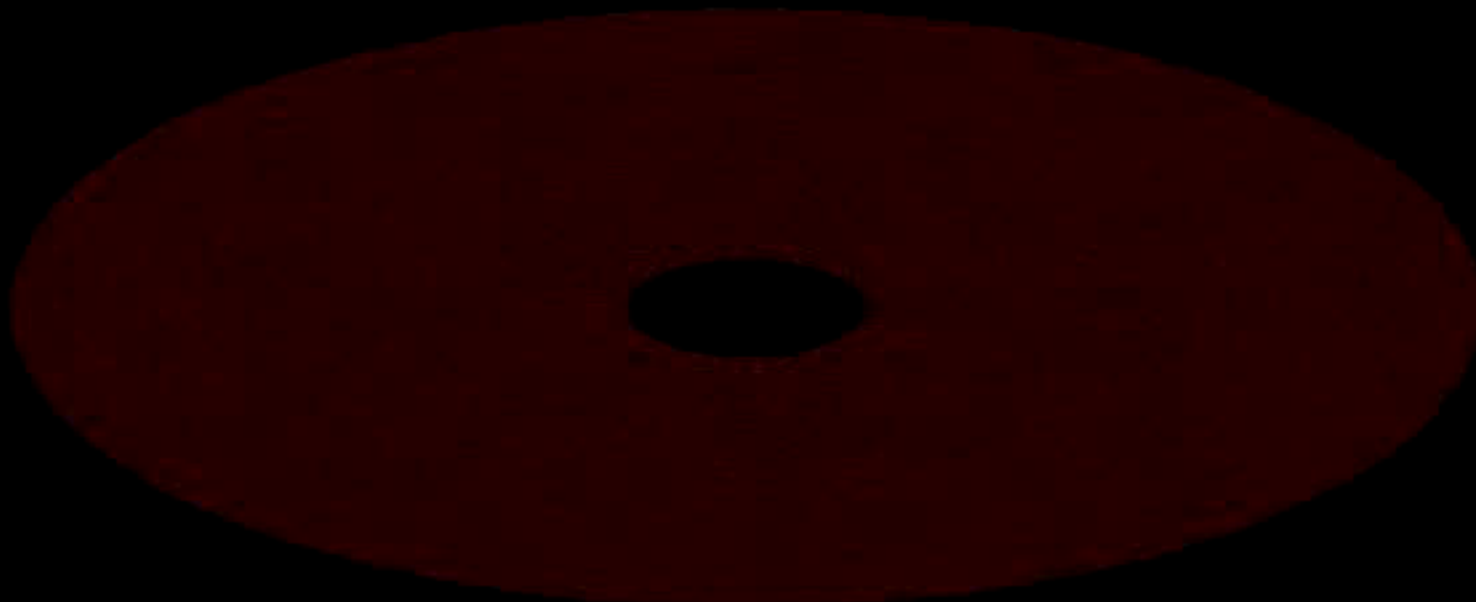
A very powerful instability.

