

Instructions

Welcome to Cume #404. This exam will address subjects related to extragalactic astronomy and cosmology and makes frequent reference to Atek+2015 (arXiv:1509.06764v2). Do not read the paper in detail; that would take too long. I recommend just reading the abstract, the figure captions, and the conclusions, and then referring back to the paper as the questions indicate.

There are 36 possible points total. The anticipated passing grade is 25.

1 Concepts

1. (1 point) What causes the Lyman break feature of high-redshift galaxy spectra?
It is caused by absorption from neutral hydrogen in the IGM at a rest-frame wavelength of $1216\text{\AA}$. As light from the galaxy travels, successively bluer wavelengths redshift into resonance with the IGM and are scattered, leading to an observed “break”.
Rubric: 1 point for mentioning the IGM and/or neutral hydrogen.
2. In this problem, you will compute the redshift at which galaxies begin to “drop out” of the F814W photometric band in Figure 1.
 - (a) (1 point) What is the rest-frame wavelength λ_{LB} of the Lyman break spectral feature?
 $1216\text{\AA}$
Rubric: 1 point for either $1216\text{\AA}$ or $n = 1 \rightarrow 2$.
No points off if they answer that it's $912\text{\AA}$ as the concept is still understood.
 - (b) (2 points) Beyond what redshift does this feature lie to the red of the F814W photometric band? If you don't know what λ_{LB} is, just leave it as a symbol.
The red end of the F814W band is at $9000\text{\AA}$ in the observed frame.

$$1216 \times (1 + z) = 9000$$

$$z \approx 6.4$$

Rubric: 1 point for identifying the filter's red end correctly; 1 point for setting up the equation to compute z . No penalty if all this is correct but z is in error.

3. The paper notes a strong connection between a galaxy's star formation rate and its UV luminosity. Let's talk about that.
 - (a) (1 point) Why is a galaxy's UV luminosity a good tracer of its star formation rate?
UV is dominated by short-lived massive stars. As long as the star formation history is smooth on 100 Myr timescales, the UV flux traces the number of such stars, which is hence proportional to the SFR.
We can think of this slightly more formally as follows: Suppose a fraction f_M of all stellar mass goes into massive stars that live for τ years. Then if a galaxy forms stars steadily for a long time compared to τ , the creation rate of massive stars is

$\text{SFR} \times f_M$ and the death rate is M_M/τ . When these are equal, $M_M = \text{SFR} \times f_M/\tau$ is the mass in massive stars, which is proportional to the UV luminosity.

Rubric: 1 point for mentioning massive stars or their short main-sequence lifespans.

- (b) (2 points) The major uncertainties in converting from measured UV luminosity to star formation rate are believed to be the amount of dust in the galaxy's interstellar medium, and the initial mass function. Explain why each of these is a source of uncertainty.

Dust suppresses observable UV flux at given SFR, making galaxies appear to have lower SFR than they actually do. The initial mass function relates the star formation in massive stars (which are visible in the UV) to the total star formation rate (including low-mass stars that don't radiate in the UV). We don't know what the IMF was at $z = 6$. It could have had relatively more or fewer low-mass stars per massive star than local measurements, leading to uncertainty in the relationship between UV flux and total star formation rate.

Rubric: 1 point for dust extinction on the UV; 1 point for the distinction between low-mass and massive stars within the IMF.

4. Figure 3 compares a histogram of the number of galaxies identified as a function of redshift versus completeness. Let's interpret it.

- (a) (1 point) Completeness is a ratio of two things. What are they?

Number of galaxies identified divided by total number of galaxies in the field.

Rubric: 1 point for mentioning both of these.

- (b) (2 points) Assume that the completeness calculations are correct and that the intrinsic galaxy abundance does not vary from $z=5$ to $z=9$ (it does, but never mind that). Why would there still be fewer galaxies detected at $z = 8$ than at $z = 6$?

If there were intrinsically just as many galaxies at $z = 8$ as at $z = 6$ one would still *detect* fewer of them because they are farther away and fainter (cosmological surface brightness dimming $\propto (1+z)^{-4}$).

Rubric: 1 point for "fainter" or "farther", 2 points for mentioning cosmological surface brightness dimming.

5. The goal of this paper is to measure the luminosity function ϕ of galaxies. ϕ is defined as the number of galaxies dN per absolute magnitude interval dM per comoving volume dV , $\phi = dN/dM/dV$. Counting the number of galaxies dN in a magnitude interval dM is conceptually straightforward, but the volume dV is more difficult. Generally, people compute an "effective volume" V_{eff} that varies with luminosity. Figure 5 shows the survey's estimated V_{eff} .

- (a) (2 points) Why does V_{eff} vary with luminosity? Point to a paragraph or equation in the paper that supports your explanation.

The completeness declines with decreasing luminosity because fainter objects are

more difficult to see at a given distance or redshift. As shown by Equation 3, $V_{\text{eff}} \propto \text{completeness } f$, so if completeness goes down then so does $V_{\text{eff}} \propto$.

Rubric: 1 point for a sensible explanation, 1 for pointing to Equation or another relevant part of the paper.

