Signatures of Planet Formation in Transition Disks

Wladimir Lyra

New Mexico State University

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Funding

AAG – 2010, 2020

TCAN – 2020
(with J. Simon)
NFDAP – 2019
XRP – 2016, 2018

NRAO - 2017

HST - 2016

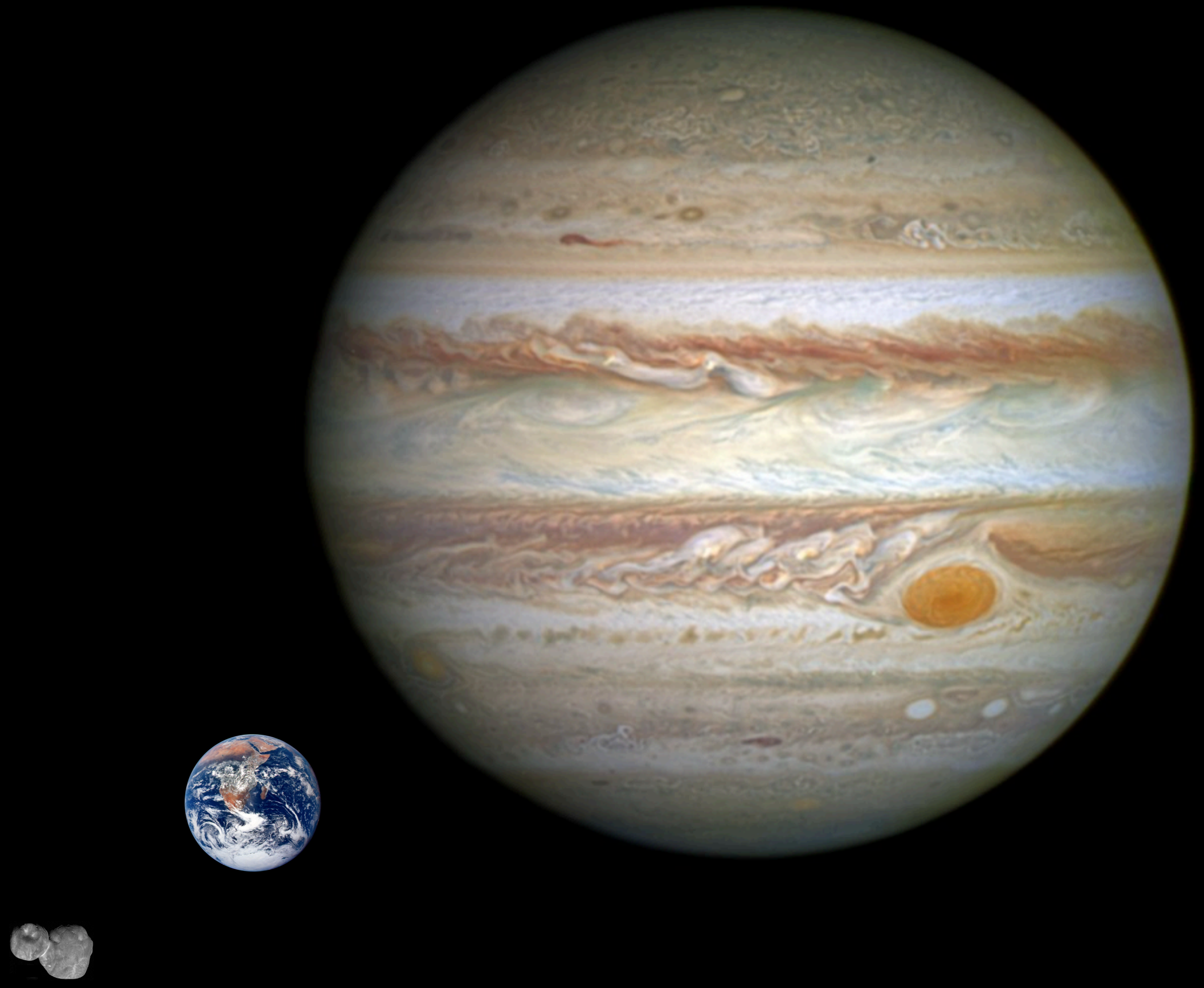
Computational Facilities

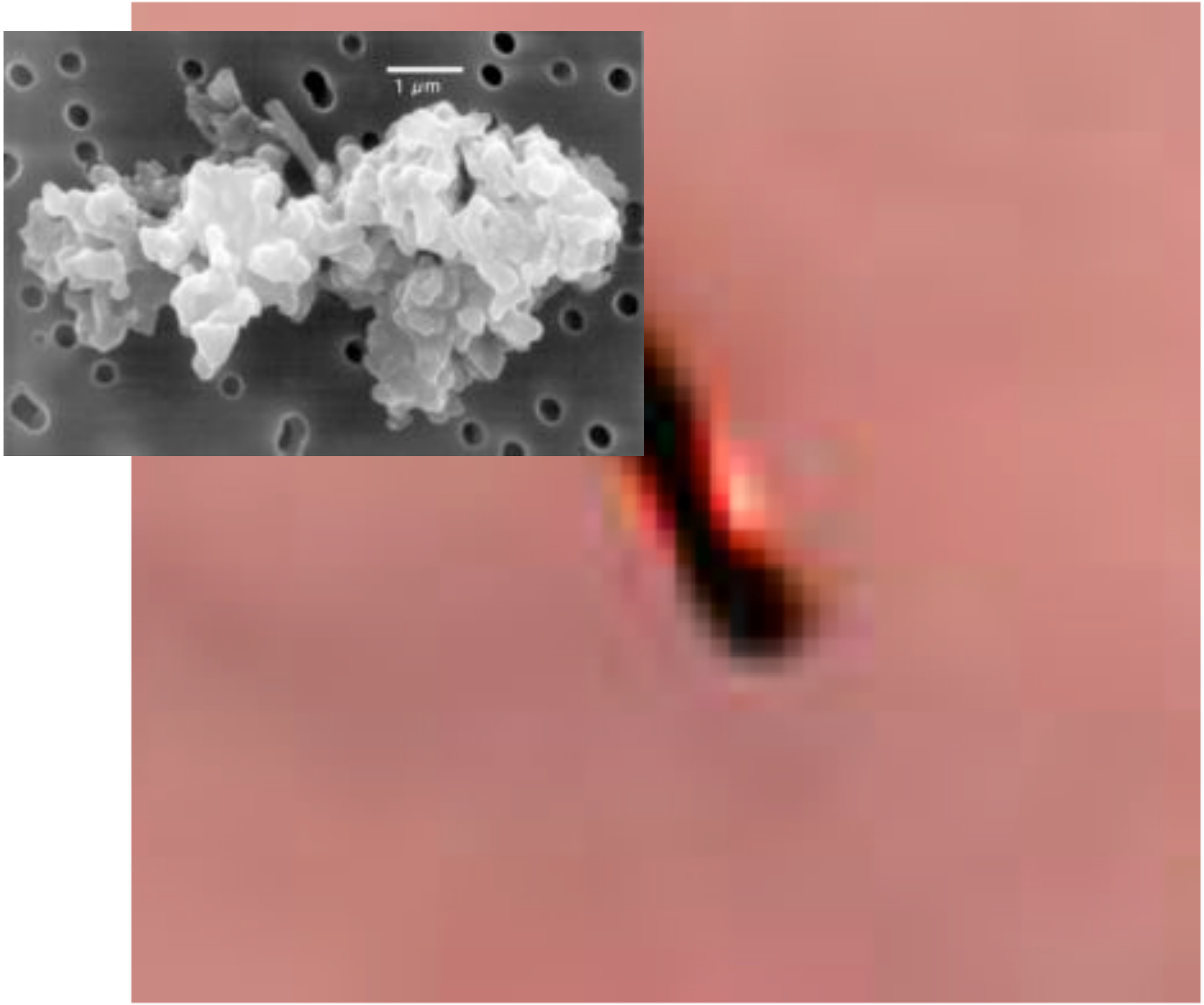
UPPSALA UNIVERSITET

Outline

- Planet Formation
- Disk Instabilities
- Disk observations

Part 1. Planet Formation




1 μm

Circumstellar/Protoplanetary Disks

PP disk fact sheet

Density: $10^{13} – 10^{15}$ cm^{-3}
(Air: 10^{21} cm^{-3})

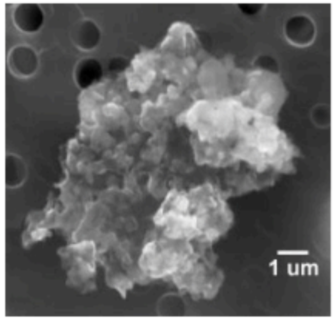
Temperature: 10-1000 K

Scale: 0.1-100AU

Mass: $10^{-3} – 10^{-1}$ M_{sun}

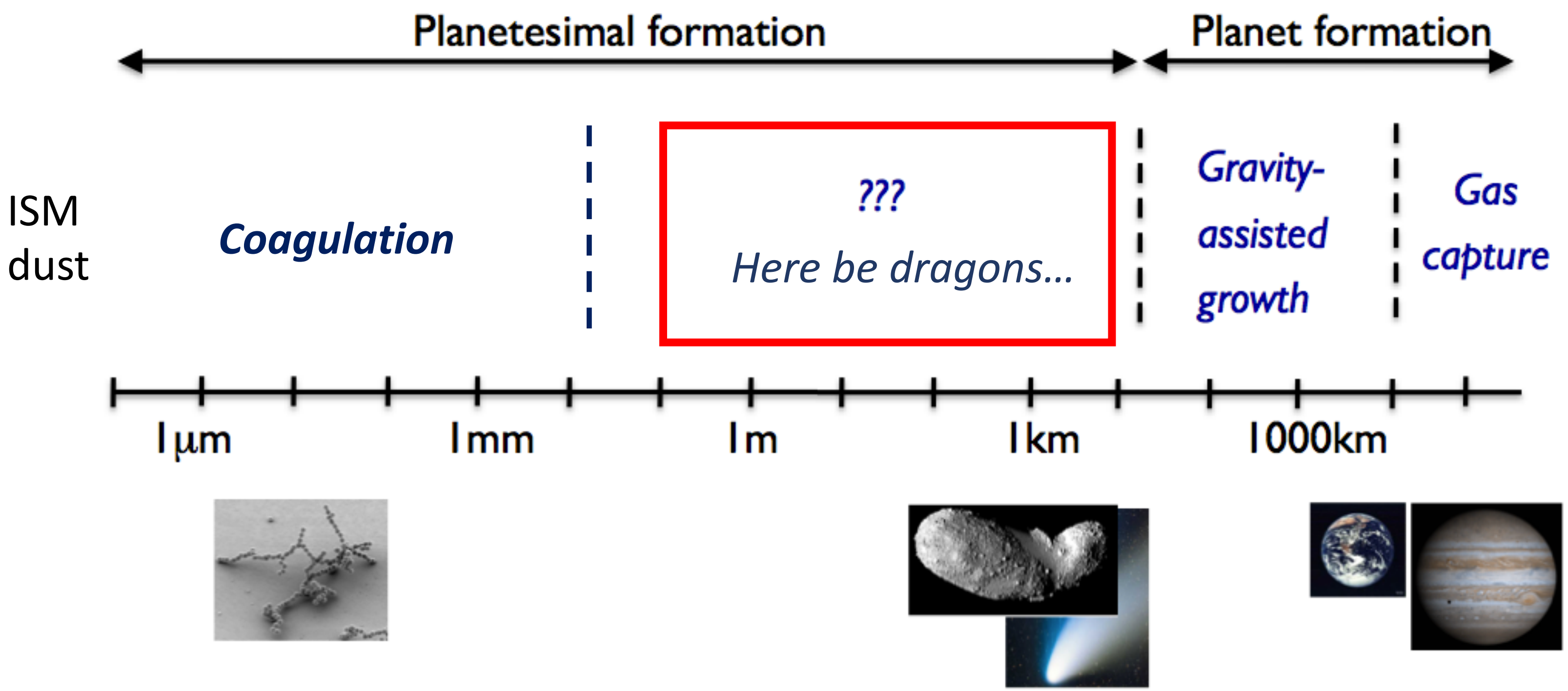
Planet Formation

"Planets form in disks of gas and dust"

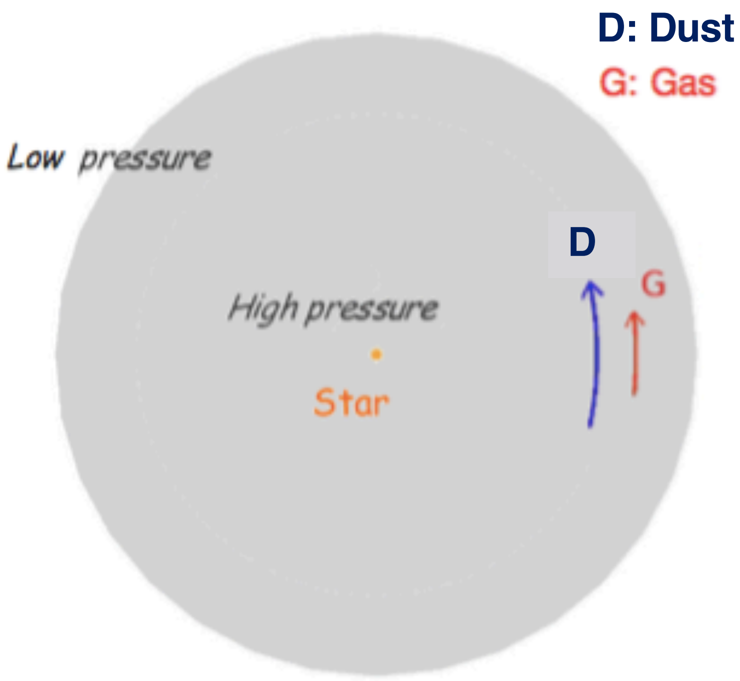


A miracle happens →

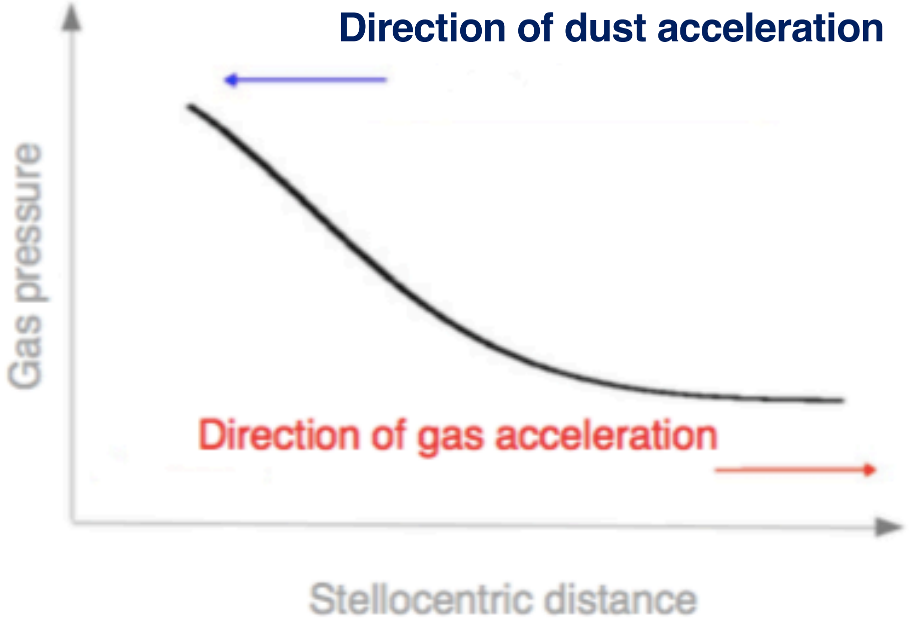
Dust evolution



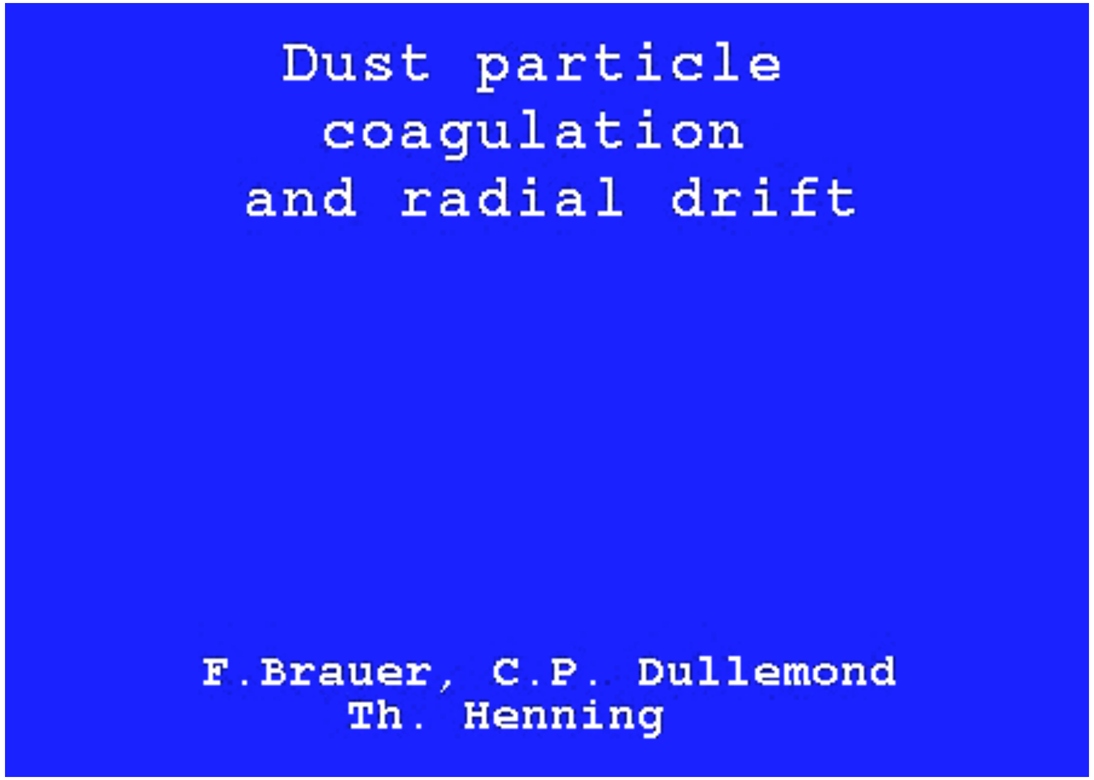
Dust Drift



Adapted from Whipple (1972)



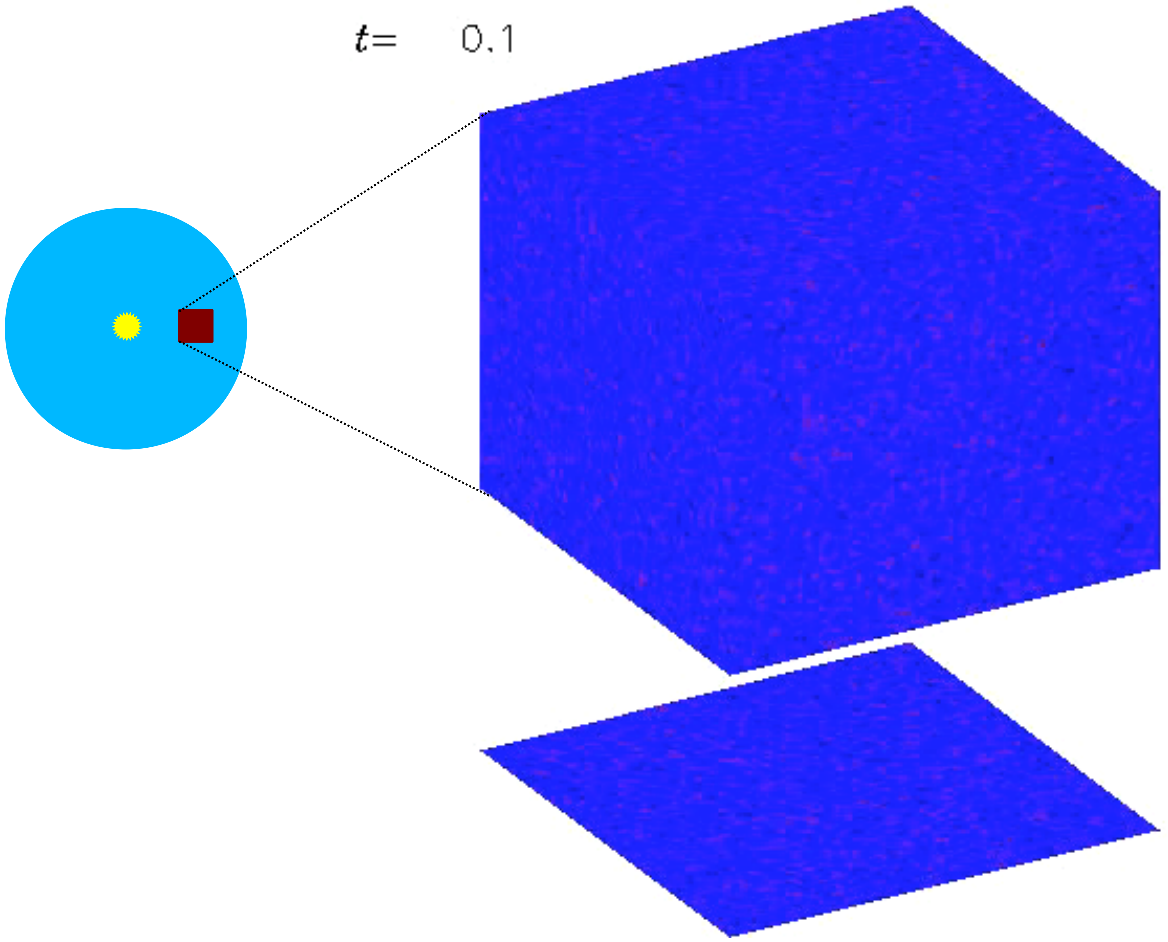
Dust coagulation and drift



Brauer et al. (2008)