The model in r-ϕ: Eccentric rings

Growth of axisymmetric modes
+
Damping of nonaxissymetric modes.
= *Rings !!!*

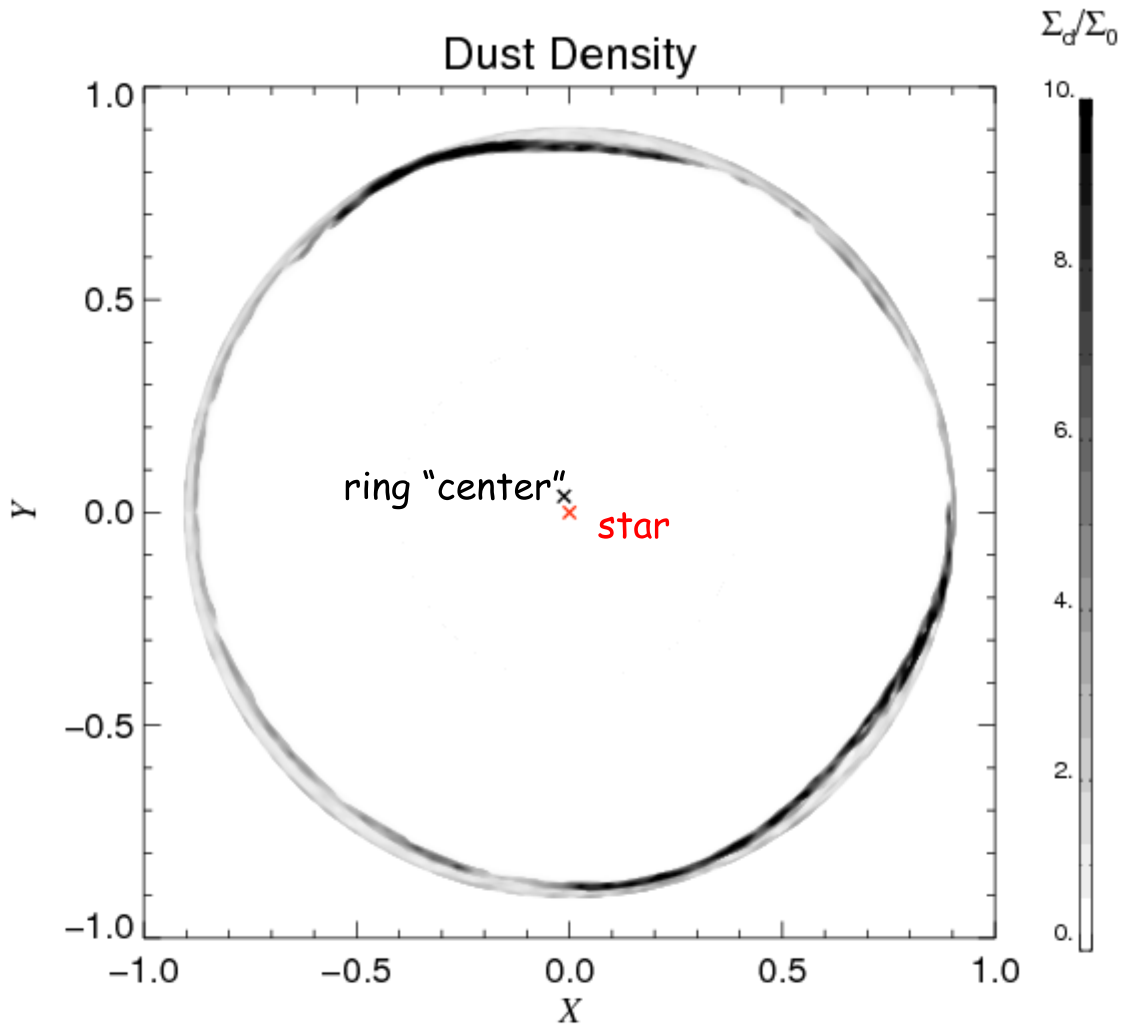


Epicyclic oscillations
make the ring appear ***eccentric !!!***



Lyra & Kuchner (2013, Nature, 499, 148)

Ring eccentricity



Eccentricity e=0.04

Conclusion

There is a robust ring-forming

photoelectric instability

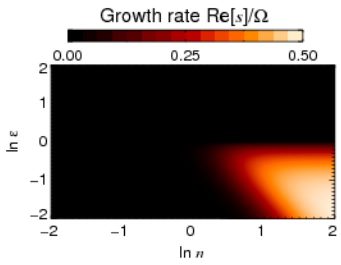
in debris disks

(in general, optically thin gas-dust disks)