- (b) (1 point) Suppose that the intrinsic number of galaxies per magnitude interval dN/dM were a constant; i.e., that there are just as many bright galaxies as faint ones. Based on Figure 5, would you expect the survey to uncover more bright galaxies or faint ones? Please explain.

More bright galaxies. The effective volume of the survey is larger for brighter galaxies; more volume = more galaxies.

- (c) (1 point) In reality, the Universe makes more faint galaxies than bright ones. According to Equation 5, the luminosity function of *faint* galaxies is proportional to $10^{0.4(\alpha+1)M}$, where α is the faint-end slope and M is the absolute magnitude. What is the condition on survey effective volume versus absolute magnitude $V_{\text{eff}}(M)$ such that the number of faint galaxies detected per magnitude interval by the survey is *constant*; that is, it does not vary with M ?

There was a typo in the exam here; $\phi \propto 10^{-0.4(\alpha+1)M}$ for faint galaxies (that is, a negative sign in the exponent was missing).

The condition that the number of identified galaxies be constant with absolute magnitude is

$$\begin{aligned} \frac{dN}{dM dV}(M) \times V_{\text{eff}}(M) &\sim \text{constant} \\ &= 10^{-0.4(\alpha+1)M} V \end{aligned}$$

If $V_{\text{eff}}(M) \propto 10^{0.4(\alpha+1)M}$, then this is satisfied. In practice, magnification doesn't help *this* much, but it still helps.

Rubric: 1 point for either verbally or formally stating this condition. (No points off if they were thrown off by the typo!)

- (d) (1 point) The left column of Figure 5 shows effective volume versus absolute magnitude in survey fields that have a foreground galaxy cluster, while the right column shows "parallel" fields. Why are the effective volumes different in the presence of foreground clusters?

Magnification by foreground clusters increases the surface area corresponding to background regions. The magnified regions tend not to be overdense, so magnification increases effective volume more for faint galaxies than for bright ones. In practice, the field sizes are also different, hence the parallel fields may have both a larger overall effective volume and a different trend of $V_{\text{eff}}(M)$.

Rubric: 1 point for any of the above.

6. Section 5.2 discusses the sources of uncertainty affecting the derived luminosity function.

- (a) (1 point) Why is cosmic variance a larger problem for bright galaxies than for faint ones?

Bright galaxies are rare and clustered, hence a larger volume must be surveyed in order to identify a representative sample.

Rubric: 1 point for any of the above.

- (b) (2 points) Why is incompleteness a larger concern for faint galaxies than for bright ones? Point to a figure that supports your explanation.

Faint galaxies are more subject to photometric errors scattering them out of the sample, or confusion with foreground sources, hence they have higher incompleteness and lower V_{eff} (Figure 5).

Rubric: 1 point mentioning errors or confusion; 1 point for referring to Figure 5 or something similarly relevant.

2 The Recombination Rate at $z = 7$

In this section, you will derive the right-hand side of Equation 6. The Friedmann equation relates the Hubble constant, H_0 , to the total proper energy density ϵ . For a spatially-flat Universe,

$$H_0^2 = \frac{8\pi G \epsilon_c}{3c^2}$$

where ϵ_c is the *critical energy density* as measured today. We can also interpret this energy density as a critical mass density $\rho_{\text{crit}} \equiv \epsilon_{\text{crit}}/c^2$:

$$H_0^2 = \frac{8\pi G \rho_c}{3}$$

The baryon density is just $\Omega_b \rho_{\text{crit}}$.

- (3 points) Using the same values $H_0 = 100h$ and $\Omega_b h^2$ as the paper, compute the current (today) mean hydrogen number density in cgs units. You will need to assume that the cosmic hydrogen mass fraction $X_H = 0.76$, that the proton mass is 1.67×10^{-24} t, and $G = 6.67 \times 10^{-8} \text{ cm}^3 \text{ g}^{-1} \text{ sec}^{-2}$.

$$\begin{aligned} n_H &= \Omega_b \rho_c X_H / M_H \\ &= \Omega_b \frac{X_H}{M_H} \frac{3H_0^2}{8\pi G} \\ &= (\Omega_b h^2) \frac{0.76}{1.67 \times 10^{-24}} \frac{3 \times 100^2}{8\pi G} (\text{km sec}^{-1} \text{Mpc}^{-1})^2 \\ &= 1.7 \times 10^{-7} \left(\frac{\Omega_b h^2}{0.02} \right) \text{cm}^{-3} \end{aligned}$$

Rubric: 1 point for setting it up correctly; 1 for finding $\Omega_b h^2 = 0.02$ in the paper; 1 for solving correctly.

2. (2 points) Given this mean hydrogen number density at $z = 0$, what is the mean hydrogen number density at $z = 7$, $\langle n_H \rangle(z = 7)$ (If you couldn't complete the previous question, just use $n_H = 10^{-7} \text{ cm}^{-3}$.)

$$\begin{aligned} n_H(z = 7) &= n_H(z = 0)(1 + z)^3 \\ &= 8.8 \times 10^{-5} \left(\frac{\Omega_b h^2}{0.02} \right) \left(\frac{1 + z}{8} \right)^3 \text{ cm}^{-3} \end{aligned}$$

If the student adopted 10^{-7} for $n_H(z = 0)$, then the answer is $5 \times 10^{-5} \text{ cm}^{-3}$.