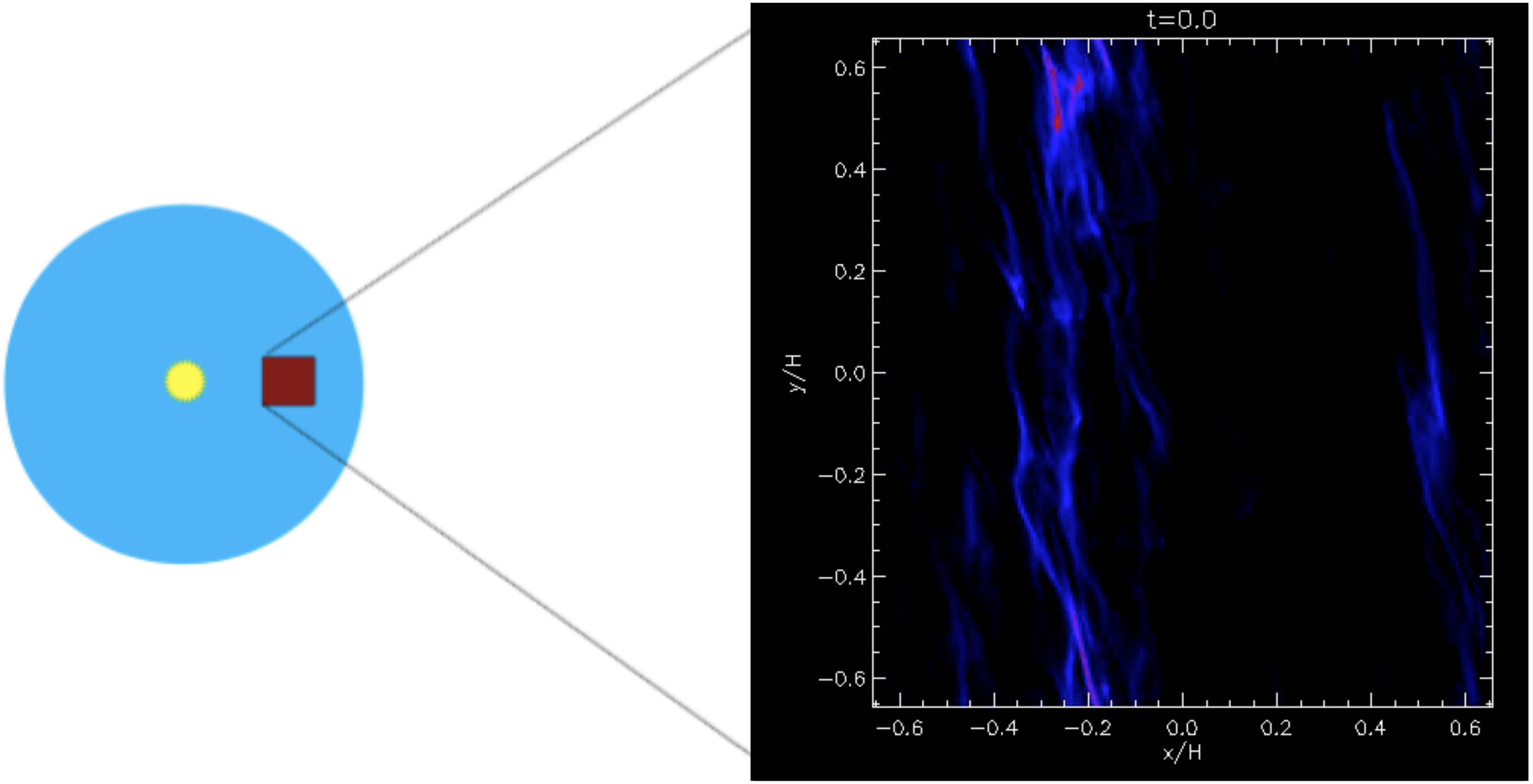
Streaming Instability

The dust drift is hydrodynamically unstable



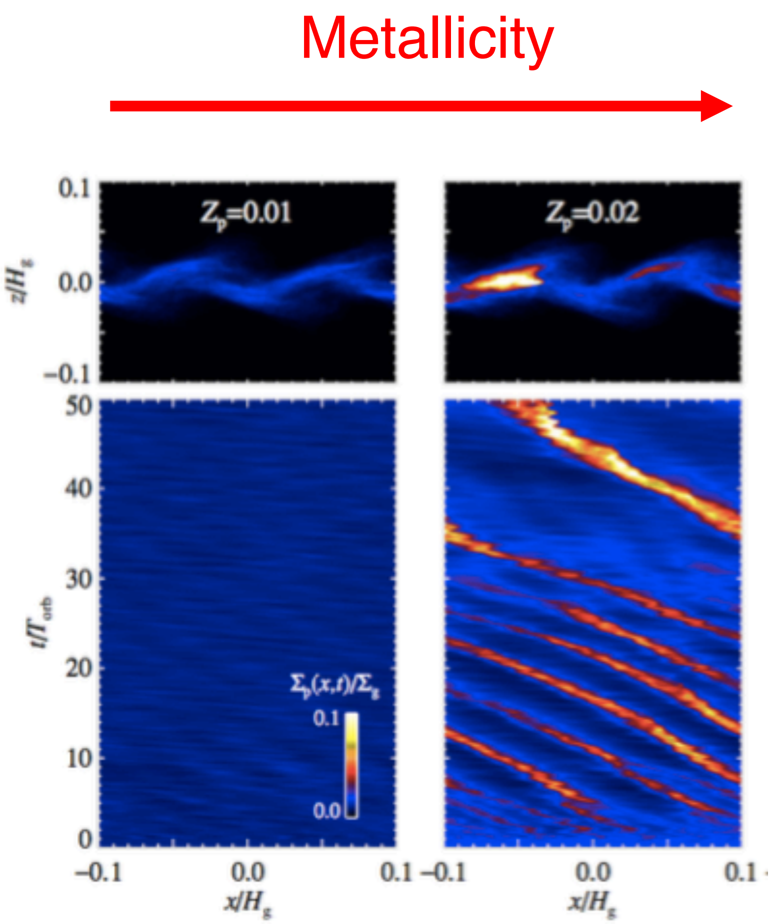
Youdin & Goodman ‘05, Johansen & Youdin ‘07, Youdin & Johansen+ ‘07, Kowalik+ ’13, Lyra & Kuchner ‘13,
Schreiber+ ‘18, Klahr & Schreiber ’20, Simon+ ‘16, ‘17, Carrera+ ‘15, ‘17, ‘20, Gole+ ’20, Li+ ‘18, ‘19, Abod+ ’19, Nesvorny+ ’19

Gravitational collapse into planetesimals

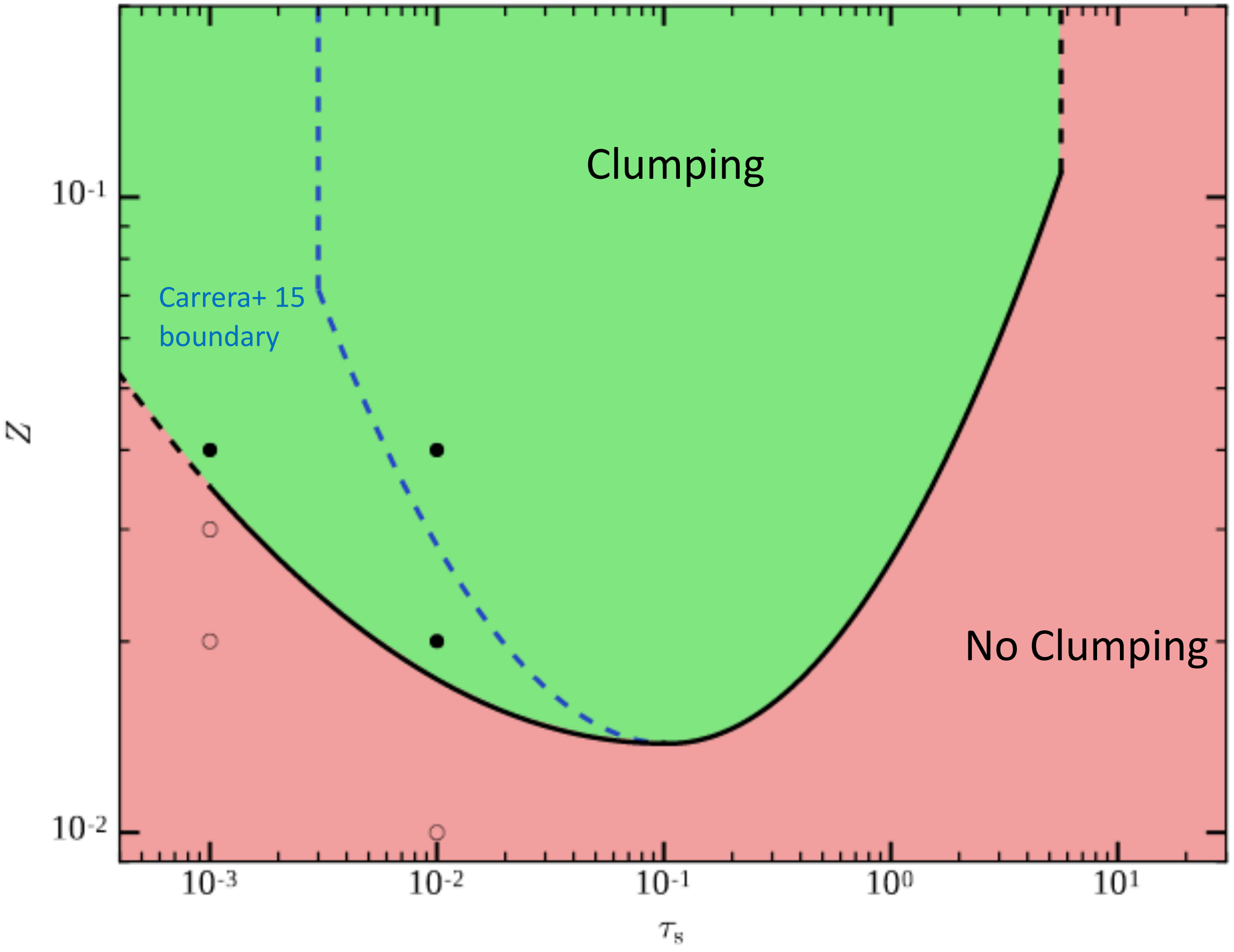


Johansen et al. (2007)

No clumping for solar metallicity



Johansen et al. (2011)

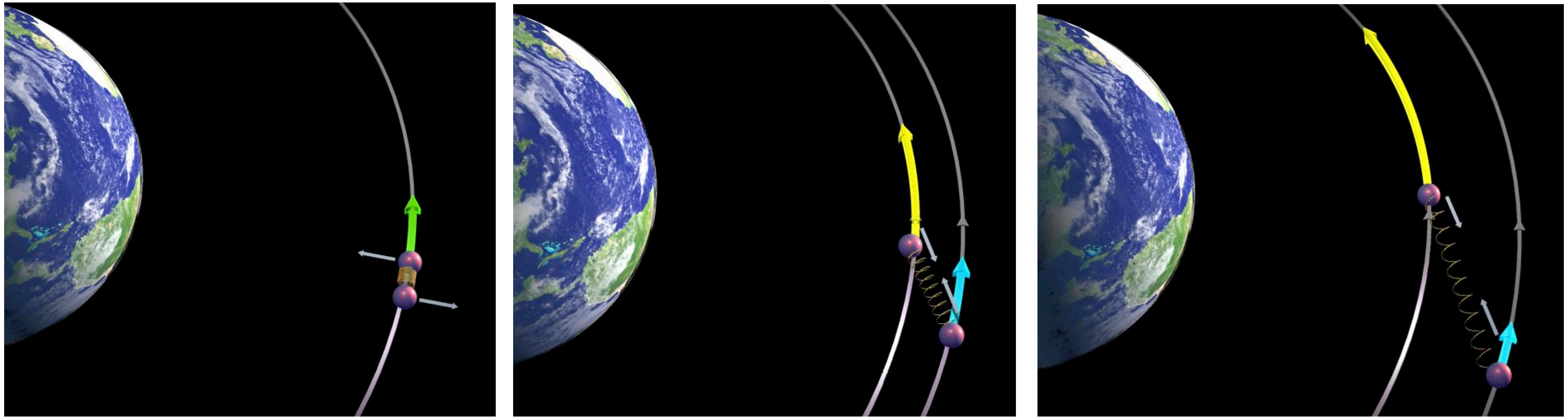


Yang+ 17 (after Carrera et al. 2015)

Turbulence and Accretion in 3D Global MHD Simulations of Stratified Protoplanetary Disk

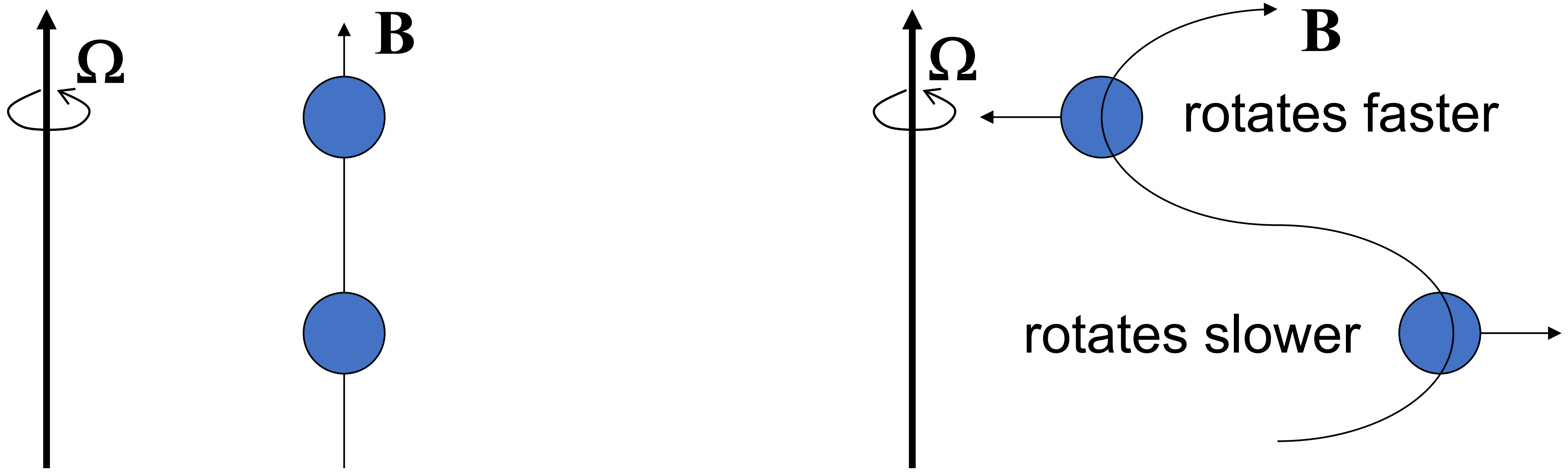
Stretching builds up tension

Tension resists shear



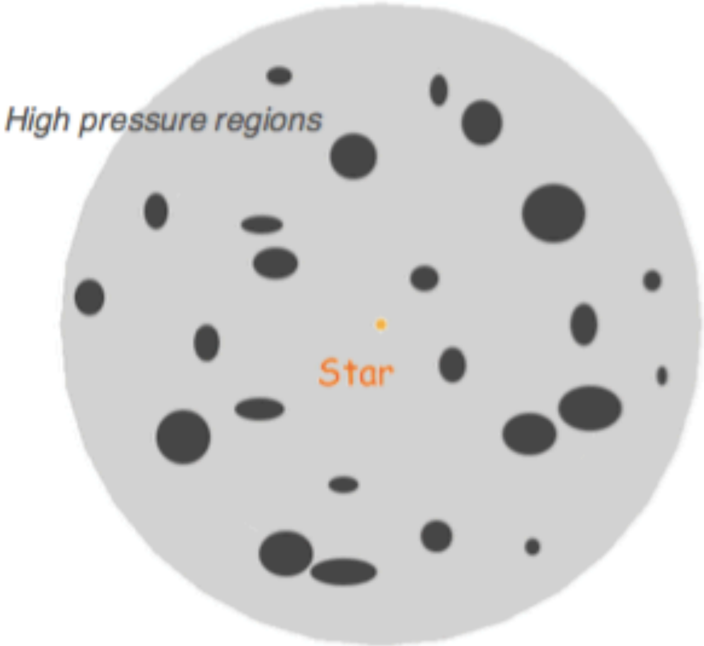
Beads exchange angular momentum

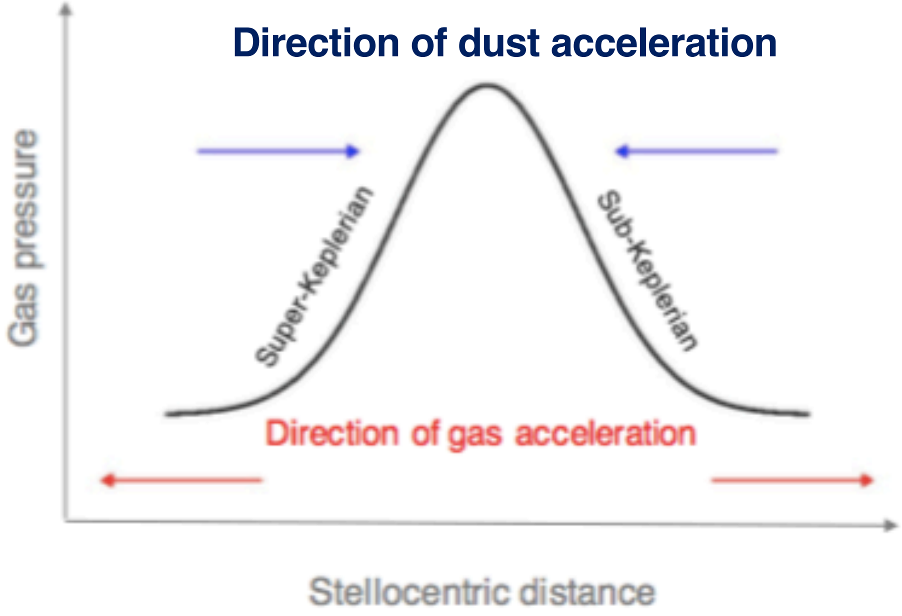
Magnetorotational Instability (MRI)



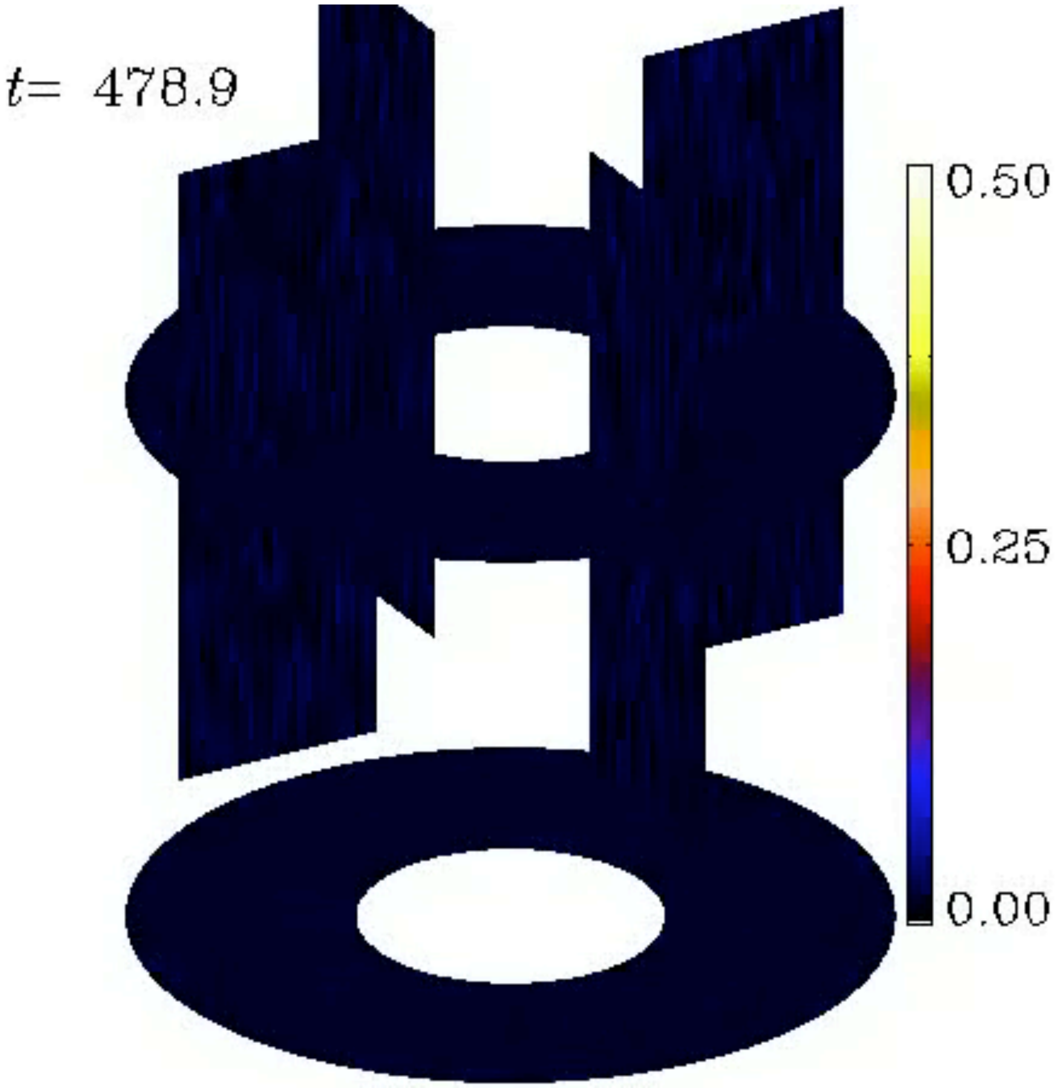
Magnetic fields
in a conducting rotating plasma behave
EXACTLY like ***springs***!