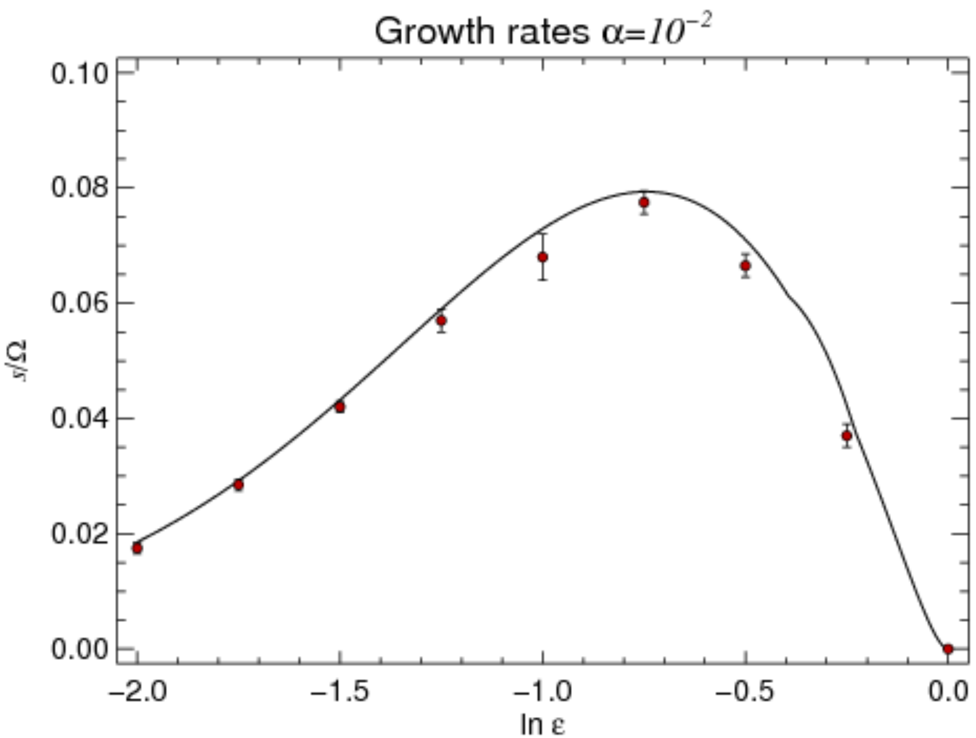
Careful

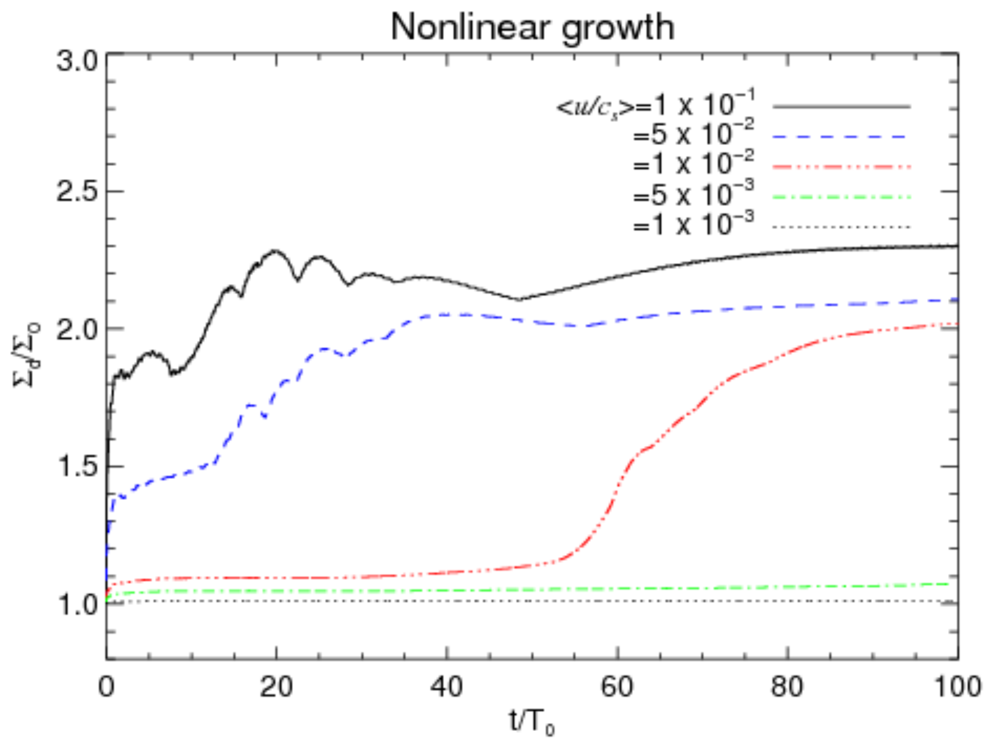
before you shout

planet!