Rubric: 1 point for setting it up correctly; 1 for getting the correct answer.

3. (2 points) The mean recombination rate per unit (proper) volume in a fully-ionized hydrogen plasma with density inhomogeneities is:

$$R = \frac{1}{V} \int_V \alpha_R(T) n_H^2 dV$$

where α_R is the recombination coefficient. Show that this total recombination rate can be written as

$$R = \alpha_R(T) C \langle n_H \rangle^2$$

where the *clumping factor* C is

$$C \equiv \langle n_H^2 \rangle / \langle n_H \rangle^2$$

and $\langle \rangle$ indicate volume-averages.

$$\begin{aligned} \frac{R}{V} &= \frac{1}{V} \int_V \alpha n_H^2 dV \\ &= \alpha \langle n_H^2 \rangle \\ &= \alpha \frac{\langle n_H^2 \rangle}{\langle n_H \rangle^2} \langle n_H \rangle^2 \\ &= \alpha C \langle n_H \rangle^2 \end{aligned}$$

Rubric: 1 point for clear thinking in the setup or 2 points for getting it all

4. (2 points) Assume that $\alpha_R = 2.59 \times 10^{-13} \text{ sec}^{-1} \text{ cm}^{-3}$, and leaving C as a variable, compute the recombination rate of a fully-ionized IGM at $z = 7$ in units of recombinations per time per proper Mpc^3 .

Plugging in numbers, I get:

$$\begin{aligned} \frac{R}{V} &= 2 \times 10^{-21} C \left(\frac{\Omega_b h^2}{0.02} \right)^2 \left(\frac{1 + z}{8} \right)^6 (\text{cm}^{-3} \text{ sec}^{-1}) \\ &= 5.9 \times 10^{52} C \left(\frac{\Omega_b h^2}{0.02} \right)^2 \left(\frac{1 + z}{8} \right)^6 \text{ pMpc}^{-3} \text{ sec}^{-1} \end{aligned}$$

Rubric: 1 point for setting it up correctly; 1 point for solving correctly

5. (2 points) What is this in recombinations per time per *comoving* Mpc^3 ? (Check: If you set $C = 30$, $z = 5$, you should recover something close to the the right-hand side of Equation 6 in the paper.)

One comoving Mpc^3 corresponds to $(1+z)^{-3}$ proper Mpc^3 at redshift z , so we divide by 8^3 to get

$$\frac{R}{V} = 1.2 \times 10^{50} C \left(\frac{\Omega_b h^2}{0.02} \right)^2 \left(\frac{1+z}{8} \right)^3 (\text{cMpc}^{-3} \text{sec}^{-1})$$

Rubric: 1 point for setting it up correctly; 1 point for solving correctly

3 The Ionization Rate at $z = 7$

The ionizing photon production rate per unit comoving volume from galaxies is assumed in this paper (and elsewhere) to be

$$\dot{N}_{\text{ion}} = L_{\text{UV}} \xi f_{\text{esc}}$$

ξ and f_{esc} are parameters that we will ignore. In this problem, you will step through how to compute the luminosity density L_{UV} .

1. (1 point) Set up an integral which could be used to compute L_{UV} using the Schechter function in Equation 5. Reminders:

- (a) A galaxy's specific luminosity L_ν relates to its absolute AB magnitude M as:

$$M = -2.5 \log \left(\frac{L_\nu}{(10 \text{pc})^2} \right) - 48.6$$

- (b) The luminosity density comes from the integral:

$$L_{\text{UV}} = \int \frac{dN}{dM dV} L_\nu(M) dM$$

Setting up the integral:

$$\begin{aligned} L_{\text{UV}} &= \int_{M^*}^{M_{\text{min}}} \frac{dN}{dM dV} L_\nu(M) dM \\ &= \int_{M^*}^{M_{\text{min}}} \frac{\ln(10)}{2.5} \phi^* 10^{0.4(\alpha+1)(M^*-M)} \exp(-10^{0.4(M^*-M)}) 10^{[(M+48.6)/-2.5]} (10 \text{pc})^2 dM \end{aligned}$$

Rubric: 1 point for getting most of the above.

2. (2 points) It would take you a while to evaluate that integral, so don't. However, considering only the galaxies that are fainter than the characteristic magnitude M_* , show that the L_{UV} is proportional to

$$L_{UV} \propto \int_{M_*}^{M_{\min}} 10^{-0.4(\alpha+2)M} dM$$

where M_* is the characteristic magnitude and $M_{\min}$ is the faintest galaxy under consideration.

If we drop all of the constant terms and ignore the exponential term, then what is left is:

$$\begin{aligned} L_{UV} &\propto \int_{M_*}^{M_{\min}} 10^{-0.4(\alpha+1)M} 10^{-0.4M} dM \\ &= \int_{M_*}^{M_{\min}} 10^{-0.4(\alpha+2)M} dM \end{aligned}$$

Rubric: 1 point for a solid effort, 1 point for getting it right. No points for just copying the answer.

3. (2 points) Show that, for $\alpha = -2$, the integral grows arbitrarily large as $M_{\min}$ is taken to be fainter. (Remember that, in the magnitude system, highly negative M corresponds to a very bright source; that is, -20 is *brighter* than -15.25).