Turbulence


High pressure regions
Star

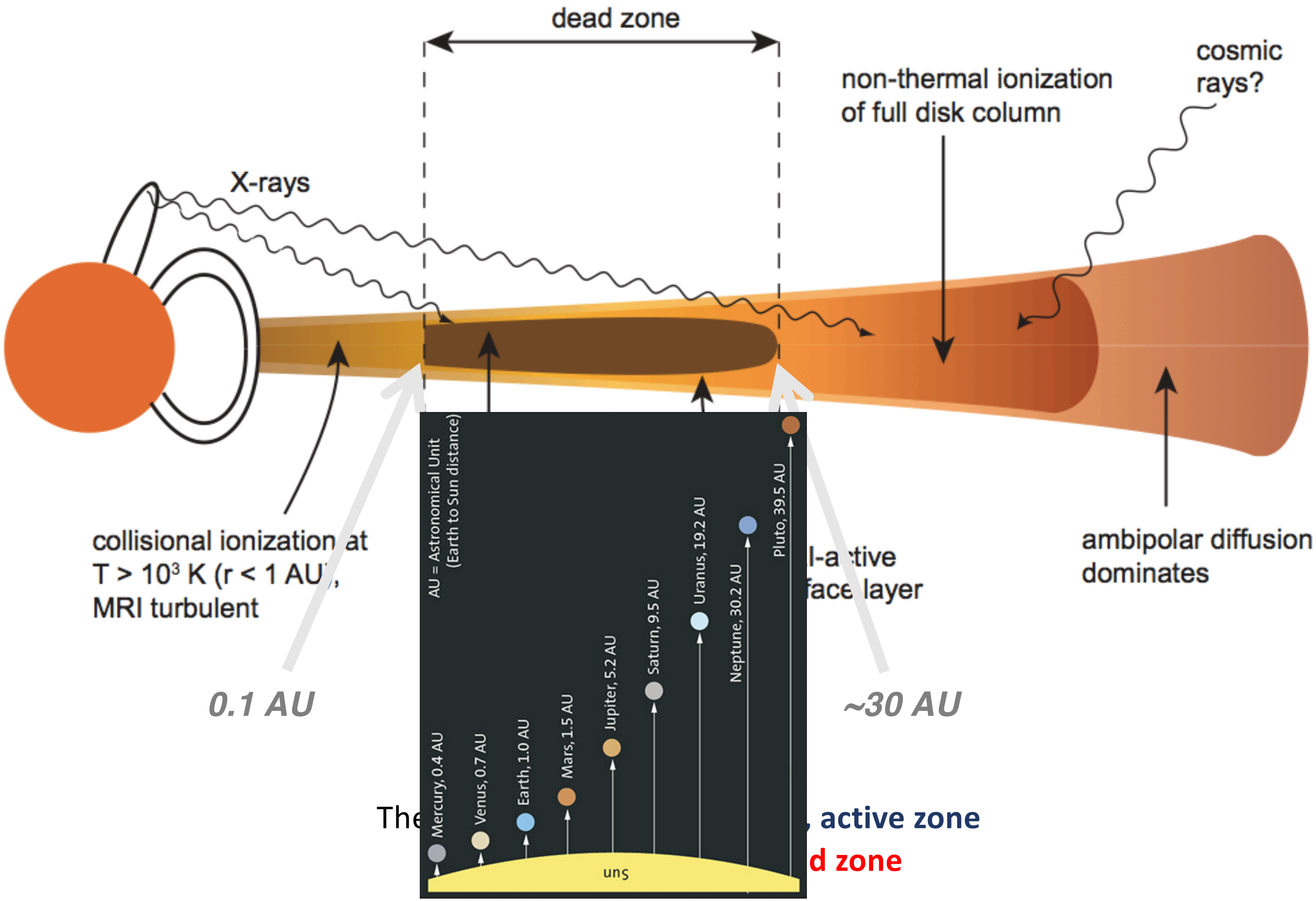

Direction of dust acceleration
Gas pressure
Super-Keplerian
Sub-Keplerian
Direction of gas acceleration
Stellocentric distance

Turbulence concentrates solids mechanically in pressure maxima



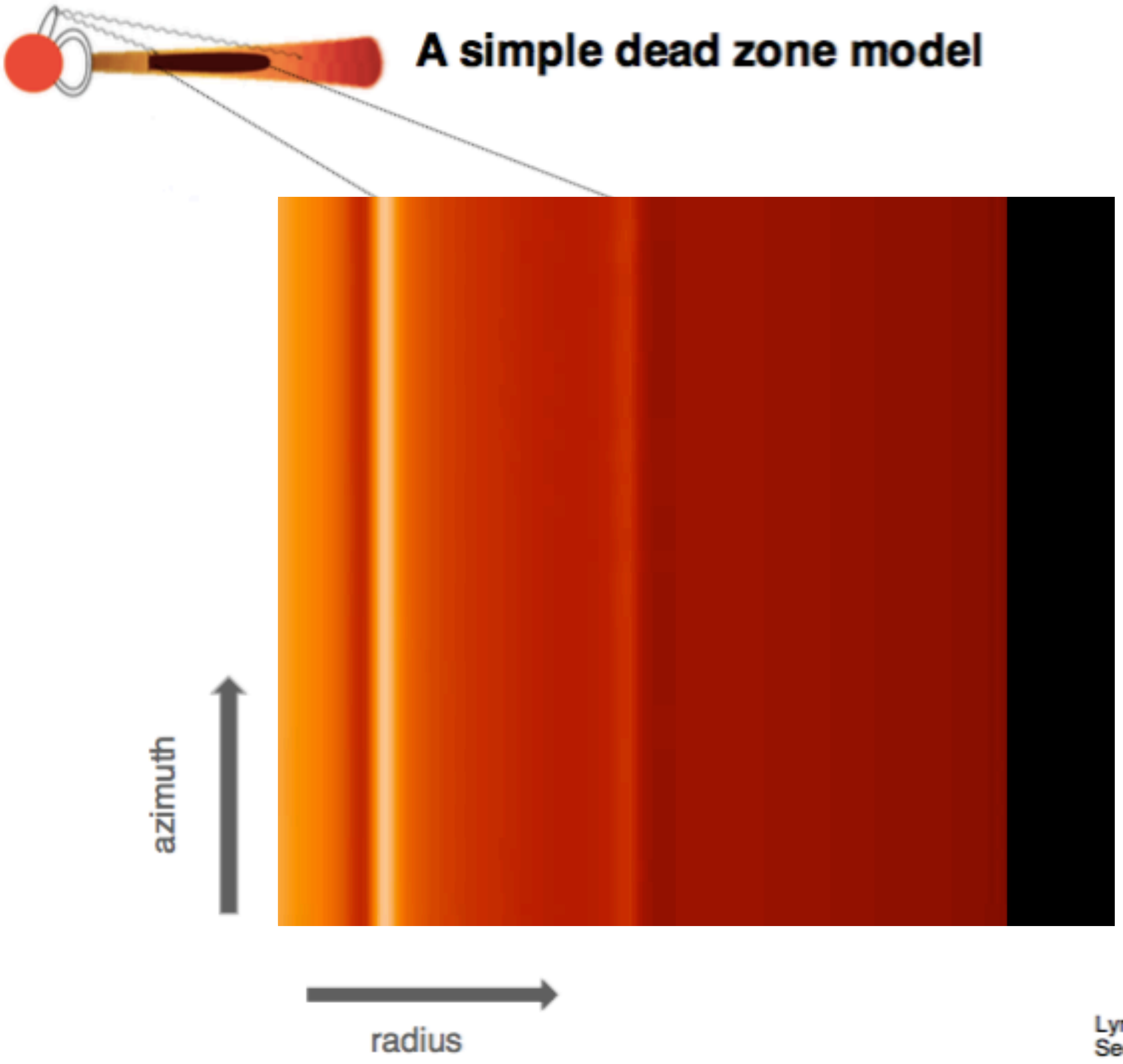
Lyra et al. (2008a)

Dead zones



The … , active zone
… d zone

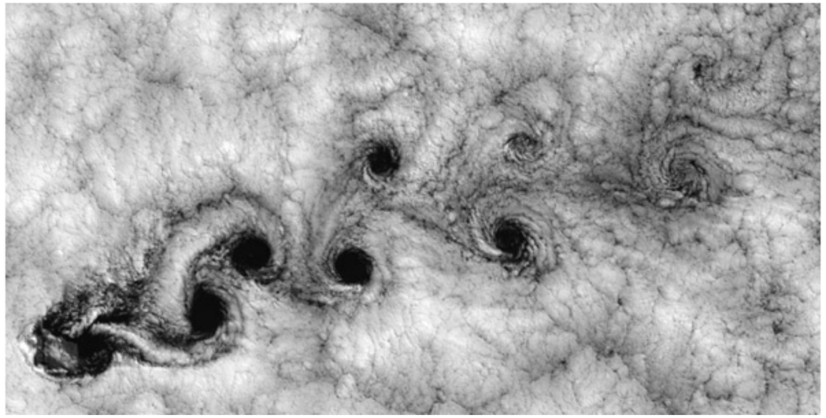
A simple dead zone model



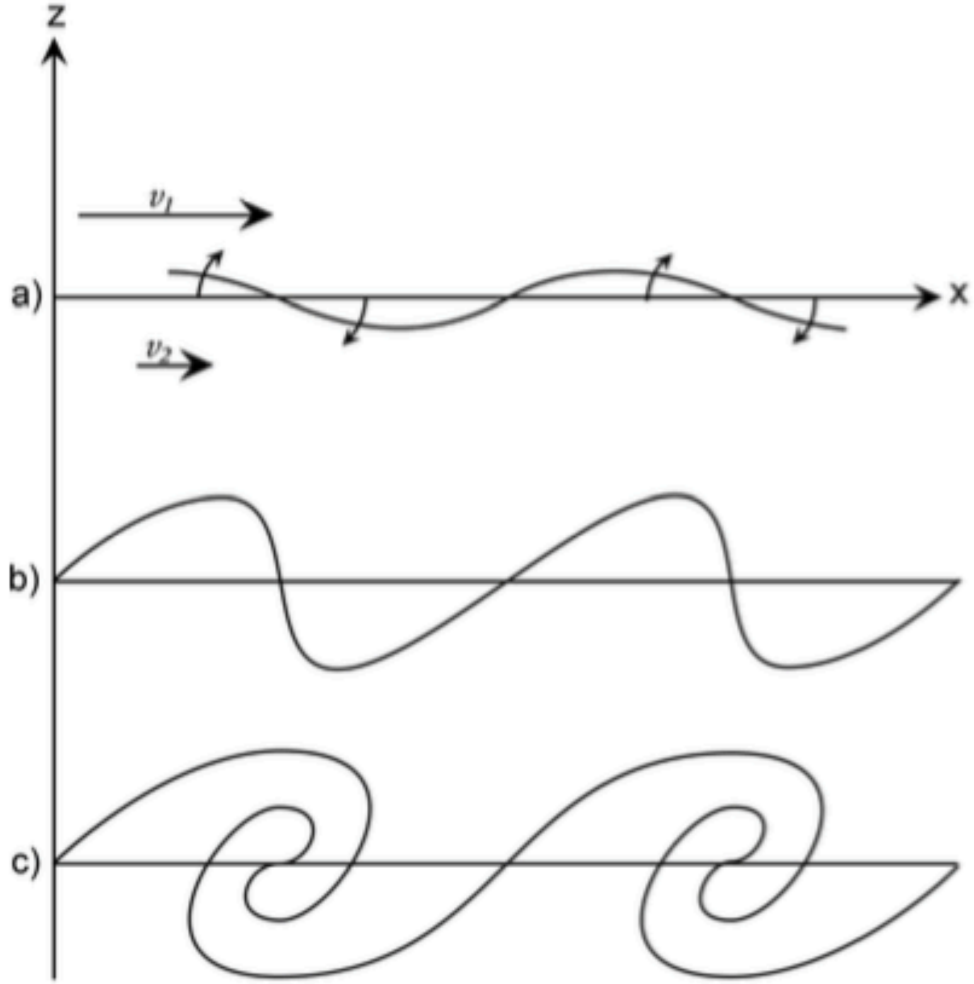
Lyra et al. (2008b, 2009a);
See also Varniere & Tagger (2006)

Vortices – an ubiquitous fluid mechanics phenomenon

Von Kármán *vortex street*

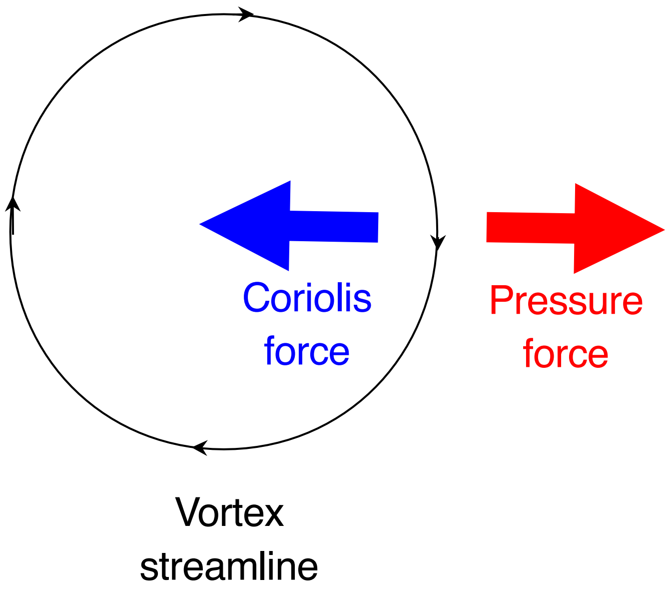


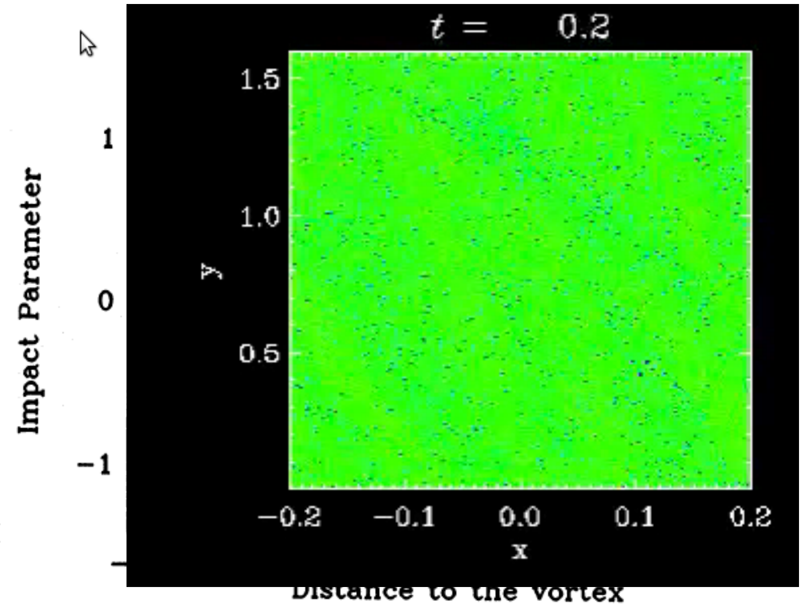
Rossby wave instability
(Kelvin-Helmholtz Instability in rotating disks)



Vortex Trapping

Geostrophic balance:





Barge & Sommeria (1995)

Grains do not feel the pressure gradient.
They sink towards the center, where they accumulate.

Aid to planet formation

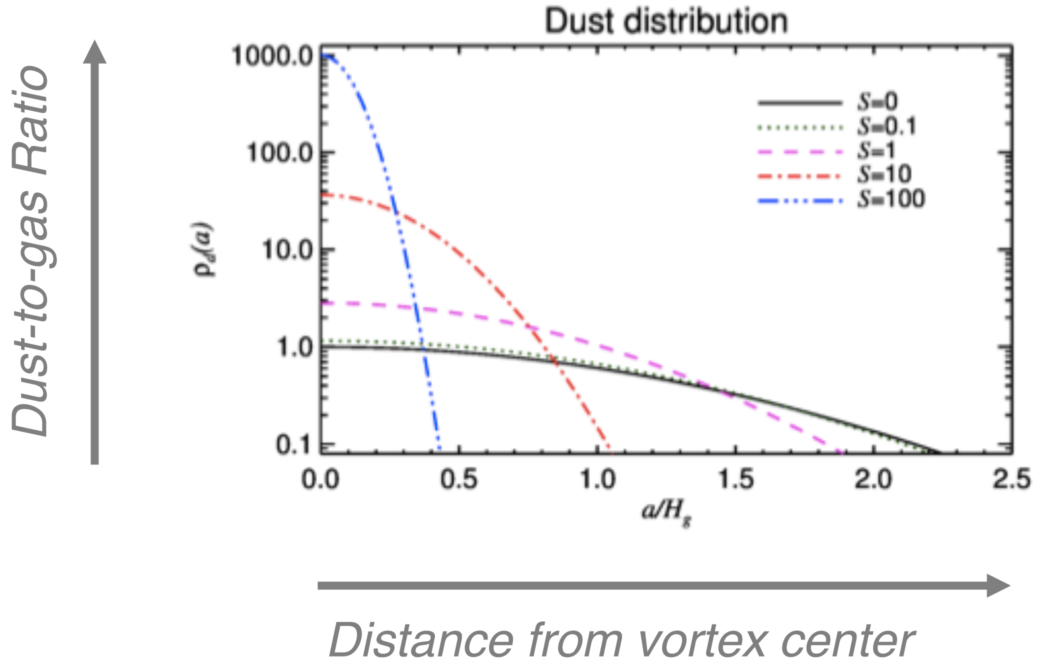
(Barge & Sommeria 1995, Tanga et al. 1996, Barranco & Marcus 2005)

Speeds up planet formation enormously

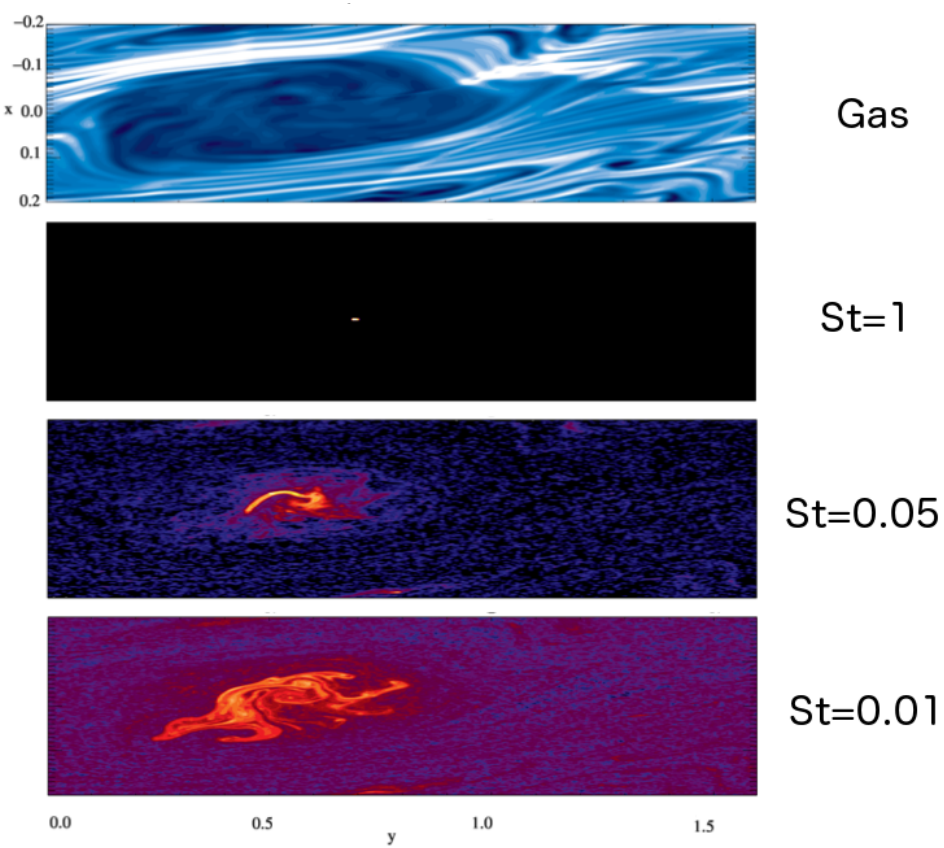
(Lyra et al. 2008b, 2009ab, Raettig et al. 2012)

Strong clumping

Easily reaches dust-to-gas ratio > 1
even for solar (and sub-solar) metallicities.

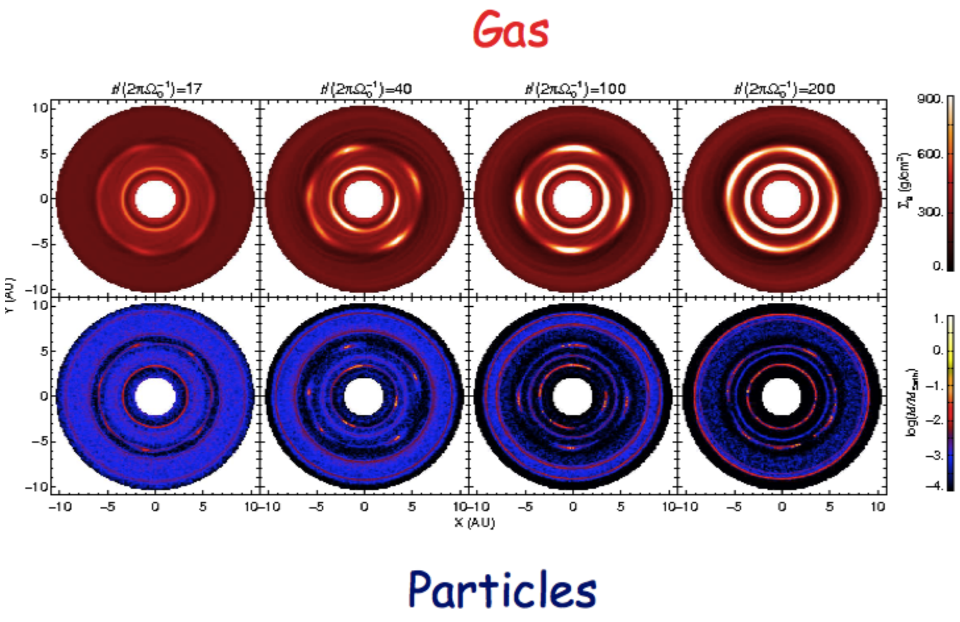


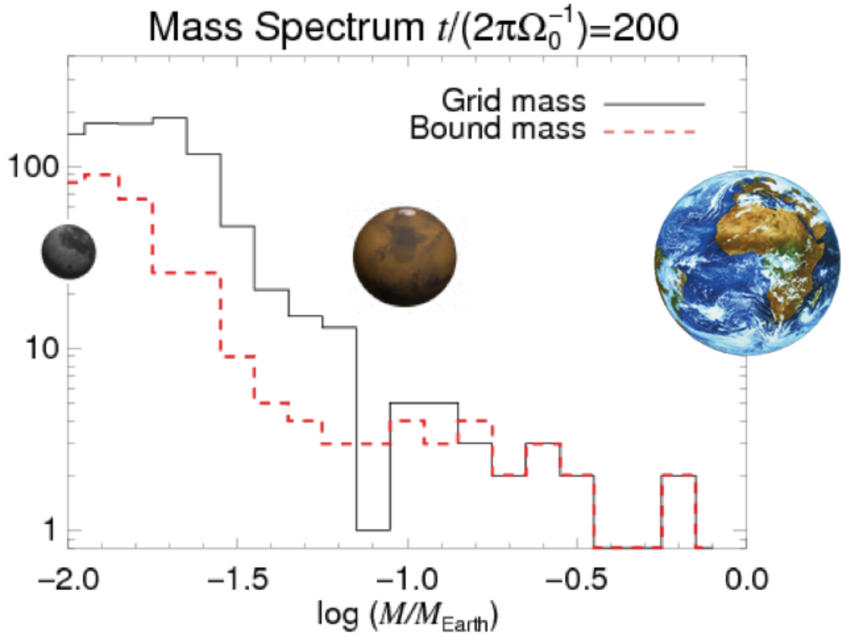
Lyra & Lin (2013)



Raettig et al (2015)

Vortices and Planet Formation





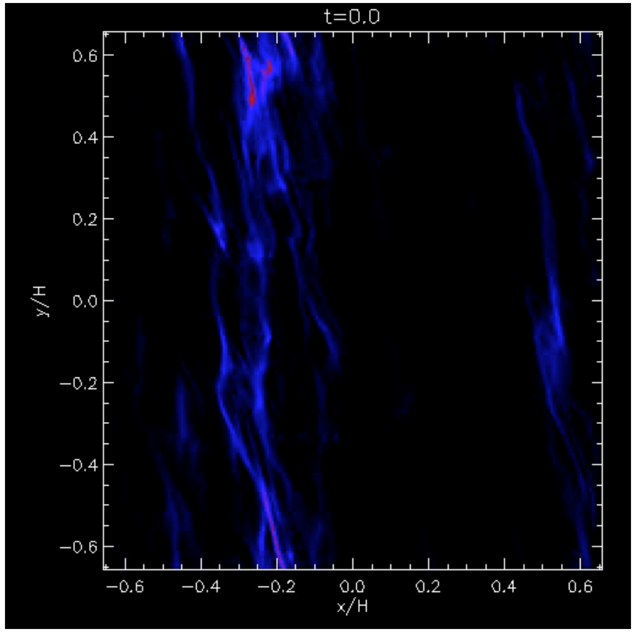
Collapse into Mars mass objects

(Lyra et al. 2008b, 2009a,
Lambrechts & Johansen 2012)