Linear and nonlinear growth



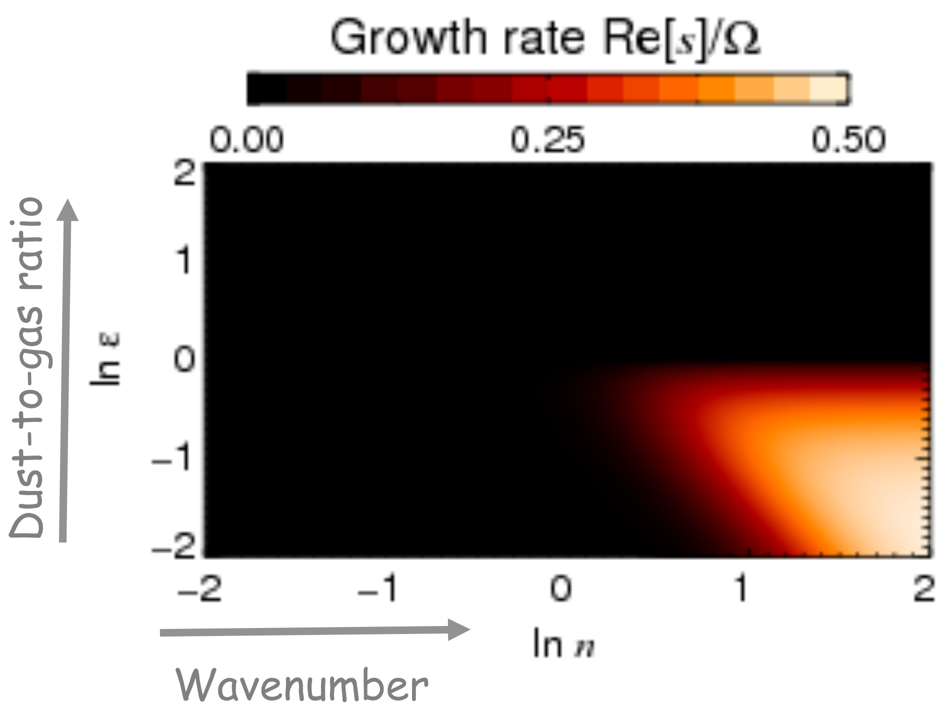


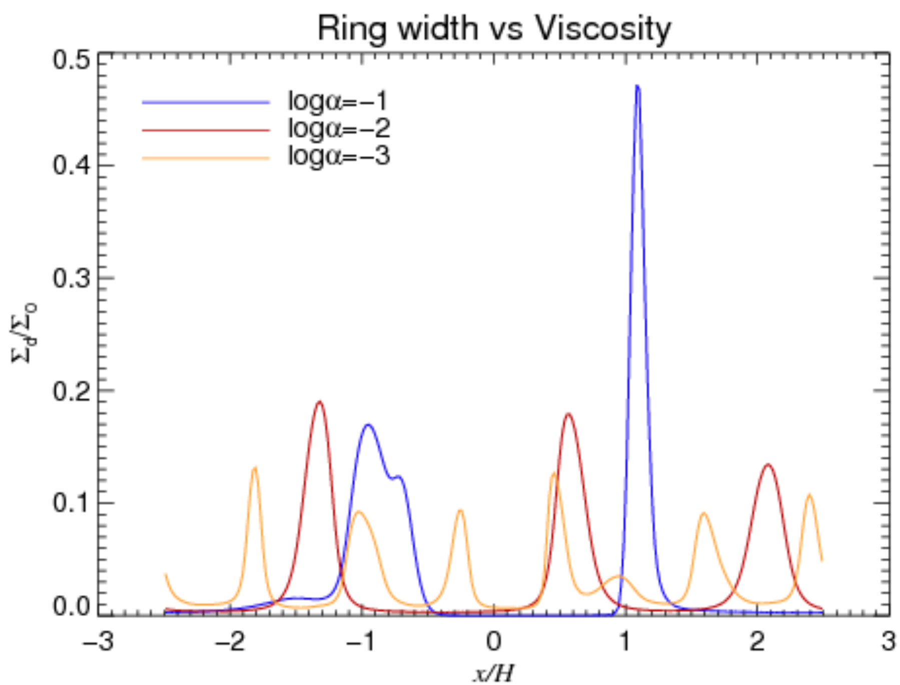
Linear growth only exists for ε < 1

But there is
nonlinear growth
beyond !