For $\alpha = -2$, all of the M dependence in the integral cancels, and what is left is that

$$L_{UV} \propto (M_{\min} - M_*)$$

This diverges as $M_{\min}$ becomes arbitrarily positive (corresponding to arbitrarily faint galaxies).

Rubric: 1 point for serious effort related to manipulating the previous expression for L ; 1 point for showing the correct answer.

4. (2 points) The current paper traces the luminosity function two magnitudes fainter than blank-field measurements. Meanwhile, two Fridays ago, we heard Steve Finkelstein announce that his group has traced the luminosity function to even fainter magnitudes. Given the answer to the previous question, why is this necessary?

The point of the study is to ask whether galaxies can provide enough ionizing photons to maintain an ionized IGM. The production rate of ionizing photons per volume per time is believed to be proportional to L_{UV} , which in turn is quite sensitive to the number of faint galaxies. Hence it is important to know how faint the UV luminosity function goes, and the only way to do that is to keep probing to fainter luminosities.

Rubric: 1 point for mentioning that L_{UV} grows if deeper observations determine that $M_{\min}$ is fainter; 1 point for mentioning something about ionization balance.

ARE ULTRA-FAINT GALAXIES AT $Z = 6 - 8$ RESPONSIBLE FOR COSMIC REIONIZATION ? COMBINED CONSTRAINTS FROM THE HUBBLE FRONTIER FIELDS CLUSTERS AND PARALLELS †

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ABSTRACT

We use deep *Hubble Space Telescope* (*HST*) imaging of the Frontier Fields (FF) to accurately measure the galaxy rest-frame ultraviolet luminosity function (UV LF) in the redshift range $z \sim 6 - 8$. We combine observations in three lensing clusters A2744, MACS0416, MACS0717 and their associated parallels fields to select high-redshift dropout candidates. We use the latest lensing models to estimate the flux magnification and the effective survey volume in combination with completeness simulations performed in the source plane. We report the detection of 227 galaxy candidates at $z = 6 - 7$ and 25 candidates at $z \sim 8$. While the total survey area is about 4 arcmin² in each parallel field, it drops to about 0.6 to 1 arcmin² in the cluster core fields because of the strong lensing. We compute the UV luminosity function at $z \sim 7$ using the combined galaxy sample and perform Monte Carlo simulations to determine the best fit Schechter parameters. We are able to reliably constrain the LF down to an absolute magnitude of $M_{UV} = -15.25$, which corresponds to $0.005L^*$. More importantly, we find that the faint-end slope remains steep down to this magnitude limit with $\alpha = -2.04^{+0.13}_{-0.17}$. We find a characteristic magnitude of $M^* = -20.89^{+0.60}_{-0.72}$ and $\text{Log}(\phi^*) = -3.54^{+0.48}_{-0.45}$. Our results confirm the most recent results in deep blank fields (Finkelstein et al. 2014; Bouwens et al. 2015b) but extend the LF measurements more than two magnitudes deeper. The UV LF at $z \sim 8$ is not very well constrained below $M_{UV} = -18$ due to the small number statistics and incompleteness uncertainties. To assess the contribution of galaxies to cosmic reionization we derive the UV luminosity density at $z \sim 7$ by integrating the UV LF down to an observational limit of $M_{UV} = -15$. We show that our determination of $\text{Log}(\rho_{UV}) = 26.2 \pm 0.13$ (erg s⁻¹ Hz⁻¹ Mpc⁻³) can be sufficient to maintain reionization with an escape fraction of ionizing radiation of $f_{esc} = 10 - 15\%$. Future HFF observations will certainly improve the constraints on the UV LF at the epoch of reionization, paving the way to more ambitious programs using cosmic telescopes with the next generation of large aperture telescopes such as the *James Webb Space Telescope* and the *European Extremely Large Telescope*.

Subject headings: galaxies: evolution — galaxies: high-redshift — galaxies: luminosity function — gravitational lensing: strong

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1. INTRODUCTION

One of the most important challenges in observational cosmology is the identification of the sources responsible for cosmic reionization. Shortly after the Big Bang the Universe was completely neutral following the recombination of hydrogen atoms, until the first sources started to reionize the neutral gas in their surroundings. Several observational results have now narrowed down the period of reionization to the redshift interval $6 < z < 12$. Observations of the Gunn-Peterson effect in the absorption spectra of quasars and gamma-ray bursts (GRB) indicate that the Universe was mostly ionized by $z \sim 6$ (Fan et al. 2006; Chornock et al. 2014). The sensitivity of Ly α emission to neutral gas is also used to probe the ionization state of the intergalactic medium (IGM) at $z > 6$. In particular, the prevalence of Ly α emitters (LAEs) among continuum selected galaxies (LBGs) appears to drop very rapidly at $z > 6.5$, which suggests an increase in the fraction of neutral hydrogen (Stark et al. 2010; Treu et al. 2013; Schenker et al. 2014a; Pentericci et al. 2014). The decline in the LAEs fraction also favors a patchy rather than a smooth reionization process. The optical depth of Thomson scattering to the cosmic microwave background (CMB) recently reported by the Planck collaboration (Planck Collaboration et al. 2015) implies a redshift of instantaneous reionization around $z_r = 8.8^{+1.3}_{-1.2}$, significantly lower than earlier determinations of $z_r = 10.6 \pm 1.1$ by the Wilkinson Microwave Anisotropy Probe (WMAP, Bennett et al. 2013). Despite these major advances, large uncertainties remain regarding the main sources that drive the reionization process.