Take home message

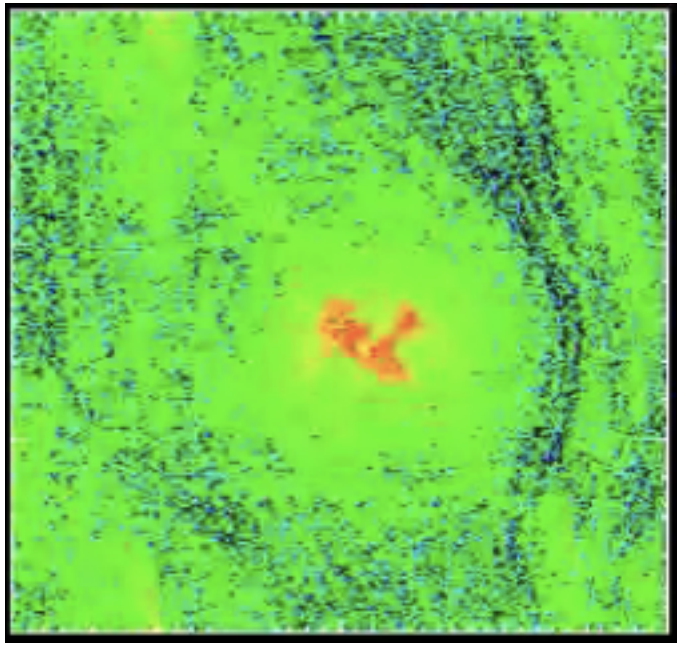
- Two routes for planet formation

Streaming Instability



Johansen+ 07

Vortex Trapping

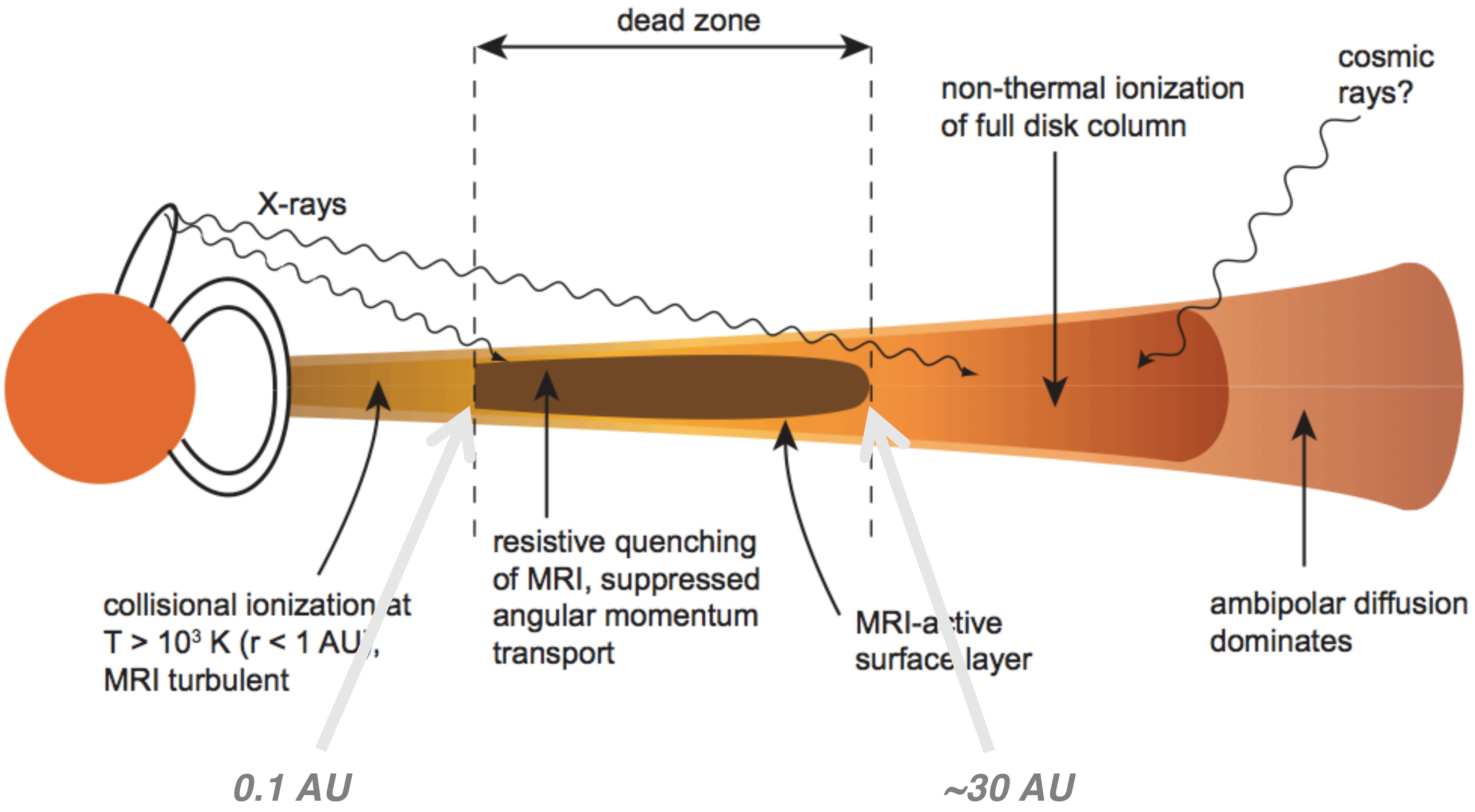


Lyra+08, Raettig+Lyra 12

- Planet formation and turbulence.
    - Does turbulence help (concentration at large scales) or hinder (diffusion at small scales)?

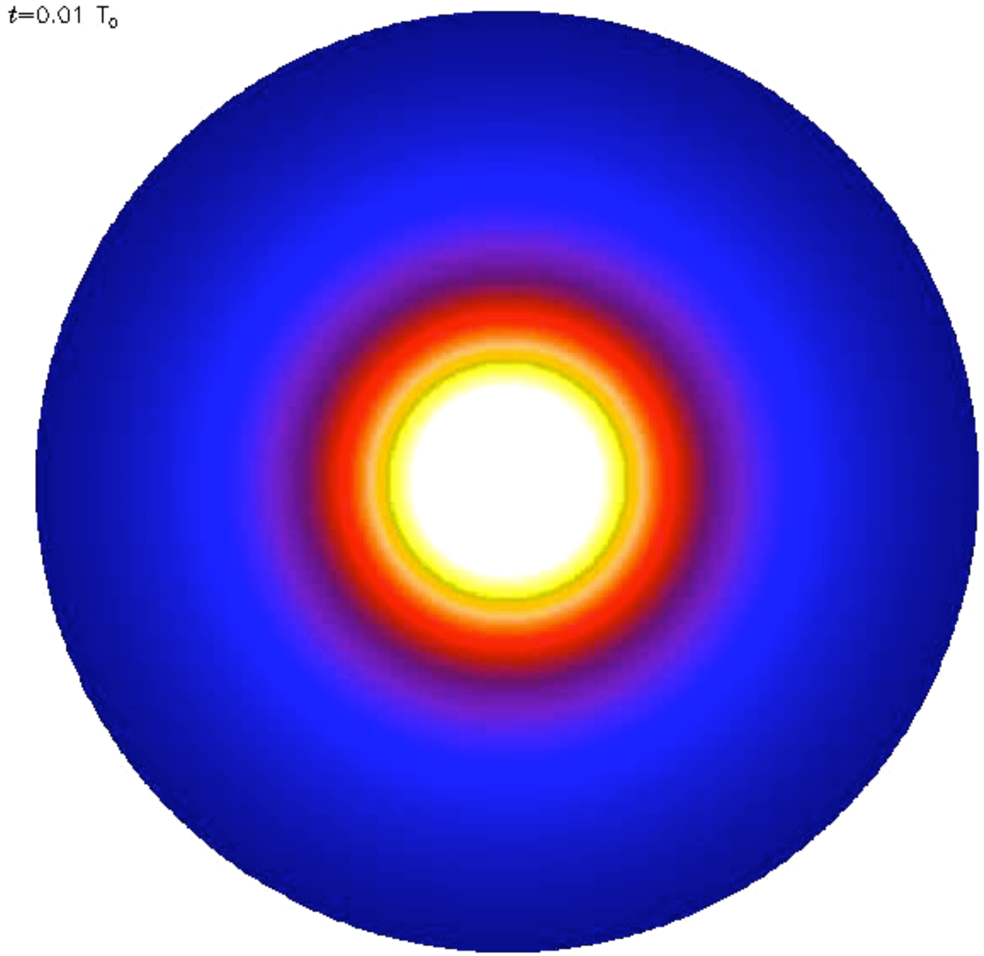
Part 2: Disk Instabilities

Dead zones



There should be a **magnetized, active zone**
and a **non-magnetic, dead zone**

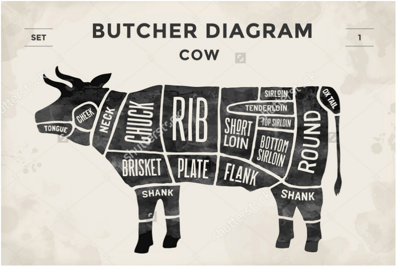
Inner (0.1 AU) active/dead zone boundary

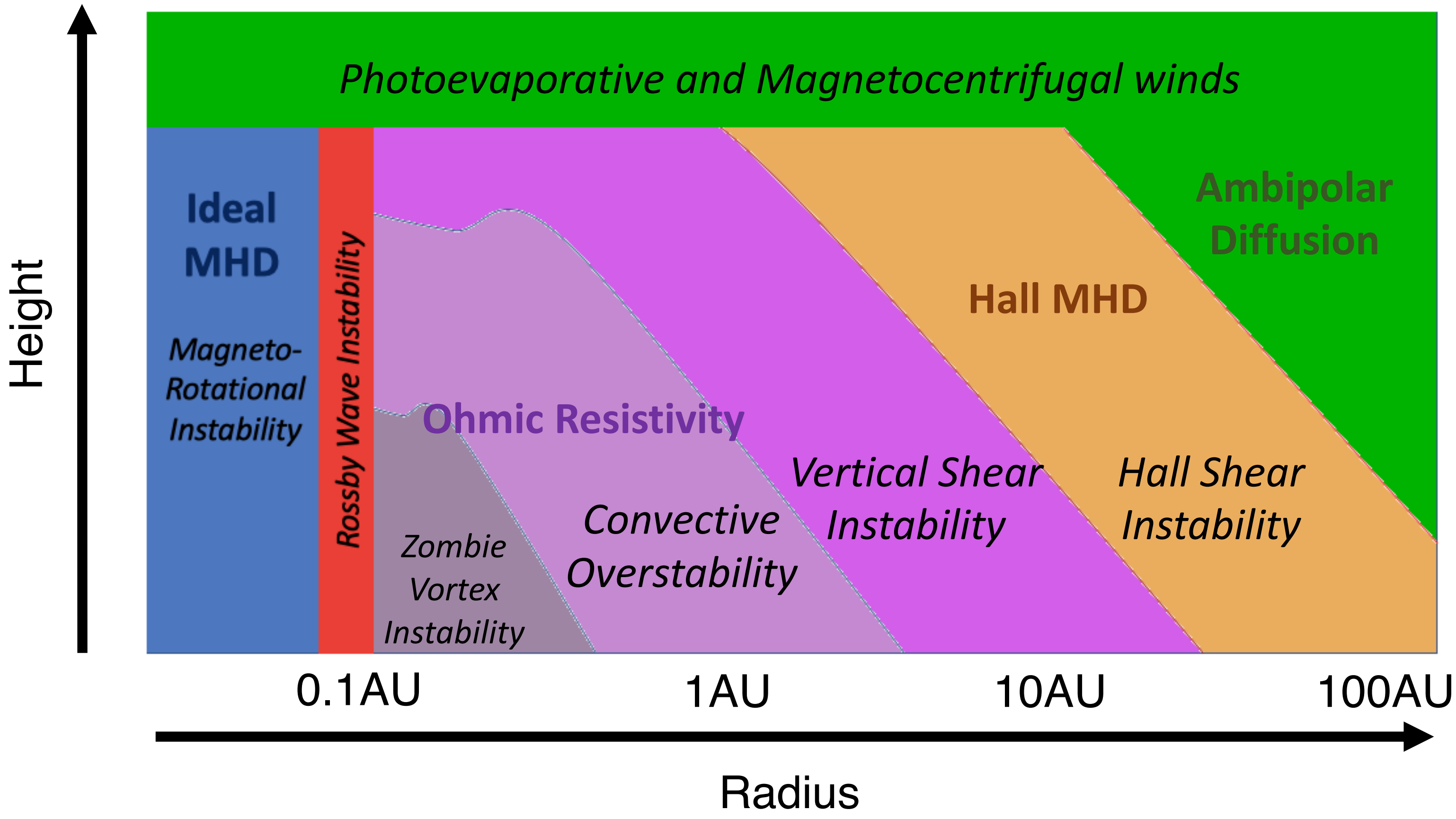


Magnetized inner disk + resistive outer disk

Lyra & Mac Low (2012)

Instability map

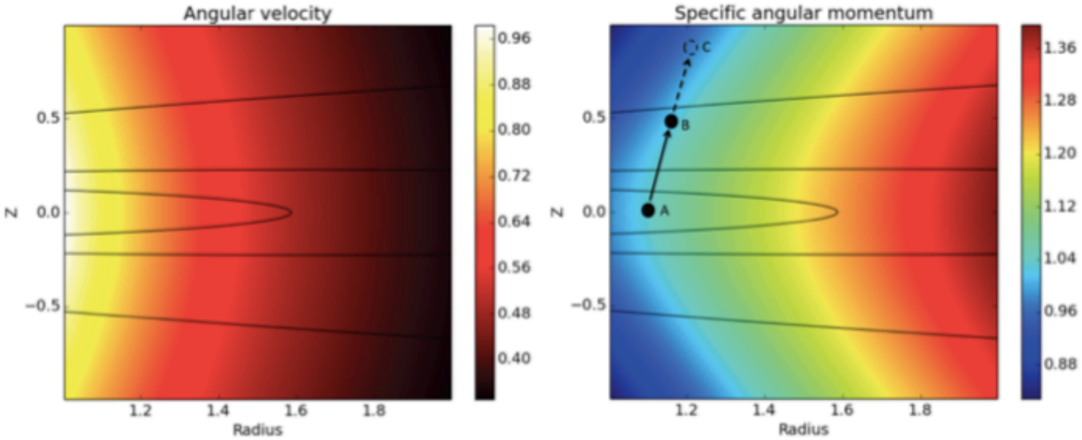




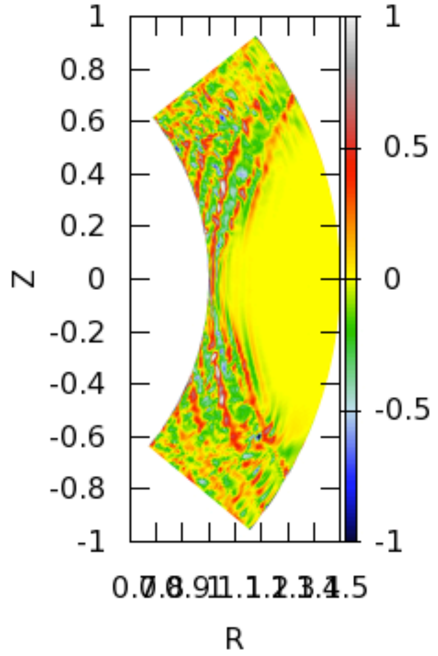
Vertical shear instability

Angular velocity not constant in cylinders: unstable

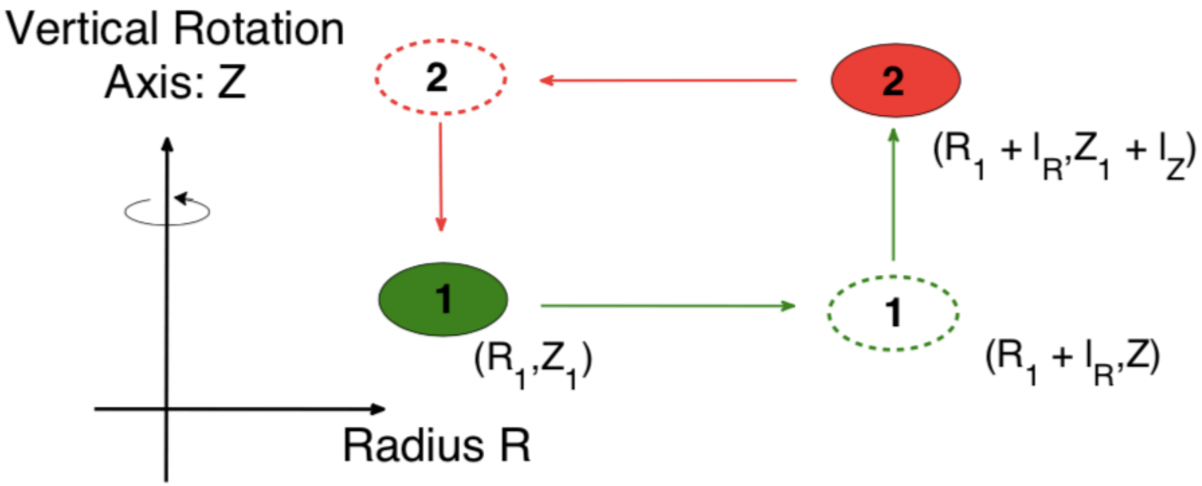
Buoyancy stabilizes! The most unstable mode is isothermal.



Fromang & Lesur (2017)

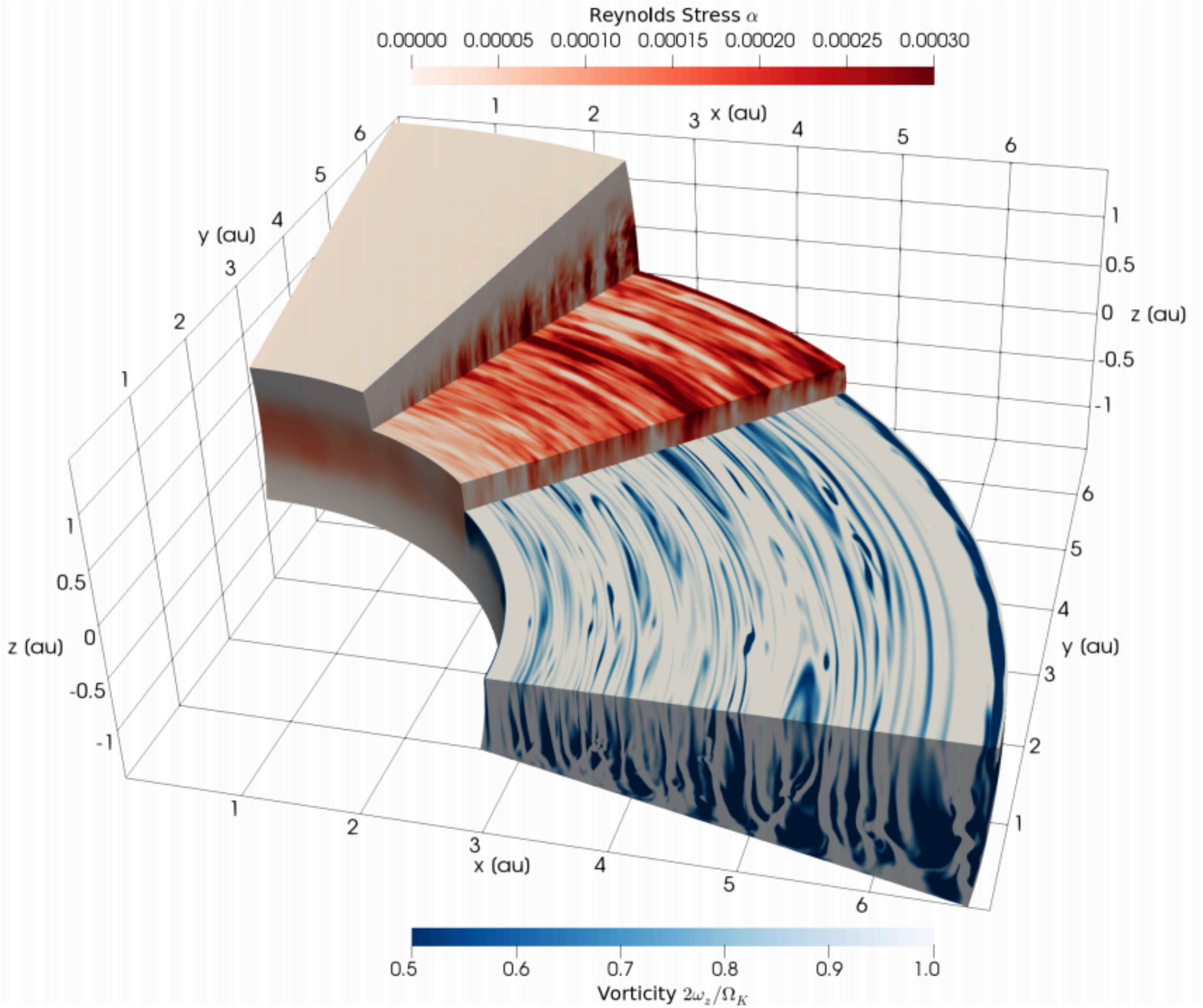


Nelson et al. (2013)



(see also Glen’s talk)

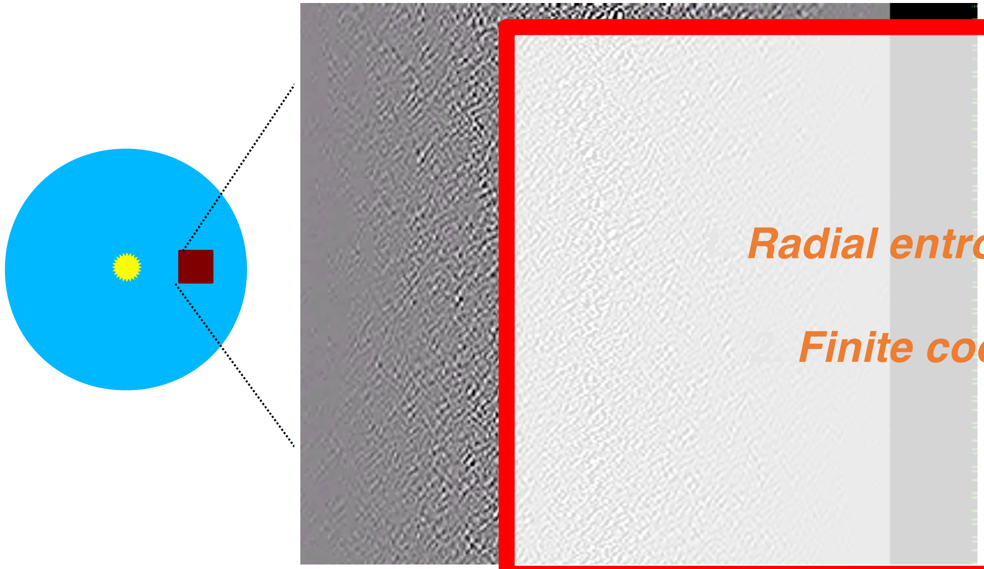
Vertical shear instability



Pfeil & Klahr (2020)