Ring Spacing

Ring spacing is determined
by the wavelength of maximum growth.





Which in turn is determined by viscosity

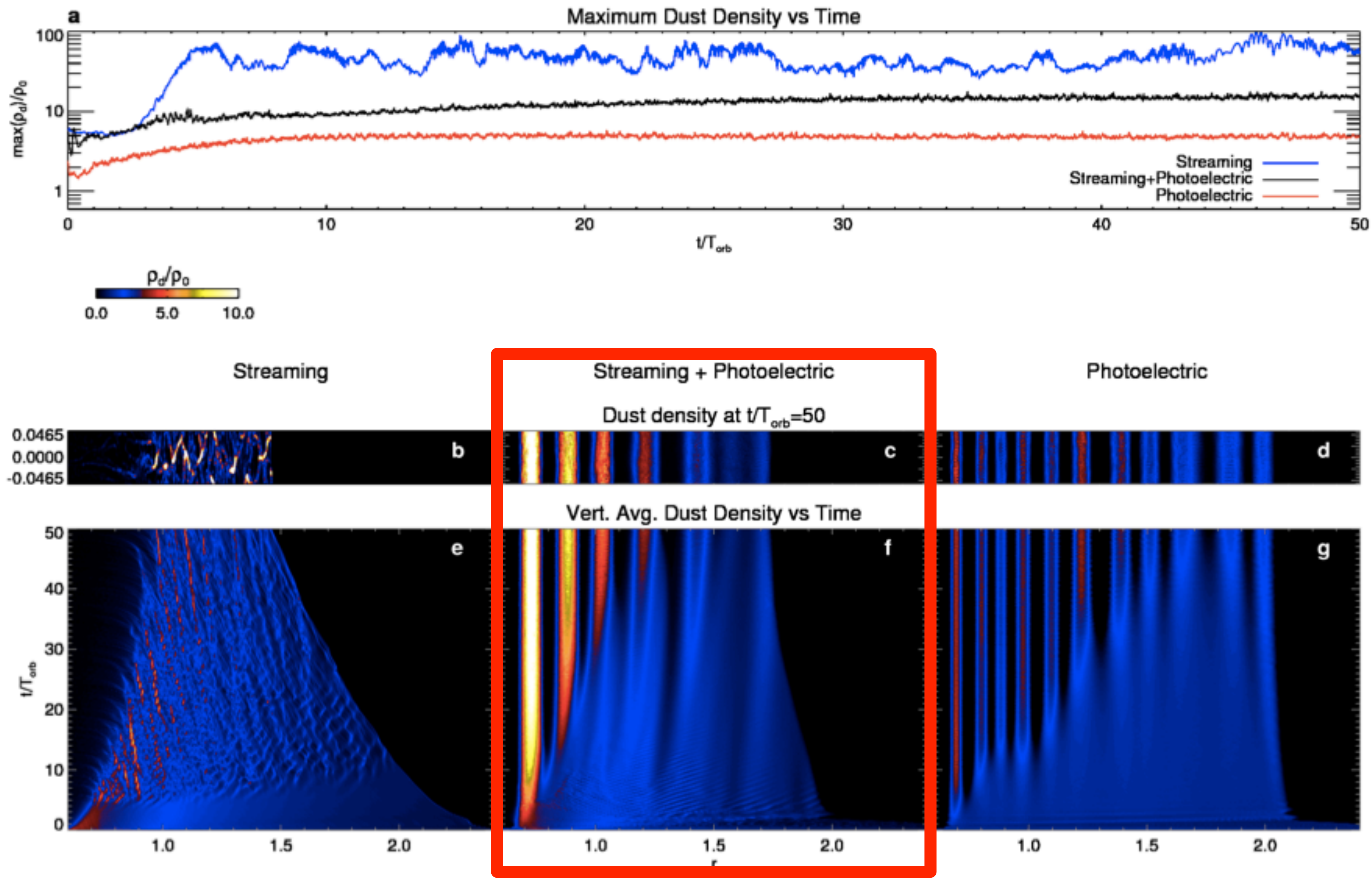
Ring spacing ~ 10 Kolmogorov lengths

Photoelectric vs Streaming Instability

Dust heats gas

Heated gas = high pressure region

High pressure concentrates dust



Model equations

Klahr & Lin (2005) used a simplified, 1-D model.