Early star-forming galaxies are now thought to be the best candidates for providing the required ionizing power (e.g. Robertson et al. 2014; Mitchell-Wynne et al. 2015; Duncan & Conselice 2015; Finkelstein et al. 2014; Mason et al. 2015). Deep imaging campaigns of blank fields with the *Hubble Space Telescope* and ground-based instrumentation have made important inroads in constraining the galaxy ultraviolet luminosity function (UV LF) out to $z \sim 10$ (e.g., Bouwens et al. 2012; Bunker et al. 2010; Oesch et al. 2012; McLure et al. 2013; Schmidt et al. 2014; Bouwens et al. 2015b; Finkelstein et al. 2014), which, in turn, encodes important information about the cosmic star formation (Bouwens et al. 2015b; Robertson et al. 2015). Indeed, the rest-frame UV radiation traces the recent star formation averaged over the past hundreds Myr. Therefore, the integration of the UV LF provides the luminosity density of galaxies at a given redshift, which can be translated to the ionizing radiation, assuming a certain star-formation history. The UV luminosity density is sensitive to two main parameters: (i) the faint-end slope of the LF, and (ii) the integration limit at the faint end.

The faint-end slope of the LF evolves with redshift, and gets as steep as $\alpha \sim -2$ at $z \sim 7$. For comparison, the faint-end slope at $z < 0.5$ slope is very shallow ($\alpha \sim -1.3$), which indicates a decreasing contribution of faint star-forming galaxies to the total star formation density towards lower redshift (Arnouts et al. 2005; Schiminovich et al. 2005). The predicted evolution of the dark matter halo mass function based on cosmological simulations predicts an even steeper slope of $\alpha \sim 2.3$ of

the UV LF at high redshift (Jaacks et al. 2012). The deepest *HST* observations of blank fields, such as the XDF, put constraints on the faint-end slope down to an absolute magnitude of $M_{UV} \sim -17.5$ AB. However, the ability of galaxies to reionize the Universe relies on the extent of the steep faint-end slope down to lower luminosities, typically around $0.001 L^*$ at $z \sim 7^{12}$. While cosmological simulations point to a halo mass limit of $10^6 M_\odot$ for early galaxy formation (Jaacks et al. 2012; Kimm & Cen 2014; Wise et al. 2014), the depth of current observations however limits the exploration of the LF to galaxies brighter than $\sim 0.1 L^*$. One particularly efficient way to push the limits of current facilities is to take advantage of the gravitational lensing offered by massive galaxy clusters, which act as natural telescopes by boosting the flux of background sources (Kneib & Natarajan 2011).

Since the discovery of the first giant arcs created by strong lensing (Soucail et al. 1987), cosmic lenses have been successfully used to detect intrinsically faint background sources and perform spatially detailed analysis of distant galaxies. Recently, using *HST* imaging of 25 X-ray-selected clusters in the Abell (Abell et al. 1989) and MACS (Ebeling et al. 2001, 2007, 2010; Mann & Ebeling 2012) catalogs, the CLASH program (Postman et al. 2012) has made important progress in the characterization of the lensing properties of clusters. Multi-wavelength observations has enabled the measurement of the total cluster mass with a precision of 10% and the detection of some of the most distant galaxies at $z > 7$ (e.g. Zitrin et al. 2011; Bradley et al. 2014; Coe et al. 2013). However, two limitations prevented such program from exploring the faint-end of the LF at the epoch of reionization. First, the relatively shallow data compared to HUDF, restricted the survey to the brighter end part of the LF, even in the case of high magnifications. In this sense, wide area surveys, such as BoRG (Trenti et al. 2011), are indeed well suited to help constrain the brighter part of the LF (Bradley et al. 2012; Schmidt et al. 2014). Second, the lack of very accurate lensing models for some of these clusters, mostly due to the low number of multiple images available to constrain the mass distribution, also thwarts the construction of a reliable LF.

The Hubble Frontier Fields (HFF) project aims at overcoming these two major limitations by obtaining deep multi-wavelength observations of six massive galaxy clusters that act as cosmic lenses. The HFF include the deepest optical and near-infrared observations of lensing clusters using *HST* director discretionary time, which are complemented by a wealth of data including *ALMA*, *Spitzer*, *Chandra*, *XMM*, *VLA*, as well as *HST* and ground-based imaging and spectroscopic follow-up (e.g. Owers et al. 2011; Ebeling et al. 2014; Rawle et al. 2015; Karman et al. 2015; Richard et al. 2015; Medezinski et al. 2015; Rodney et al. 2015; Zitrin et al. 2015; Schirmer et al. 2015; Grillo et al. 2015; Treu et al. 2015; Ogorean et al. 2015). Based on the full HFF dataset of the first cluster A2744, Atek et al. (2015) presented the first constraints on the UV LF at $z \sim 7$ and $z \sim 8$ (see also Ishigaki et al.