Convective Overstability

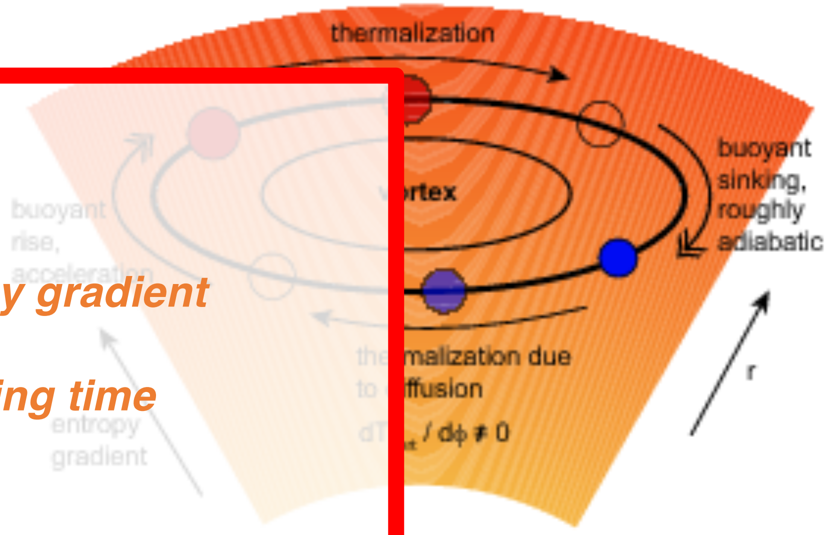
Sketch of the
Convective Overstability



Radial entropy gradient

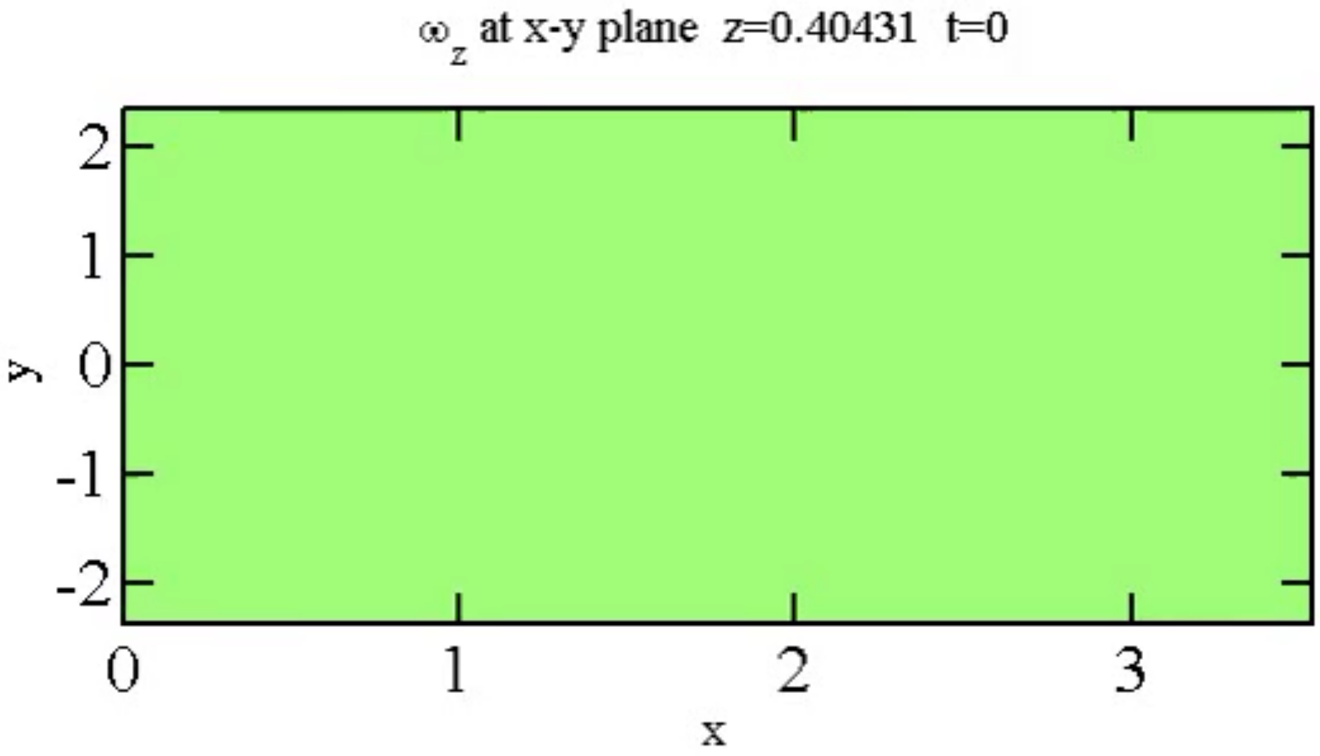
Finite cooling time

Lesur & Papaloizou (2010)
Lyra & Klahr (2011)



Armitage (2010)

Zombie Vortex Instability

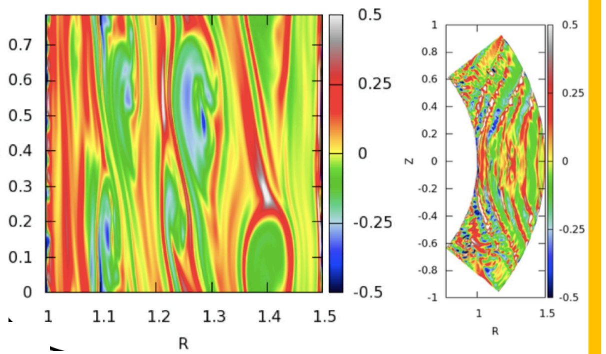


Cascade of baroclinic critical layers

Marcus et al. (2015, 2016)

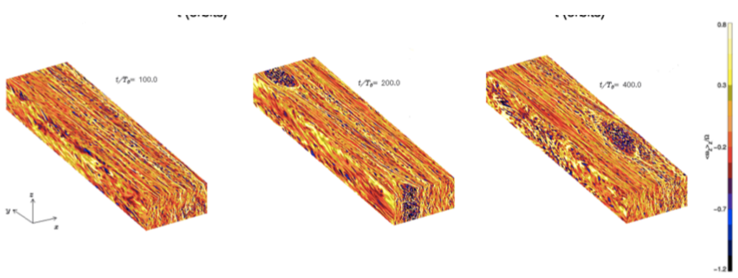
Hydrodynamical Instabilities

Vertical Shear Instability



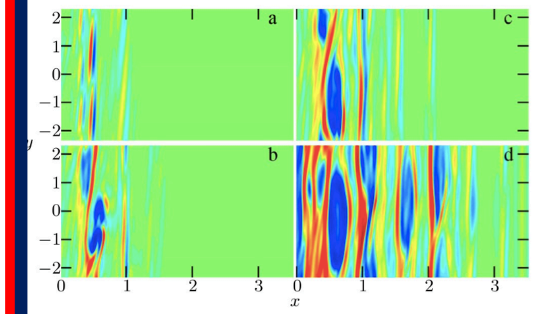
Nelson et al. 2013, Lin & Youdin 2015, Umurhan et al. 2016

Convective Overstability



Klahr 2003, Lesur & Papaloizou 2010, Lyra & Klahr 2011, Lyra 2014

Zombie Vortex Instability



Marcus et al. 2012, 2013, 2015, 2016
Umurhan et al 2016, Lesur & Latter 2016

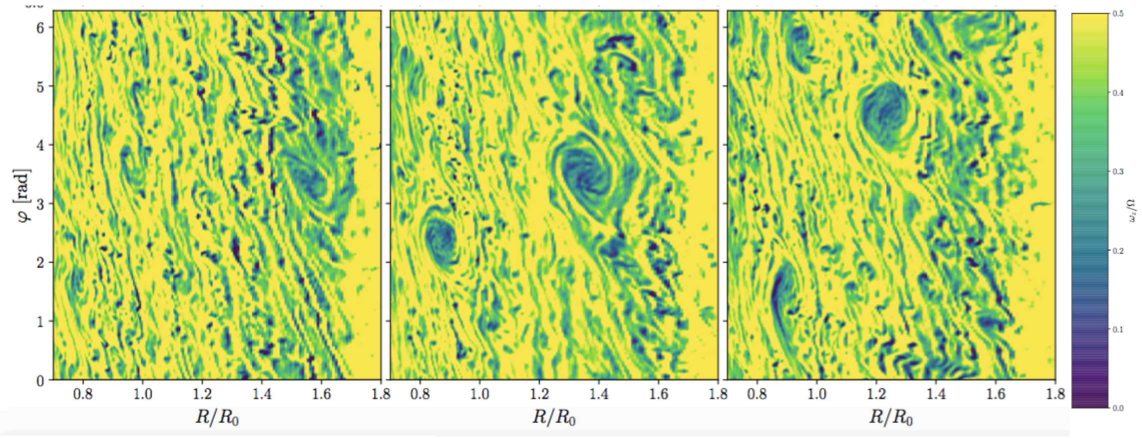
$\Omega\tau << 1$
($\kappa < 1$ cm^2/g)

$\Omega\tau \sim 1$
($\kappa \sim 1$–50 cm^2/g)

$\Omega\tau >> 1$
($\kappa > 50$ cm^2/g)

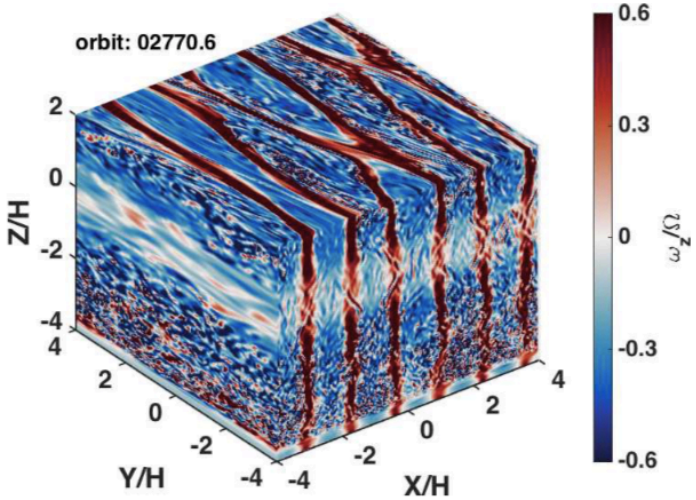
Opacity

Take-home message



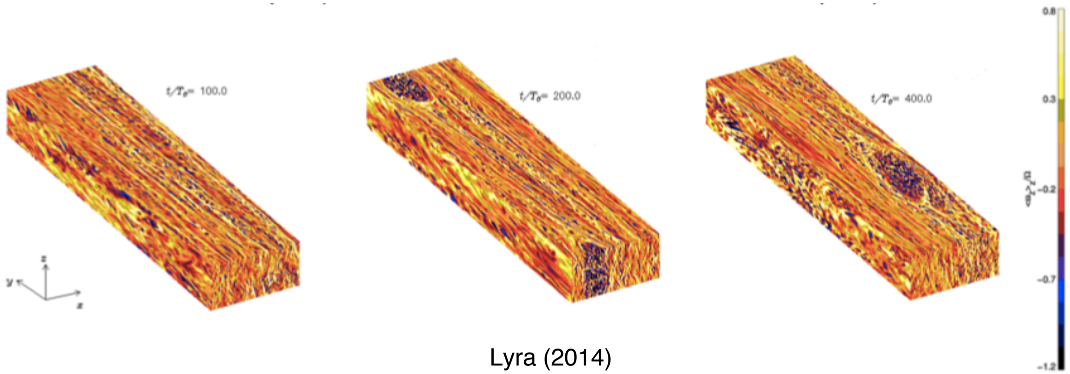
Manger & Klahr (2018)

Vertical Shear Instability
saturates into
vortices



Barranco et al. (2019)

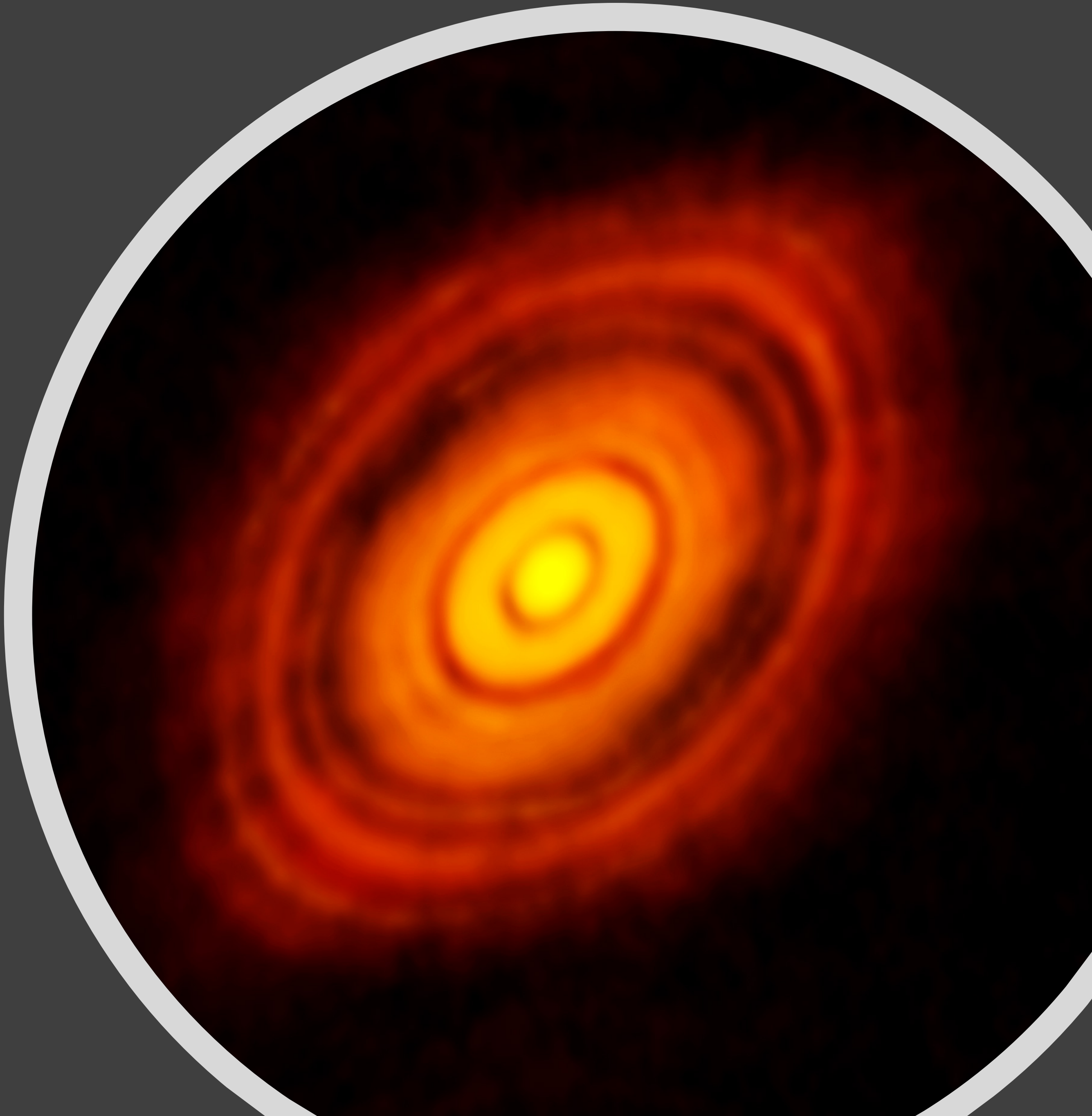
Zombie Vortex Instability
saturates into
vortices

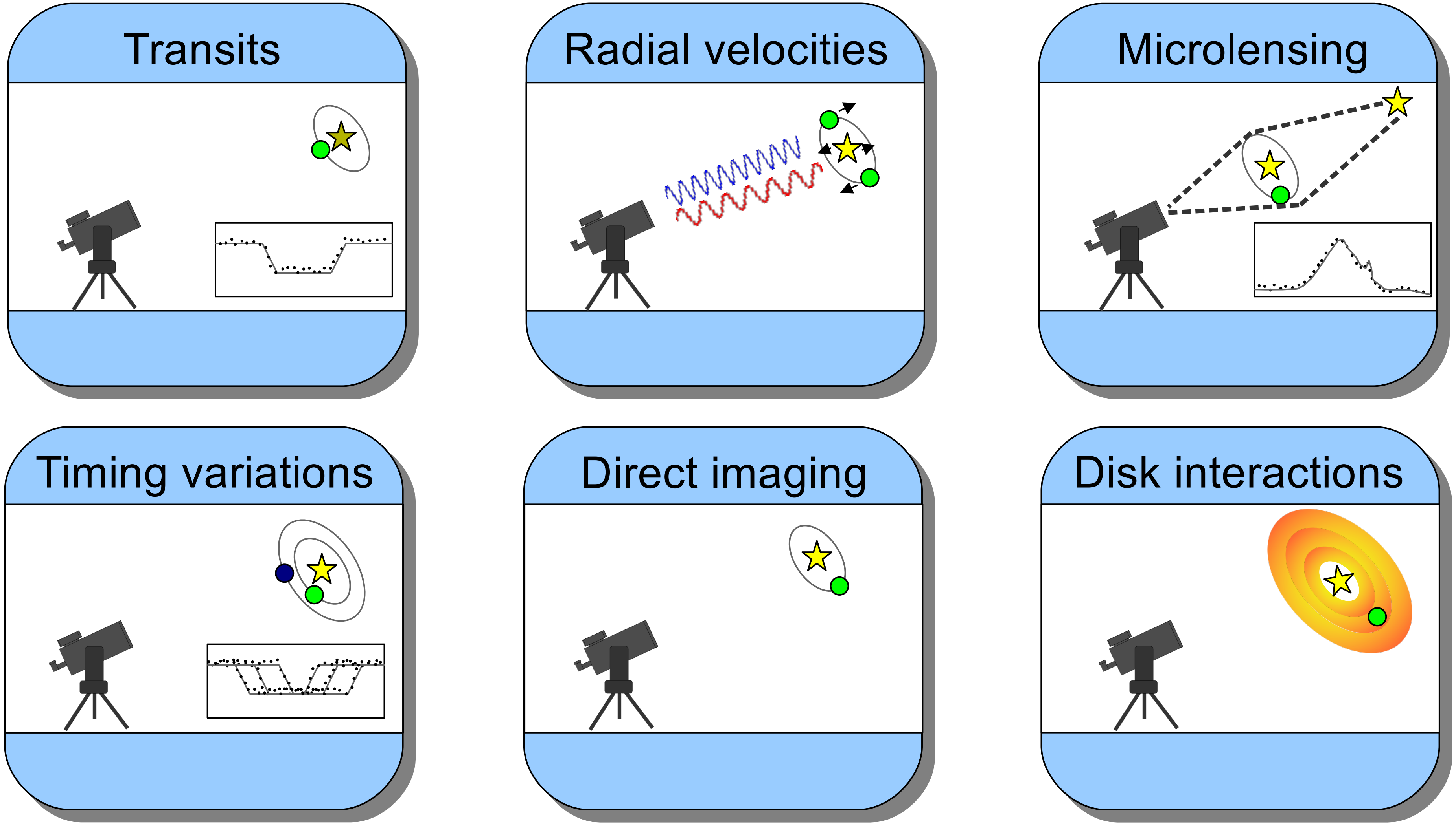


Lyra (2014)