$$\frac{\partial}{\partial t}\Sigma_d + \frac{1}{r}\frac{\partial}{\partial r} r\Sigma_d v_r = 0.$$

Continuity equation

$$V_\phi = \Omega r + \frac{1}{2\Omega\Sigma_g}\frac{\partial}{\partial r}P$$

Terminal velocity

$$T_g = T_0 \left(\frac{\Sigma_d}{\Sigma_0}\right)^\beta ,$$

Equation of state

Model equations

Our simulation adds much more physics, and works in 2D.

Klahr & Lin (2005)
1D

$$\frac{\partial}{\partial t}\Sigma_d + \frac{1}{r}\frac{\partial}{\partial r} r\Sigma_d v_r = 0.$$

$$V_\phi = \Omega r + \frac{1}{2\Omega\Sigma_g}\frac{\partial}{\partial r}P$$

$$T_g = T_0 \left(\frac{\Sigma_d}{\Sigma_0}\right)^\beta,$$

Inertia for both gas and dust

Energy equation

Drag force and drag force backreaction

Lyra & Kuchner (2012)
2D

$$\frac{\partial \Sigma_g}{\partial t} = -(\boldsymbol{u}\cdot\boldsymbol{\nabla})\Sigma_g - \Sigma_g \boldsymbol{\nabla}\cdot\boldsymbol{u}$$

$$\frac{\partial \boldsymbol{u}}{\partial t} = -(\boldsymbol{u}\cdot\boldsymbol{\nabla})\boldsymbol{u} - \frac{1}{\Sigma_g}\boldsymbol{\nabla}P - \boldsymbol{\nabla}\Phi - \frac{\Sigma_d}{\Sigma_g}\boldsymbol{f}_d$$

$$\frac{\partial S}{\partial t} = -(\boldsymbol{u}\cdot\boldsymbol{\nabla})S - \frac{c_v}{T}\frac{(T-T_p)}{\tau_T}.$$

$$\frac{d\boldsymbol{x}}{dt} = \boldsymbol{v}$$

$$\frac{d\boldsymbol{v}}{dt} = -\boldsymbol{\nabla}\Phi + \boldsymbol{f}_d$$

$$f_d = -\frac{(v-u)}{\tau_f}$$

$$T_p = T_0\frac{\Sigma_d}{\Sigma_0}.$$

Linear Analysis

Dust

$$D_w \Sigma_d = -\Sigma_d \nabla \cdot \boldsymbol{w}$$

$$D_w w_x = 2\Omega w_y - \frac{1}{\tau_f}(w_x - u_x)$$

$$D_w w_y = -\frac{1}{2}\Omega w_x - \frac{1}{\tau_f}(v_y - u_y)$$

Gas

$$D_u \Sigma_g = -\Sigma_g \nabla \cdot \boldsymbol{u}$$

$$D_u u_x = 2\Omega u_y - \frac{1}{\Sigma_g}\frac{\partial P}{\partial x} - \frac{\epsilon}{\tau_f}(u_x - w_x)$$

$$D_u u_y = -\frac{1}{2}\Omega u_x - \frac{1}{\Sigma_g}\frac{\partial P}{\partial y} - \frac{\epsilon}{\tau_f}(u_y - w_y)$$

$$\psi = \psi_0 + \psi'$$

$$\psi' = \hat{\psi}\exp(ikx + st)$$

Dispersion relation

$$A\omega^5 + B\omega^4 + C\omega^3 + D\omega^2 + E\omega + F = 0$$

$$A = 1$$
$$B = 2\epsilon + 2$$
$$C = \epsilon^2 + \epsilon(n^2 + 2) + 3$$
$$D = \epsilon^2 n^2 + \epsilon(3n^2 + 2) + 2$$
$$E = \epsilon^2(2n^2 + 1) + \epsilon(3n^2 + 2) + 2$$
$$F = \epsilon^2 n^2 - \epsilon n^2$$

$\epsilon = \Sigma_d / \Sigma_g$ $\quad n = kH$ $\quad \omega = s/\Omega$