¹² The characteristic magnitude $M_{UV}^* = -21$ AB (e.g. Bouwens et al. 2015b; Atek et al. 2014b)

2015a). The key result was the steep faint-end slope of $\alpha \sim -2.01$ that extends down to $M_{UV} = -15.5$ AB. However, large uncertainties arising from small sample size, lensing models, and cosmic variance still prevent strong conclusions on the total UV luminosity density of galaxies at the epoch of reionization.

In this paper, we combine the complete dataset of the three lensing clusters A2744, MACS0416, and MACS0717, and their respective parallel fields to search for high-redshift dropout galaxies and put stronger constraints on the UV LF at $z > 6$. The paper is organized as follows. In Section 2, we describe the observations. The sample selection method is described in Section 3. In Section 4, we briefly describe the lensing models and the multiple-image identification. The procedure and the results of the computation of the UV luminosity function are presented in Section 5. A summary is given in Section 6. Throughout the paper, we adopt a standard Λ CDM cosmology with $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_\Lambda = 0.73$, and $\Omega_m = 0.27$ to be consistent with previous studies. All magnitudes are expressed in the AB system.

2. HFF OBSERVATIONS

The HFF clusters and parallel fields were observed by *HST* with three ACS (Advanced Camera for Survey) optical filters (F435W, F606W, F814W) and four WFC3 (Wide Field Camera 3) near-IR filters (F105W, F125W, F140W, F160W). Observations were scheduled in two-epoch sequences, obtaining ACS observations of the main cluster and WFC3 observations of the parallel field in one epoch and swapping instruments in the second epoch. We use the high-level science products delivered by the Space Telescope Science Institute (STScI) through the Mikulski Archive for Space Telescopes¹³ (MAST), which include drizzled science and weight images.

Basic calibrations were performed with the standard IRAF procedures CALACS and CALWF3 for ACS and WFC3 data, respectively. Here we chose a pixel scale of 60 mas pix^{-1} for both optical and IR drizzled images. For the ACS bands we used the "self calibrated" mosaics, which contain additional corrections applied by the STScI team to better account for charge transfer inefficiency (CTI) effects. Similarly, the WFC3 bands were also corrected for a time-variable IR background. We refer the reader to a detailed explanation of the data reduction performed by STScI¹⁴ (Koekemoer et al. in prep). For the fields that have been observed prior to the HFF program, we combine all the available data using the weight maps included in the HFF data release. Table 4 summarizes the exposure times and the depth achieved in each filter for each of the fields. The limiting magnitude in each filter was calculated using $0.4''$ diameter apertures randomly distributed in the image to sample the sky variance before fitting the resulting distribution. The quoted depth is given at the $3\text{-}\sigma$ level.

3. SAMPLE SELECTION

3.1. Photometric catalogs

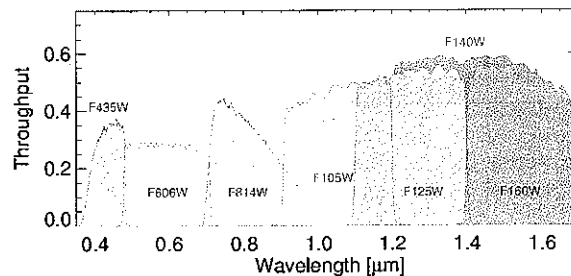


FIG. 1. Throughput curves of the HST/ACS and WFC3 filter set used for the HFF observations.

We constructed the photometric catalogs in each field using the SExtractor software (Bertin & Arnouts 1996). We first matched all the images to the same point spread function (PSF) using a model based on the largest PSF derived with TinyTim (Krist et al. 2011). In order to increase the sensitivity to faint sources, we created deep images by using inverse variance (IVM) weight maps to combine all IR frames for $z \sim 7$ sources and F125W, F140W and F160W for $z \sim 8$ sources, respectively. This deep image is used for source detection in the SExtractor dual image mode, while individual images are used for photometry, weighted by the individual IVM images.

One important limitation for the detection of the faintest sources in the cluster fields is the contamination from the intra cluster light (ICL) in the central region of the cluster. This diffuse light, concentrated primarily in the central region of the cluster, is due to the brightest cluster galaxy, (BCG) or other bright cluster members and the tidal stripping of stars from interacting galaxies during the merging history of the cluster (e.g. Montes & Trujillo 2014). In order to mitigate the contamination of sources close to the cluster center we subtract a median filtered image from the detection one with a filter size of $\sim 2'' \times 2''$, while the photometry is performed on the original image. The adopted filter size is a tradeoff between the removal of extended bright emission and the appearance of artifacts close the bright galaxies of the cluster core. We note that the median filtered detection image allows us to detect five more sources on average in cluster fields compared to the original images. This is also important for the visual inspection of galaxy candidates as the background flux is much lower in these corrected images.