Convective Overstability
saturates into
vortices

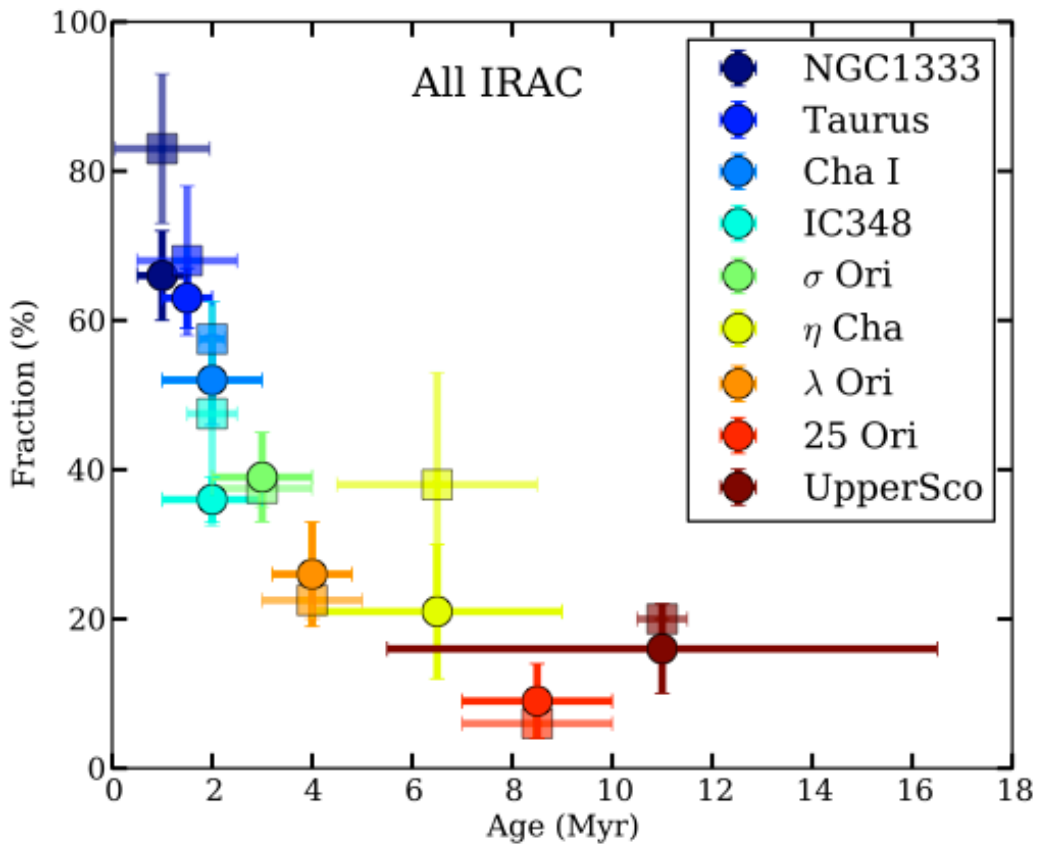
Part 3: Disk Observations





Slide by A. Richert

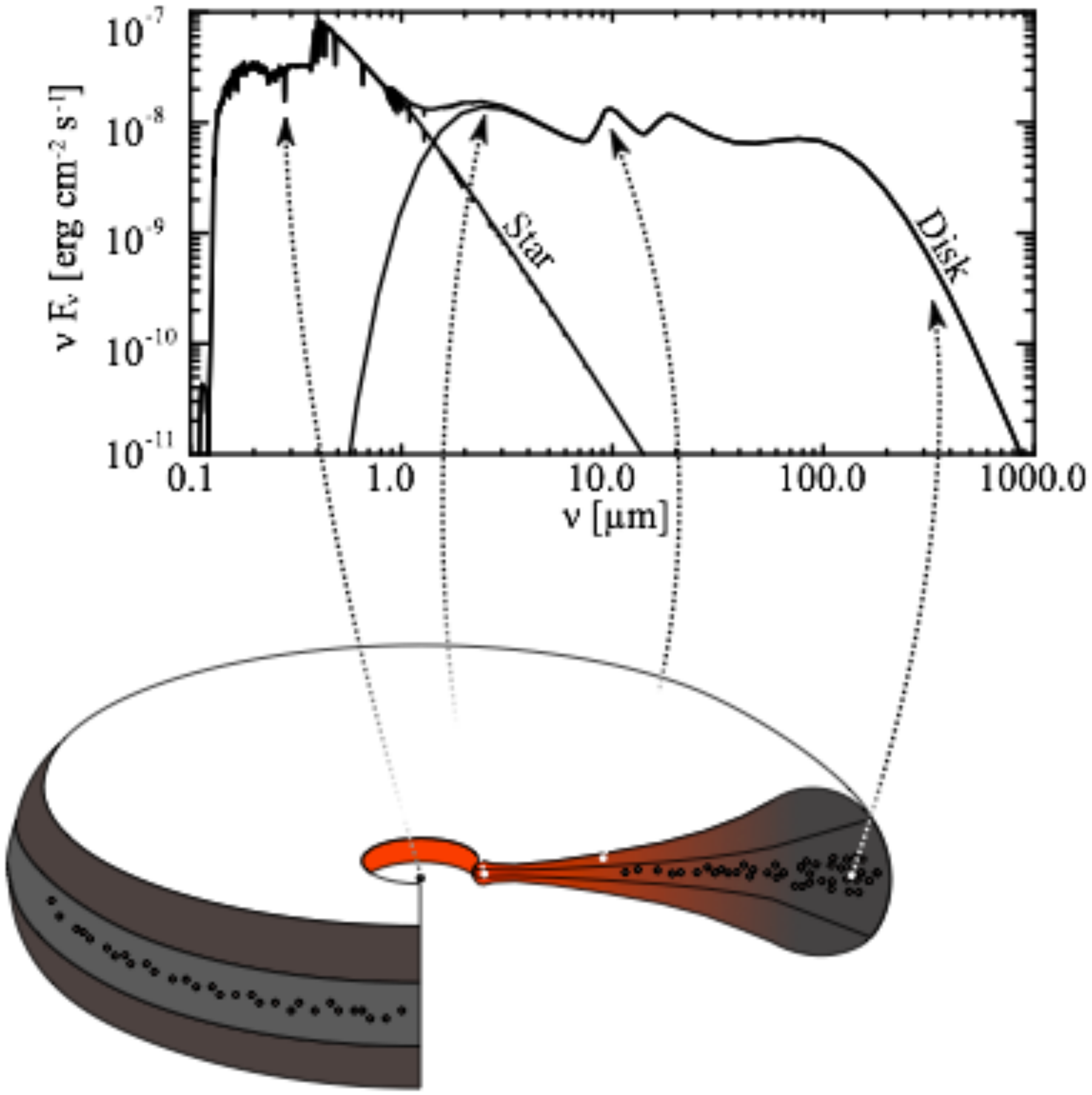
Disk lifetime



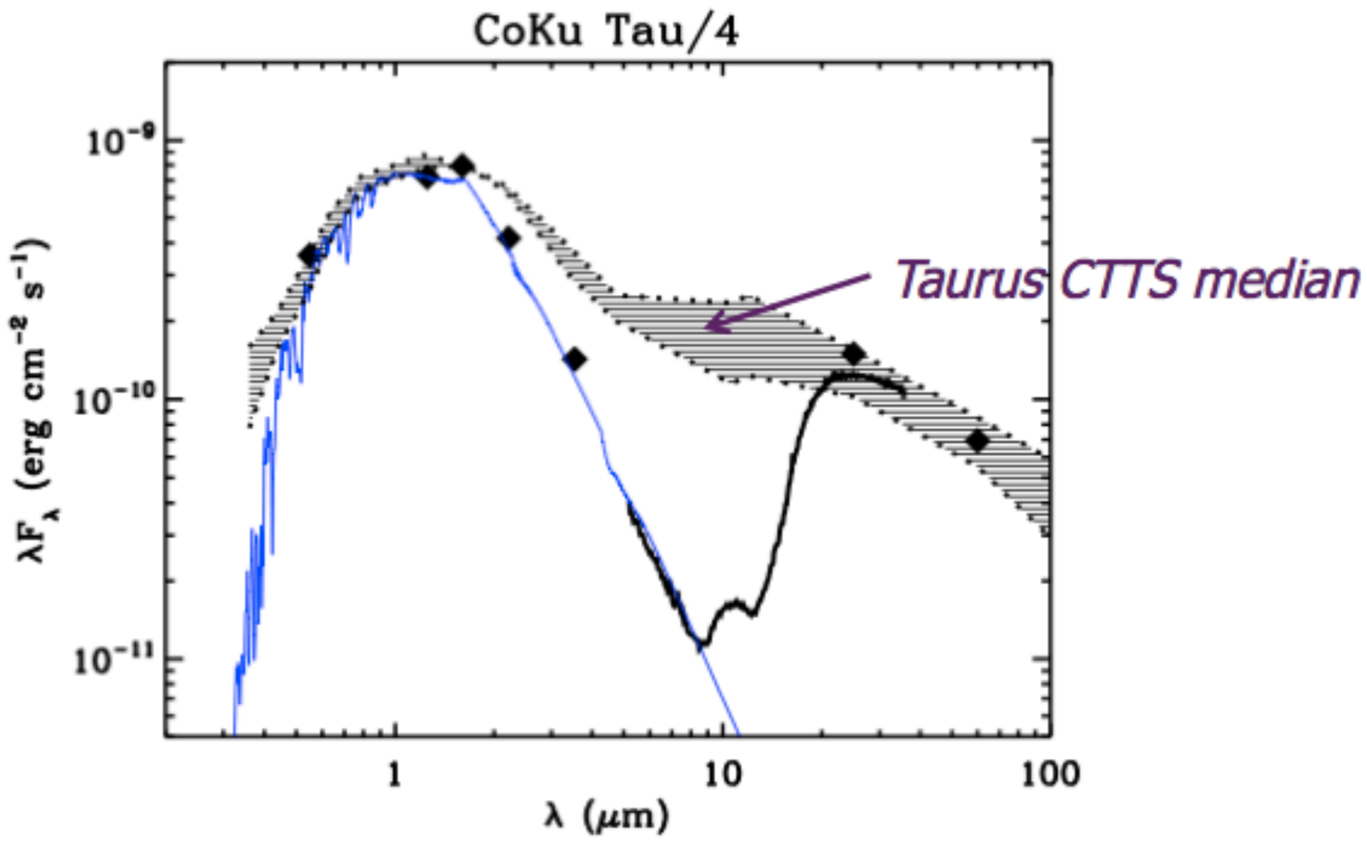
(Ribas et al. 2014)

Disks dissipate within ~10Myr

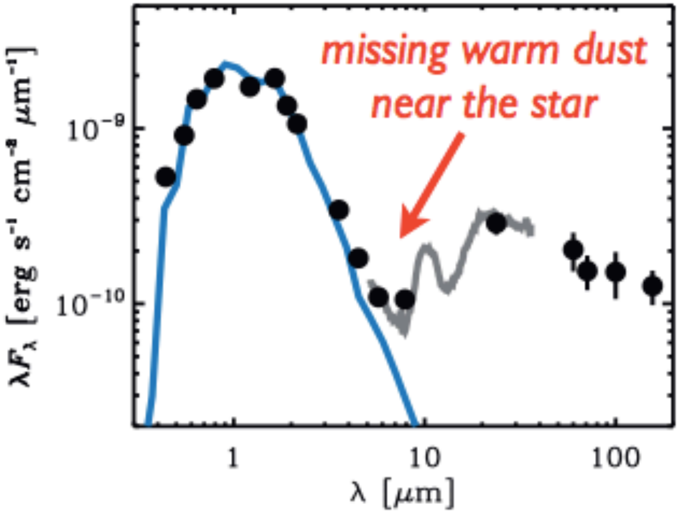
Disk spectra



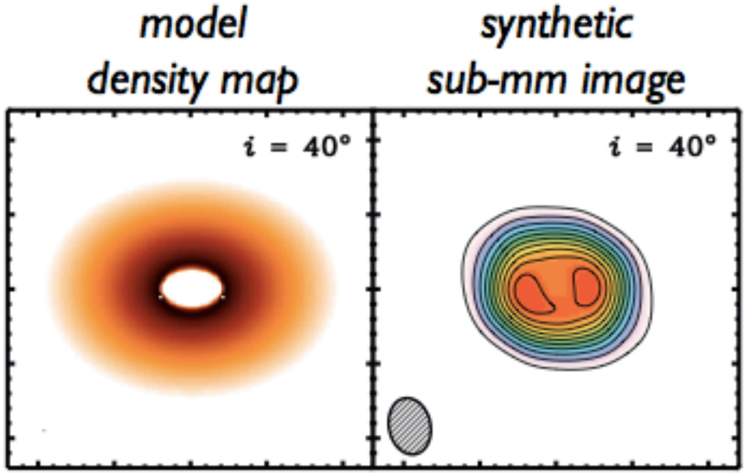
A class of disks with missing hot dust.



Disks with missing hot dust.

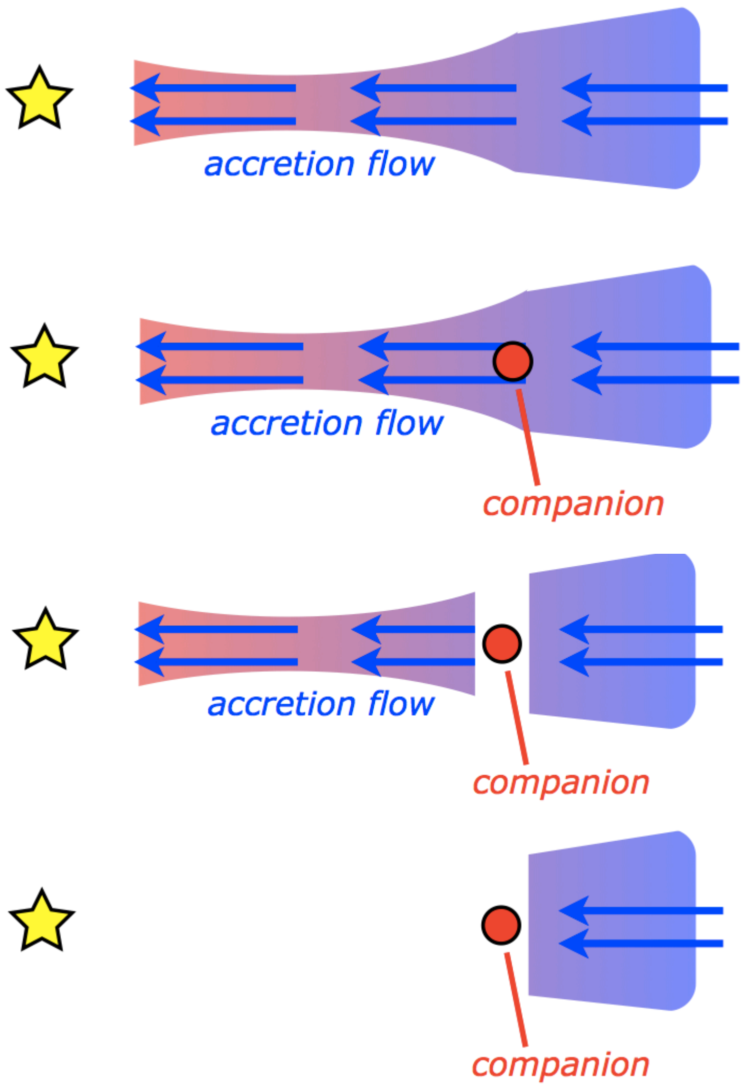


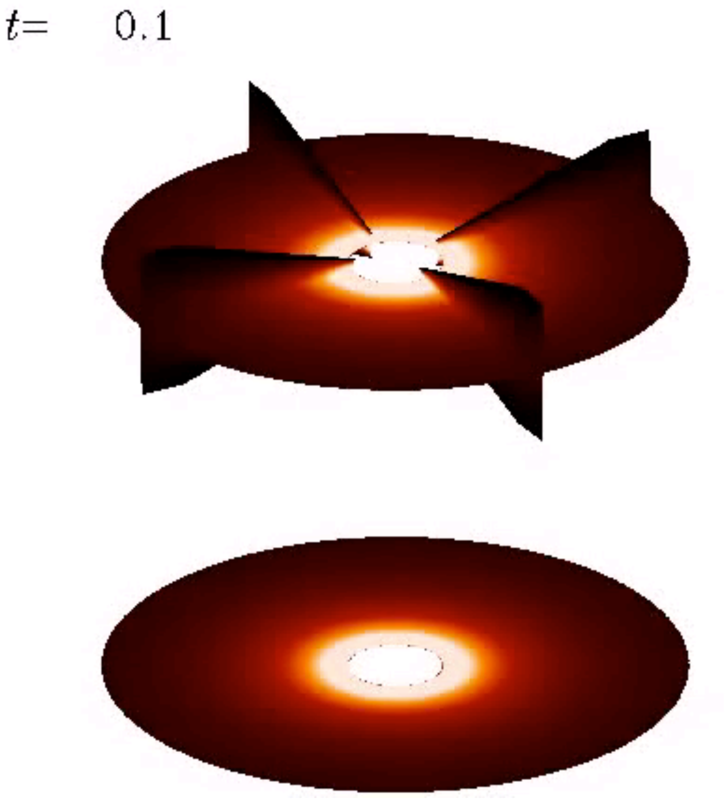
[e.g., Furlan et al. 2009]



a disk with a large reduction
in optical depth near the star
(i.e., a "cavity" or "hole")

Planetary companion

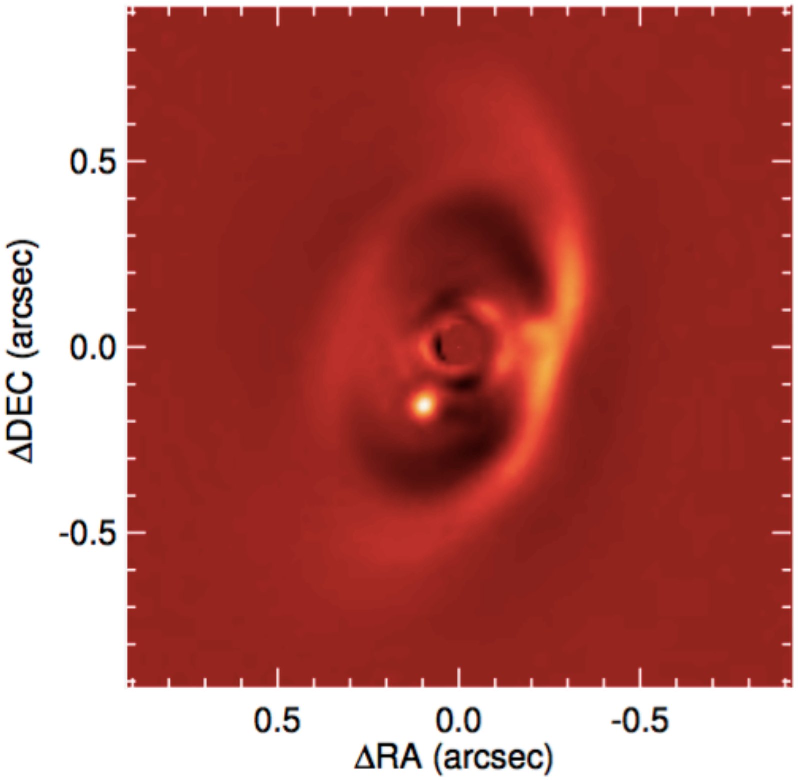




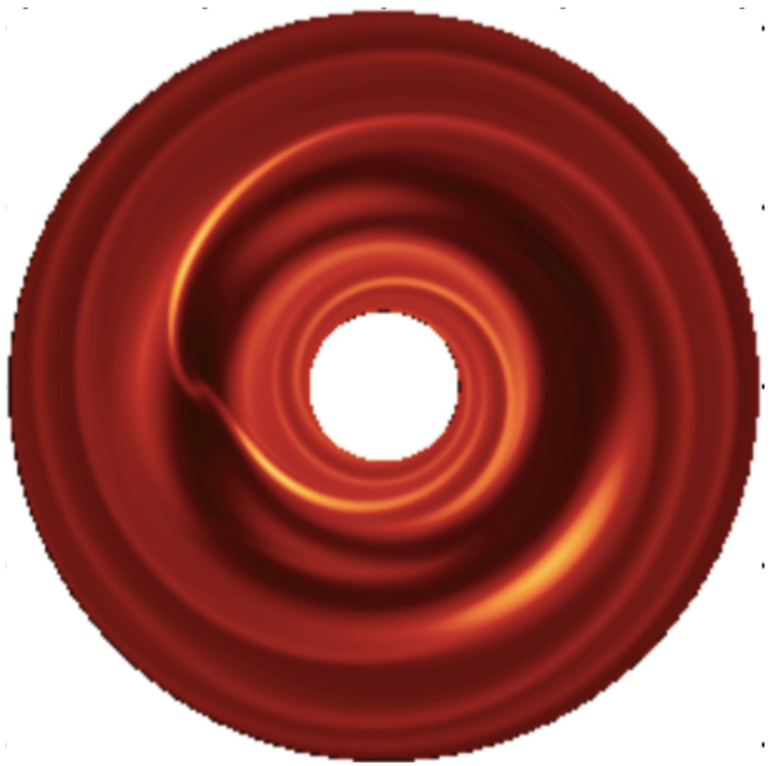
(Lyra 2009)

These cavities may be the telltale signature of forming planets

PDS 70 and PDS 70b



(Muller et al. 2018)



(Lyra et al. 2009b)

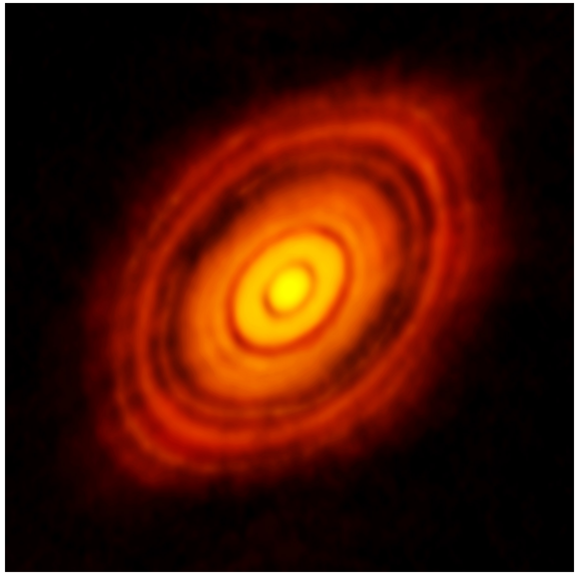
A way to directly study planet-disk interaction

Planet-disk interaction: gaps, spirals, and vortices.

(Lyra et al. 2009b)

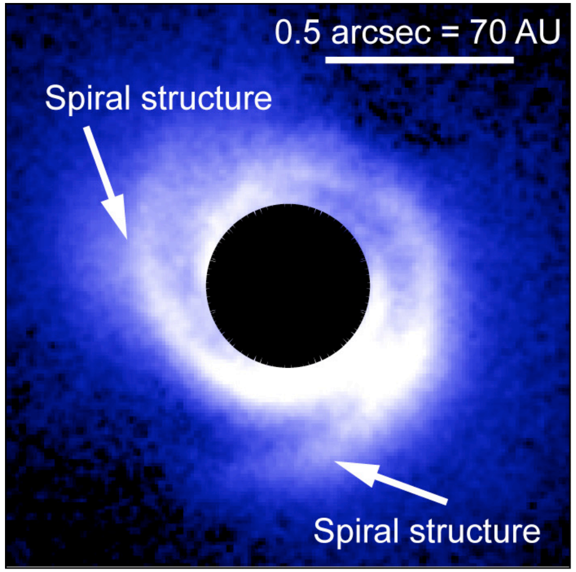
Observational evidence: gaps, spirals, and vortices

HL Tau



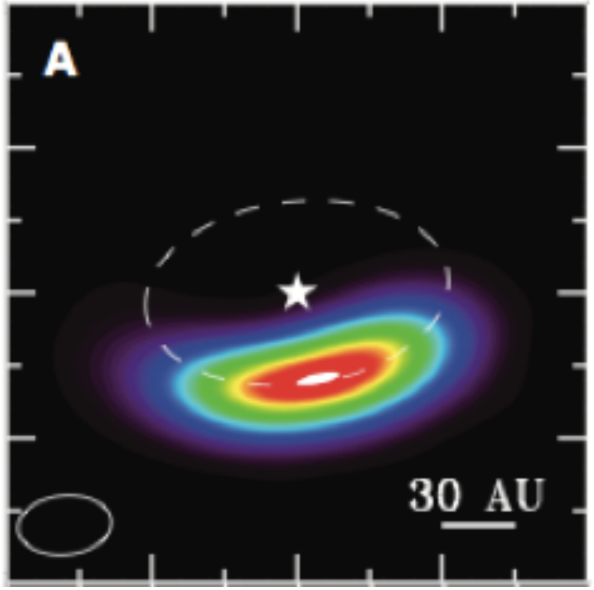
The ALMA Partnership et al. (2015)

SAO 206462



Muto et al. (2012)

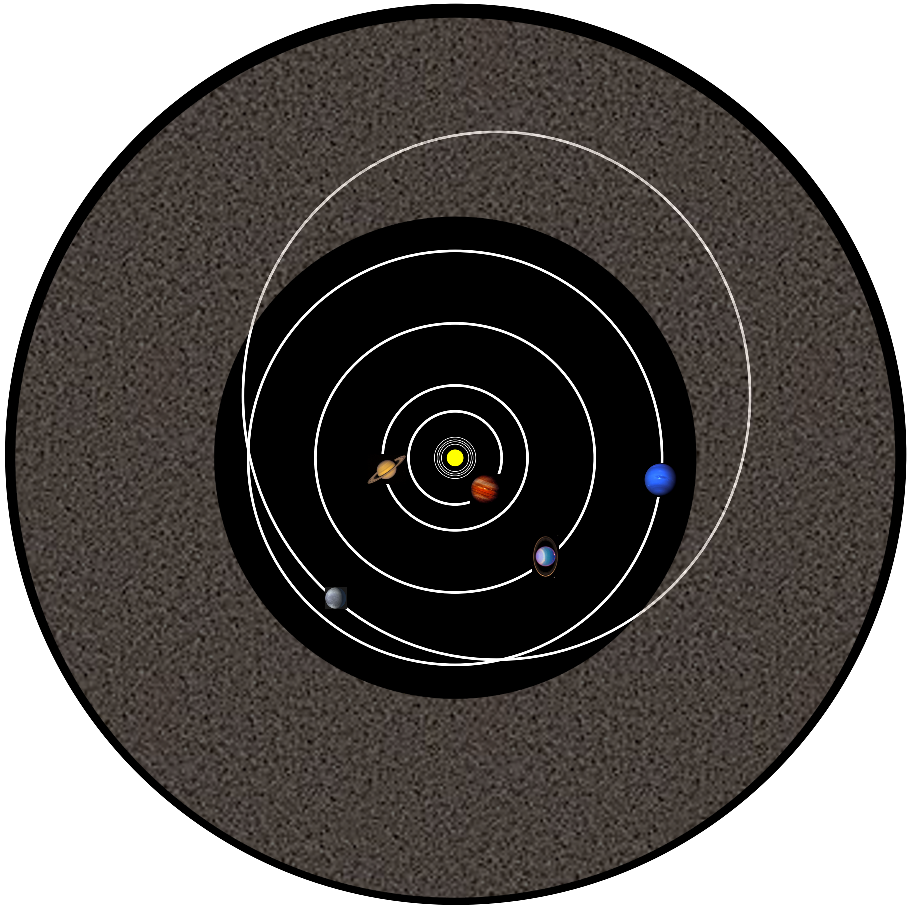
Oph IRS 48

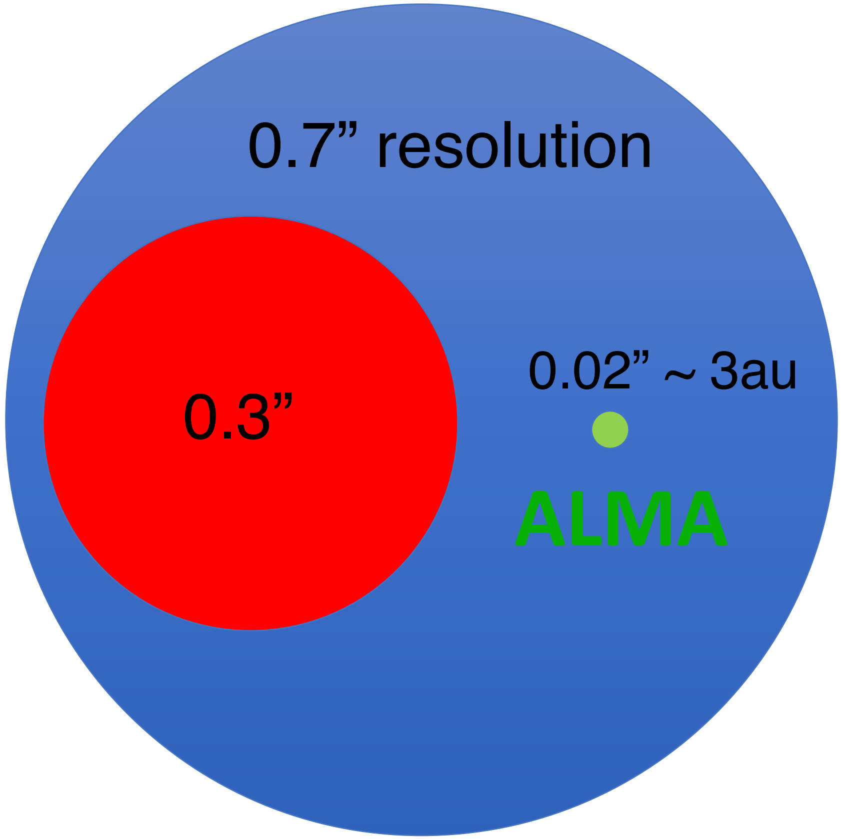


van der Marel et al. (2013)

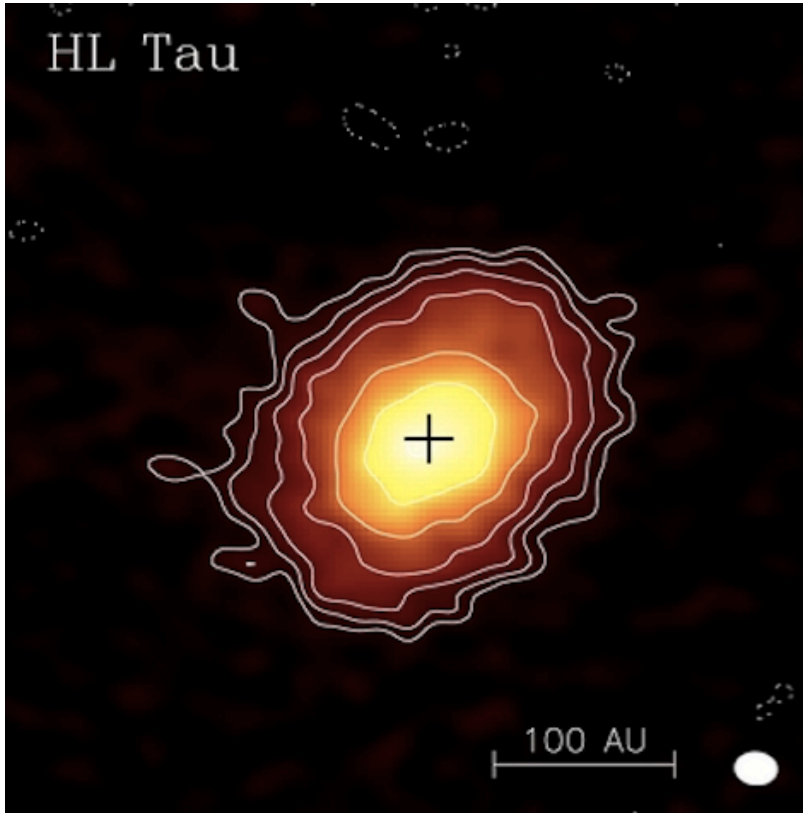
The Atacama Large (sub-)Millimeter Array (ALMA)

The ALMA Re*S*olution

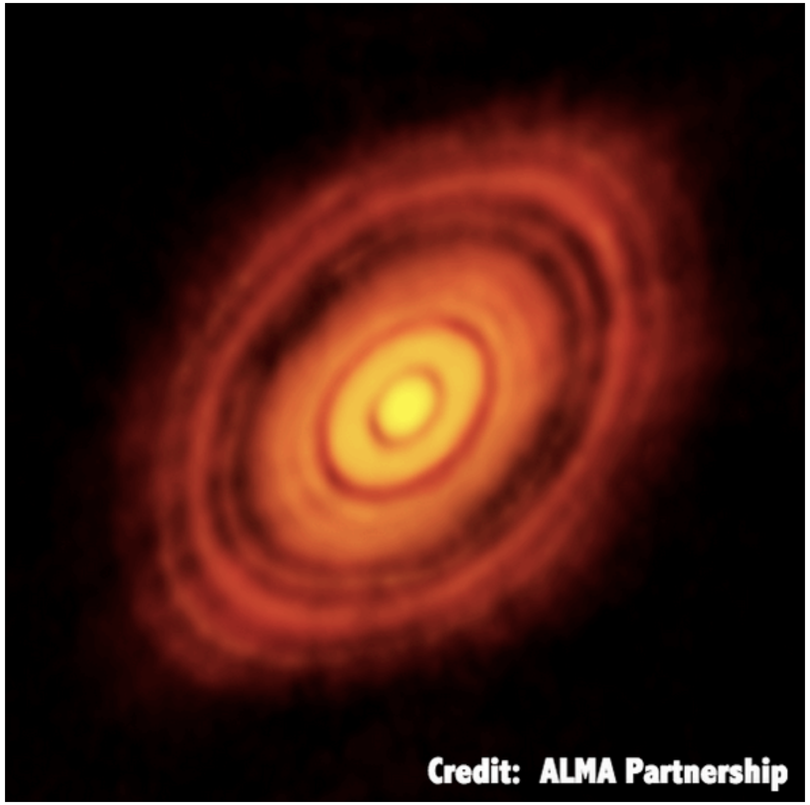




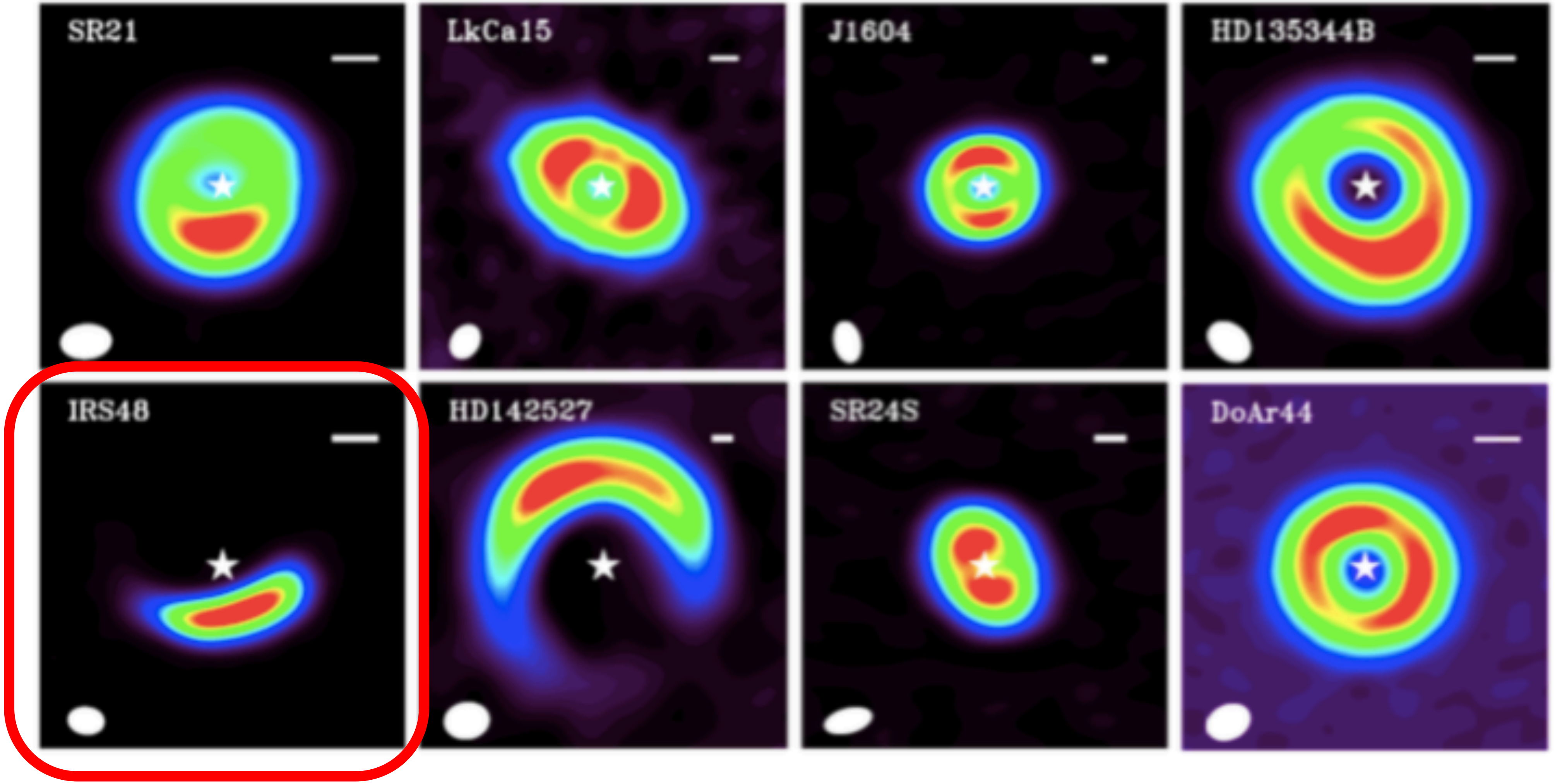
Before ALMA



ALMA



Dust traps in disks: ALMA Cycle 0 (2012)



Oph IRS 48

A Major Asymmetric Dust Trap in a Transition Disk

Nienke van der Marel,[1*] Ewine F. van Dishoeck,[1,2] Simon Bruderer,[2] Til Birnstiel,[3] Paola Pinilla,[4] Cornelis P. Dullemond,[4] Tim A. van Kempen,[1,5] Markus Schmalzl,[1] Joanna M. Brown,[3] Gregory J. Herczeg,[6] Geoffrey S. Mathews,[1] Vincent Geers[7]

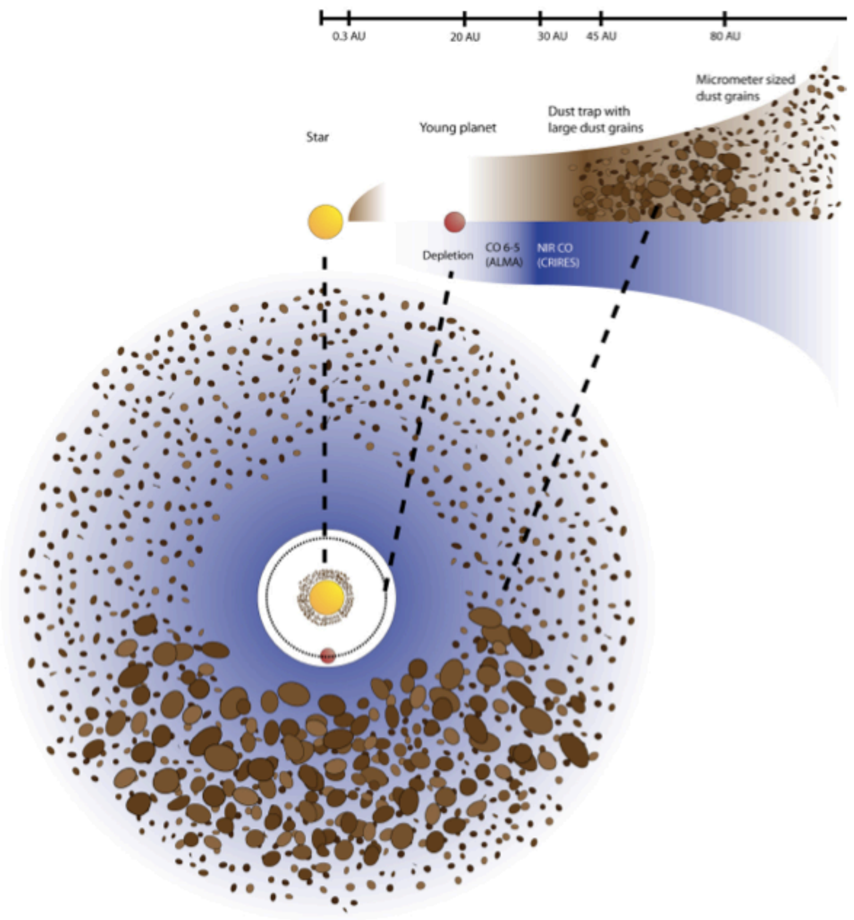
The statistics of discovered exoplanets suggest that planets form efficiently. However, there are fundamental unsolved problems, such as excessive inward drift of particles in protoplanetary disks during planet formation. Recent theories invoke dust traps to overcome this problem. We report the detection of a dust trap in the disk around the star Oph IRS 48 using observations from the Atacama Large Millimeter/submillimeter Array (ALMA). The 0.44-millimeter–wavelength continuum map shows high-contrast crescent-shaped emission on one side of the star, originating from millimeter-sized grains, whereas both the mid-infrared image (micrometer-sized dust) and the gas traced by the carbon monoxide 6-5 rotational line suggest rings centered on the star. The difference in distribution of big grains versus small grains/gas can be modeled with a vortex-shaped dust trap triggered by a companion.

Although the ubiquity of planets is confirmed almost daily by detections of new exoplanets (*1*), the exact formation mechanism of planetary systems in disks of gas and dust around young stars remains a long-standing problem in astrophysics (*2*). In

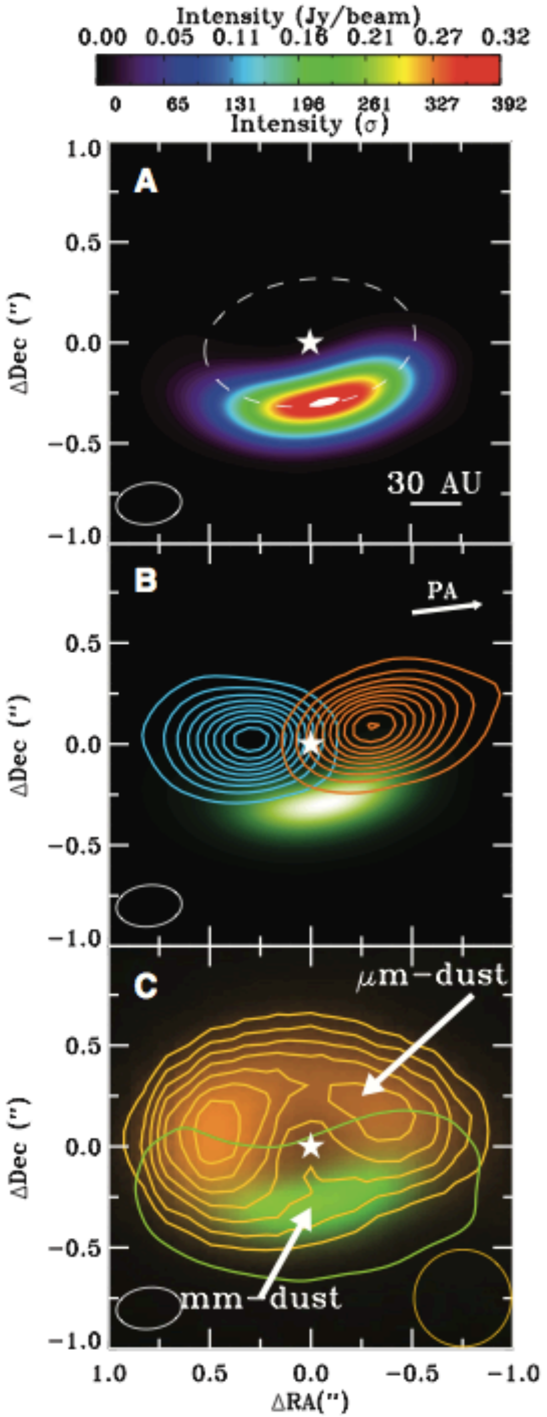
van der Marel+ ‘13

A huge vortex observed with ALMA

The Oph IRS 48 "comet formation factory"



van der Marel+. '13

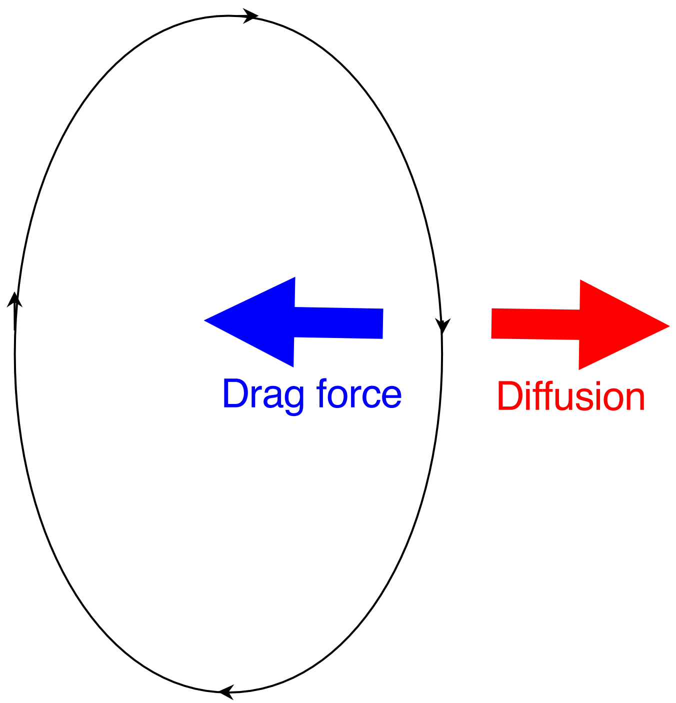


asymmetric
mm dust
at 63 AU

Gas detection:
Keplerian rotation

Micron-sized
dust follows gas

Drag-Diffusion Equilibrium

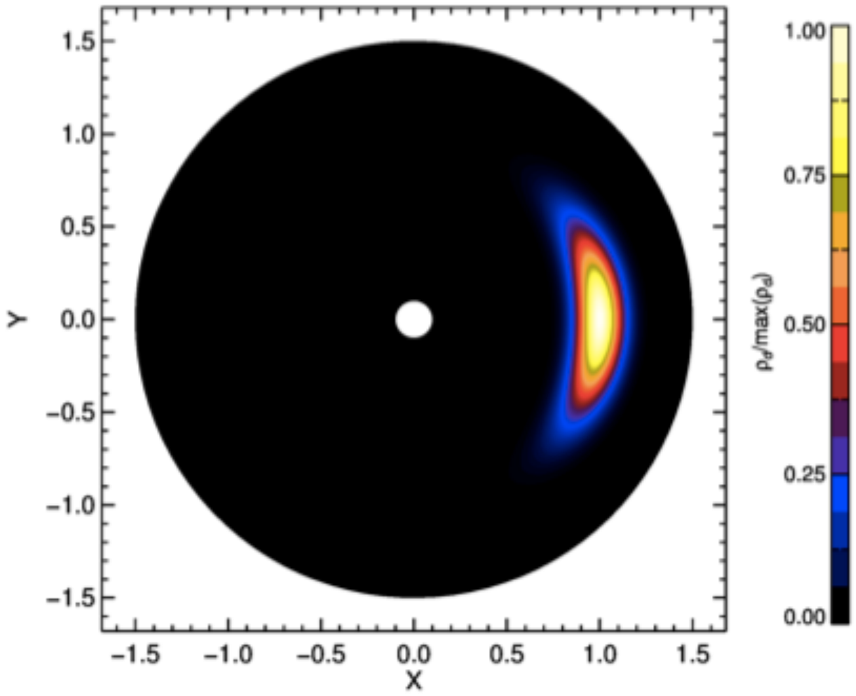


Trapped particle

Dust continuity equation

$$\frac{\partial \rho_d}{\partial t} = \underbrace{-(\boldsymbol{v} \cdot \boldsymbol{\nabla})\rho_d}_{\text{advection}} \underbrace{- \rho_d \boldsymbol{\nabla} \cdot \boldsymbol{v}}_{\text{compression}} + \underbrace{D \nabla^2 \rho_d}_{\text{diffusion}},$$

Analytical Solution for dust in Drag-Diffusion Equilibrium



Solution for

H/r=0.1 χ*=4* *S*=1

Steady-state solution

$$\rho_d(a,z) = \varepsilon\rho_0\,(S+1)^{3/2}\,\exp\left\{-\frac{[a^2 f^2(\chi)+z^2]}{2H^2}(S+1)\right\}$$

Lyra & Lin '13

$$S = \frac{St}{\delta}$$

$$\delta = v_{\rm rms}^2/c_s^2,$$

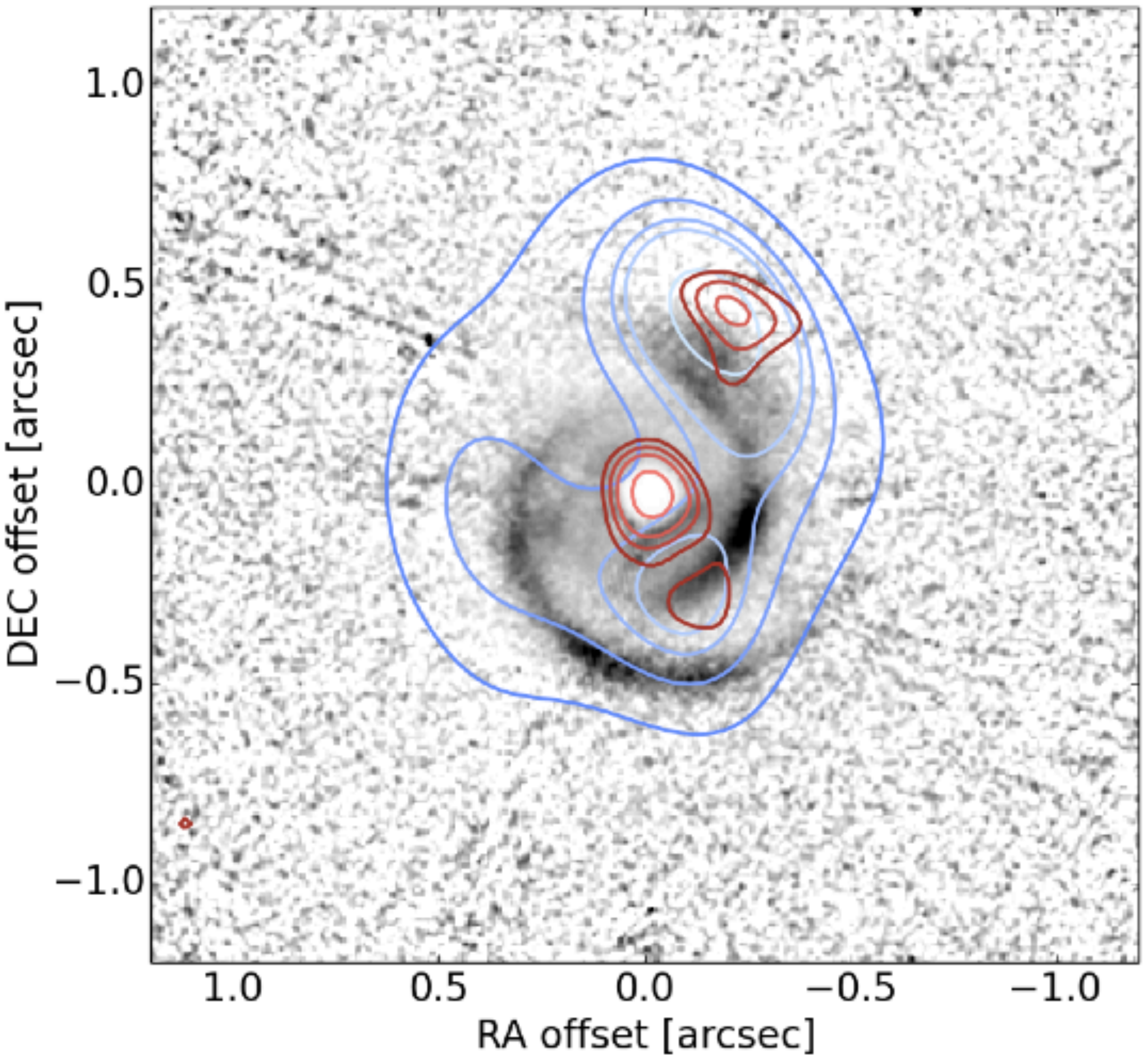
a = vortex semi-minor axis
H = disk scale height (temperature)
χ = vortex aspect ratio
δ = diffusion parameter
St = Stokes number (particle size)
f(χ*)* = model-dependent scale function