In order to estimate the total flux errors, we ran simulations of galaxies with different sizes and profiles (cf. the completeness simulations in Section 5.1) and compared the recovered fluxes to the input values. We find that the median-filtering approach achieves an uncertainty of 0.5 mag for the faintest sources ($H_{140} \sim 28 - 29$ mag), by using a local background estimate with `back_size = 6` in SExtractor for photometry. In addition, the following extraction parameters were set to improve the detection of the faintest sources in the field: `detect_minarea=2` and `detect_thresh=1.5` for the detection and `deblend_nthresh=16` for source deblending. Two types of photometry are used throughout this work. The isophotal magnitude (ISO) is used to compute the colors and ensure that the same aperture is used across the filter set. The total flux is measured within the Kron radius using the AUTO magnitude, which hereafter is

¹³ <https://archive.stsci.edu/prepds/frontier/>

¹⁴ <http://www.stsci.edu/hst/campaigns/frontier-fields/>

TABLE 1
HST OBSERVATIONS OF THE HFF FIELDS

A2744				MACS0416			MACS0717		
Filter	#Orbits	Depth ^a	Obs Date	#Orbits	Depth	Obs Date	#Orbits	Depth	Obs Date
WFC3/F160W	24	28.3	Oct/Nov 2013	24	29.1	Jul/Sep 2014	26	28.8	Feb/Mar 2015
WFC3/F140W	10	29.1	Oct/Nov 2013	10	28.8	Jul/Sep 2014	12	28.5	Feb/Mar 2015
WFC3/F125W	12	28.6	Oct/Nov 2013	12	28.8	Jul/Sep 2014	13	28.6	Feb/Mar 2015
WFC3/F105W	24	28.6	Oct/Nov 2013	24	29.2	Jul/Sep 2014	27	28.9	Feb/Mar 2015
ACS/F814W	42	29.4	Jun/Jul 2014	50	29.2	Jan/Feb 2014	46	29.3	Sep/Nov 2014
ACS/F606W	10	29.4	Jun/Jul 2014	13	29.1	Jan/Feb 2014	11	28.6	Sep/Nov 2014
ACS/F435W	18	28.8	Jun/Jul 2014	21	30.1	Jan/Feb 2014	19	29.5	Sep/Nov 2014

A2744				MACS0416			MACS0717		
Filter	#Orbits	Depth	Obs Date	#Orbits	Depth	Obs Date	#Orbits	Depth	Obs Date
WFC3/F160W	24	28.5	Jun/Jul 2014	28	28.5	Jan/Feb 2014	26	28.7	Sep/Nov 2014
WFC3/F140W	10	28.2	Jun/Jul 2014	12	28.6	Jan/Feb 2014	10	28.3	Sep/Nov 2014
WFC3/F125W	12	28.34	Jun/Jul 2014	12	28.6	Jan/Feb 2014	14	28.5	Sep/Nov 2014
WFC3/F105W	24	28.6	Jun/Jul 2014	28	28.9	Jan/Feb 2014	24	28.6	Sep/Nov 2014
ACS/F814W	42	29.0	Oct/Nov 2013	42	28.9	Jul/Sep 2014	43	28.7	Feb/Mar 2015
ACS/F606W	16	28.9	Oct/Nov 2013	10	28.7	Jul/Sep 2014	16	28.4	Feb/Mar 2015
ACS/F435W	28	29.7	Oct/Nov 2013	18	29.1	Jul/Sep 2014	26	29.1	Feb/Mar 2015

^aThe depth of the images are 3σ magnitude limits measured in a $0.4''$ aperture.

used as the total magnitude. We modified the magnitude errors to account for pixel-to-pixel noise correlations in the drizzled images following Casertano et al. (2000). Finally, the individual catalogs were matched into a master photometric catalog and cleaned from spurious sources.

3.2. High-redshift Dropout Selection

We adopted the Lyman break selection technique (Steidel et al. 1996; Giavalisco 2002) to detect high-redshift galaxy candidates. The selection is based on color-color criteria to sample the UV continuum dropout of the IGM absorption blueward of $\text{Ly}\alpha$ caused by the intervening hydrogen along the line of sight, and minimize contamination by low-redshift red objects at the same time. Following (Atek et al. 2015), we use the following criteria to select $z = 6 - 7$ galaxies:

$$\begin{aligned}
 (I_{814} - Y_{105}) &> 1.0 \\
 (I_{814} - Y_{105}) &> 0.6 + 2.0(Y_{105} - J_{125}) \\
 (Y_{105} - J_{125}) &< 0.8
 \end{aligned} \tag{1}$$

In addition, we require all sources to be detected in the deep IR image and in at least two IR bands with 4σ significance or higher. We also reject any galaxy that shows up at a significant (at 1.5σ) level in the deep optical combination of $B_{435} + V_{606}$ images. To preserve the Lyman break criterion in the case of non detection in the I_{814} image, we assign a 2σ limiting magnitude to this band. Therefore, we select only galaxies with Y_{105} at least 1 mag brighter than the I_{814} depth. Similarly, we select $z \sim 8$ galaxies that satisfy:

$$\begin{aligned}
 (Y_{195} - J_{125}) &> 0.5 \\
 (Y_{105} - J_{125}) &> 0.3 + 1.6(J_{125} - H_{140}) \\
 (J_{125} - H_{140}) &< 0.5
 \end{aligned} \tag{2}$$

We require a 4σ detection in a deep $J_{125} + H_{140} + H_{160}$ image and the detection in the stacked deep optical image (four ACS filters) to be less than 1.5σ . The deep optical image is about one magnitude deeper than the faintest