Disk Tomography
SPHERE-ALMA-VLA overlay of MWC 758

SPHERE (μm)

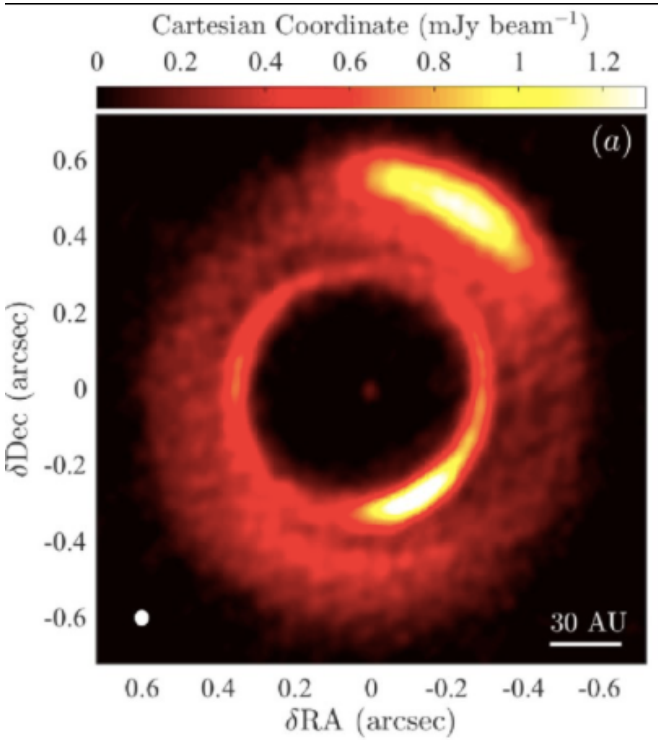
ALMA (~ mm)

VLA (cm-m)



Marino+Lyra '15

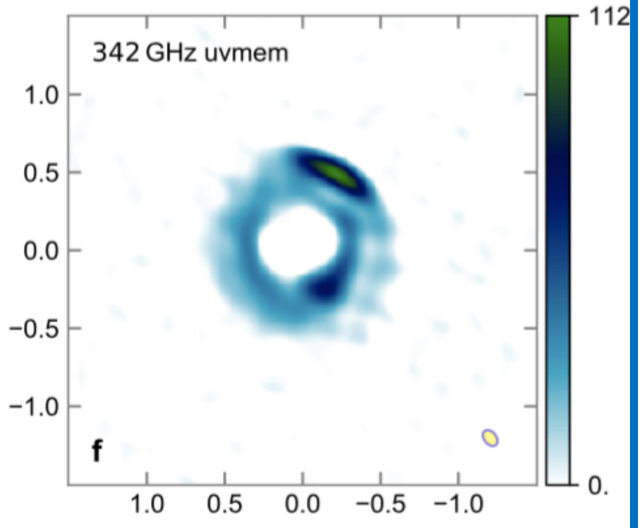
ALMA



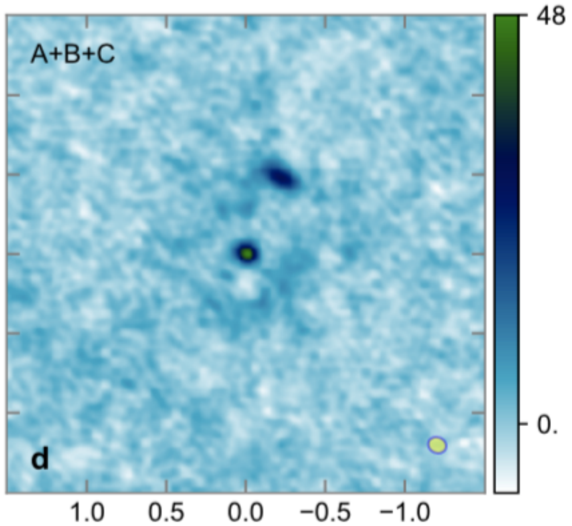
Dong+ '18

Pebble trapping

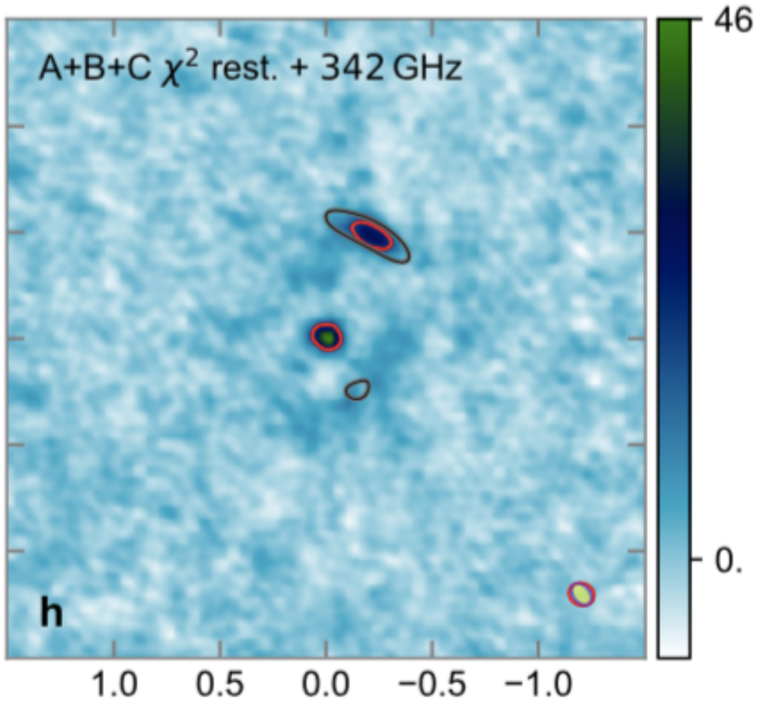
ALMA
(mm)



VLA
(cm-m)

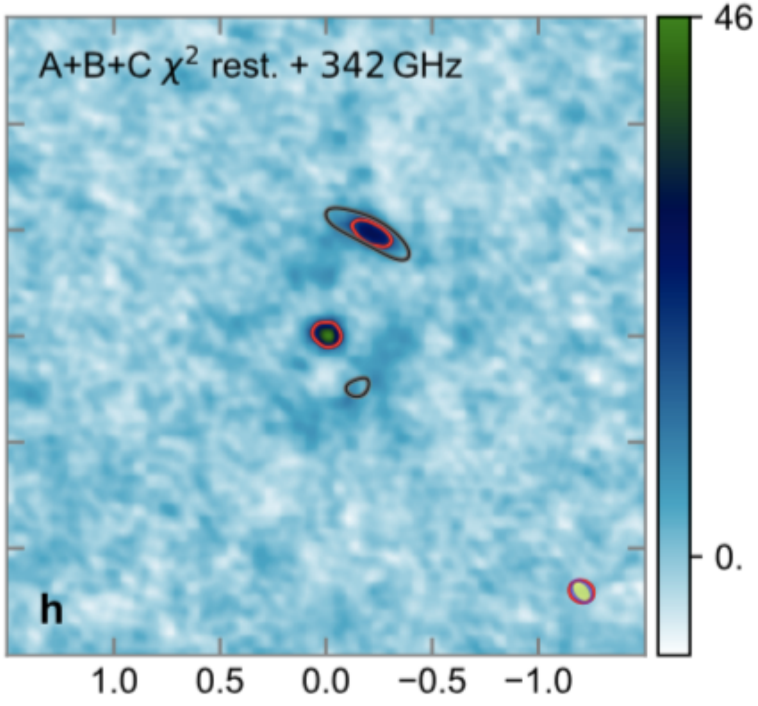


Overlay

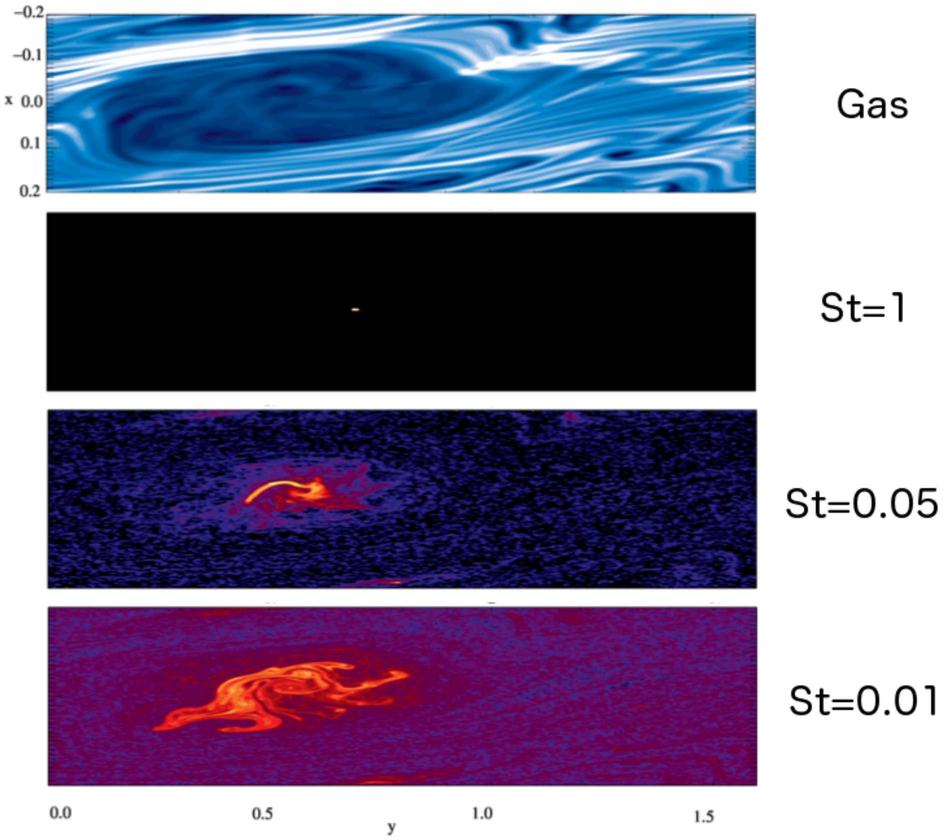


Casassus+Lyra '19

Model vs Observation



Casassus+Lyra '19



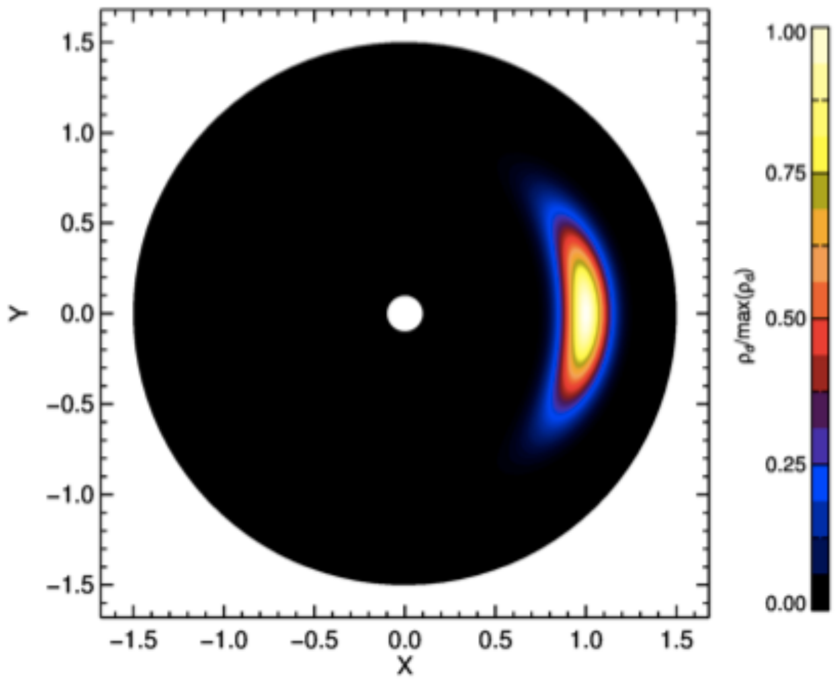
Raettig+Lyra '15

Take home message

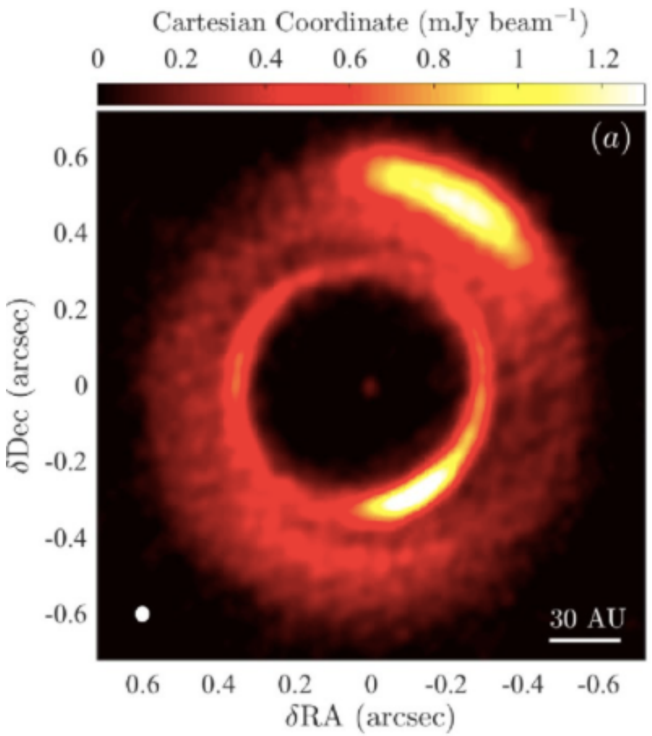
- Vortex-trapped dust in drag-diffusion equilibrium explains the observations

$$\rho_d(a,z) = \varepsilon \rho_0 (S+1)^{3/2} \exp\left\{-\frac{[a^2 f^2(\chi) + z^2]}{2H^2}(S+1)\right\}$$

Lyra-Lin solution



Observed Disk



The future

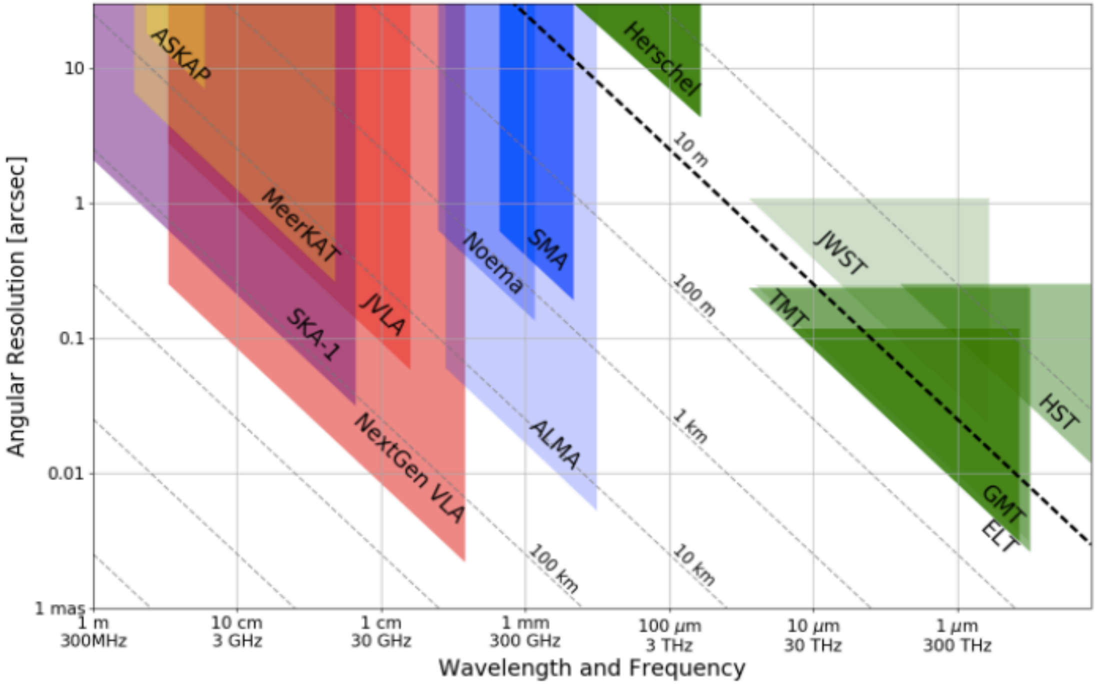
After 8 years of ALMA…

Nearly all nearby disks observed at <0.1” (< 20-30AU) show substructures.

3 main types of substructures

- Crescent-shaped
- Spiral arms
- Rings/Gaps

Next Generation Very Large Array (ngVLA)



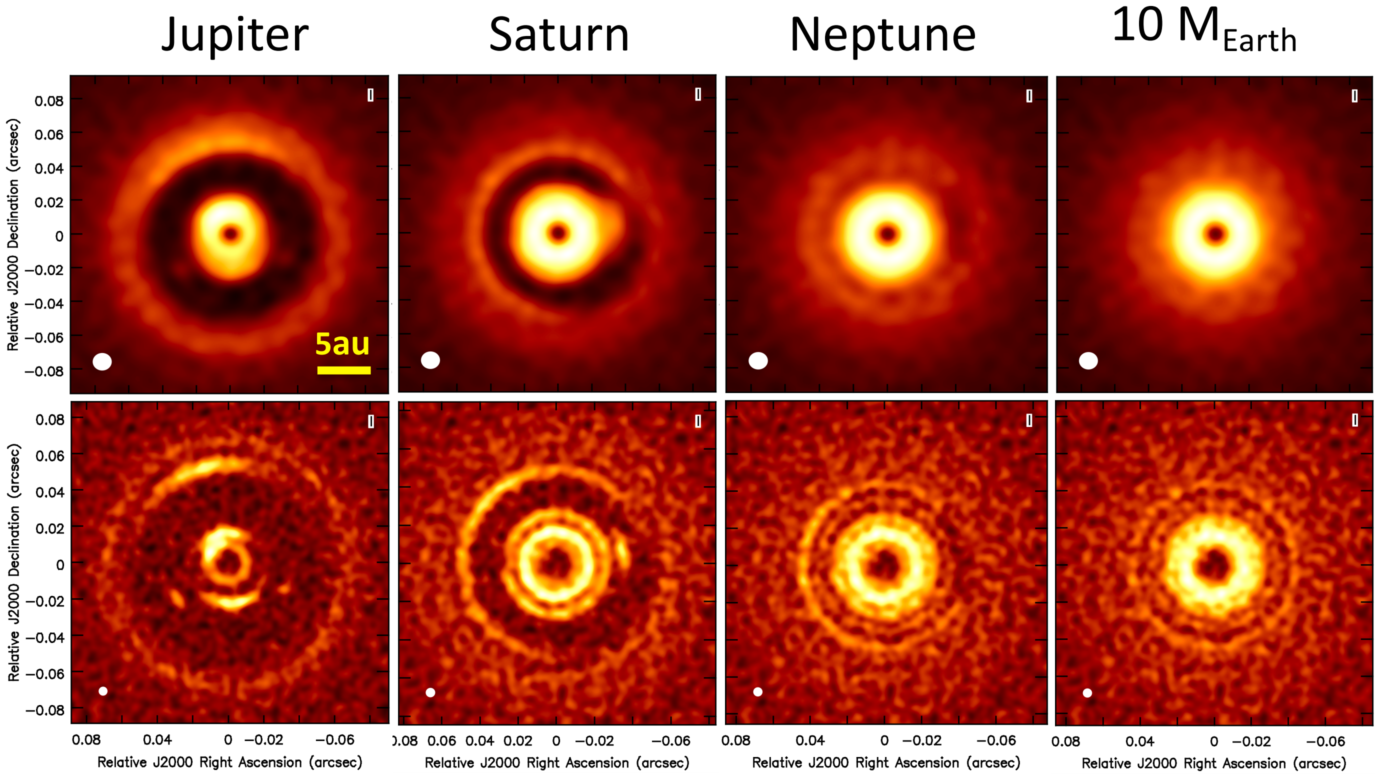
Planets at 5AU

ALMA @ 0.87mm

ngVLA @ 3mm

5 mas = 0.7 AU

rms = 5×10^{-7} Jy/beam

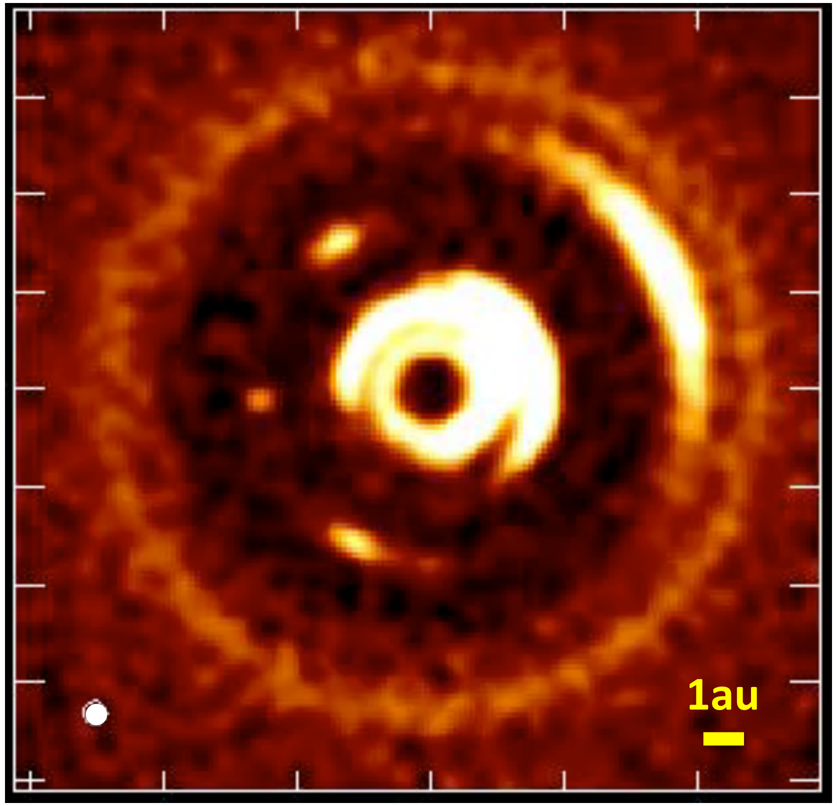


Ricci et al. 2018

ngVLA identifies gaps/substructures down to ~5-10 M_{Earth}

ngVLA: Proper motions

Jupiter at 5 AU



Conclusions

- Two routes for planet formation (streaming instability and vortices, complementary)
- Does turbulence help (concentration at large scales) or hinder (diffusion at small scales)?

- Three dynamical instabilities in the Ohmic dead zone

    - Different regimes of opacity, operate in different regions
    - Saturate into vortices

  - Dust trapped in drag-diffusion equilibrium explains the observations

- **Issues**:
  - Are the dynamical instabilities responsible for the observed crescents?
  - Overlap unclear
  - Global model of COV needed
  - Relevance of ZVI unclear/unlikely.
  - Planet formation properties / Synergy with streaming instability

	ZVI	COV	VSI
Global model	✗	✗	✓
Vertical Stratification	✓	✗	✓
Boundaries with other instabilities	✗	✗	✗
Interaction with dust	✗	✓	✓
Observational Validation/Rule out	✗	✗	✗
Planet Forming Properties	✗	✗	✗