TABLE 2
NUMBER OF GALAXY CANDIDATES IN EACH FIELD

Field	$z = 6 - 7$	$z = 8$
A2744	45	7
MACS0416	33	3
MACS0717	41	3
A2744 par	44	3
MACS0416 par	33	5
MACS0717 par	31	5

sources (detected in the IR) accepted in our sample. In Figure 2 we show the result of our selection procedure. The green circles show the location of all high-redshift galaxies identified in this work in the color-color diagram. The dotted and solid lines represent the color evolution of low-redshift elliptical galaxies and high-redshift starbursts, respectively, constructed from Coleman et al. (1980) and Kinney et al. (1996) templates. An attenuation of $A_V = 1, 2, 3$ is applied to the blue, orange, and red curves, respectively. The shaded region represent our adopted selection window, which was chosen to minimize contamination from low-redshift interlopers and red objects such as cool stars represented by magenta points. Our final sample, combining cluster and parallel fields, contains 227 galaxies at $z \sim 6 - 7$, according to the selection based on Eq. 1, and 25 galaxies at $z \sim 8$ based on the selection of Eq. 2. We find roughly the same number of candidates in the cluster and the parallel fields. While the survey area is smaller in the lensed field, the magnification bias balances the number density, allowing the detection of much fainter galaxies than in the parallel fields.

3.3. Sources of Contamination

We now discuss the main sources of contamination for our high-redshift sample of galaxies. A first possibility is spurious sources produced by detector artifacts, diffraction features, or photometric errors scattering into the color selection space. We have visually inspected all of our candidates to identify such contaminants. Most of the spurious sources were already cleaned using the

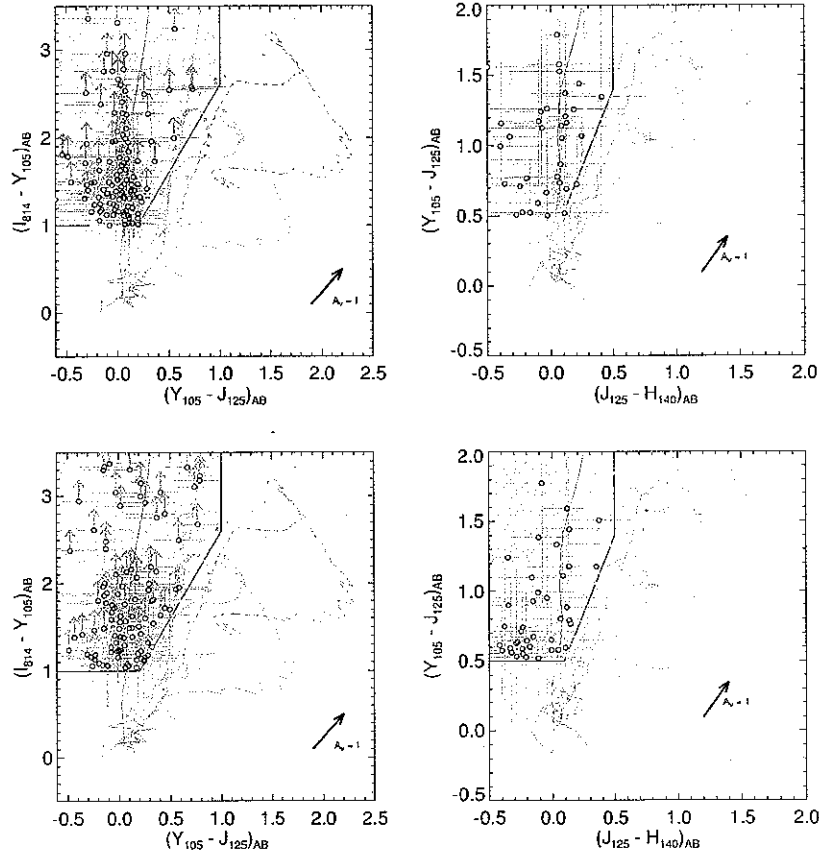


FIG. 2. The Color-color selection windows (represented by the shaded regions) of high- z dropout candidates. The left and right panels show the selection of $z \sim 7$ and $z \sim 8$ galaxies, respectively. The top panels are for cluster fields and bottom panels for parallel fields. The green circles (with associated $1-\sigma$ uncertainties) represent a compilation of all galaxies satisfying our selection criteria and included in our $z > 6$ samples. The dotted line show the color track, i.e. the redshift-evolution of colors, of low-redshift elliptical galaxies generated from Coleman et al. (1980) templates, whereas solid lines show starburst galaxies generated from Kinney et al. (1996) templates. The color code, from blue to red, illustrates the impact of extinction in steps of $A_V = 1$. Finally, the magenta points denotes the color track of stars generated from Chabrier et al. (2000) templates.

weight maps that exclude the frame edges and the IR blobs identified by STScI calibrations. We find that the remaining artifacts are mostly diffraction spikes of bright stars and few very bright cluster galaxies in the field. Fake sources due to photometric noise are also minimal since we require the detection in at least two different bands.

The second potential source of contamination are low-redshift galaxies that show similar colors to high-redshift candidates. According to our stringent spectral break criteria, such sources need to show an extremely red continuum or a large Balmer/4000Å break. Dust obscured galaxies would be excluded by our selection because they would exhibit relatively red colors redward of the break. The remaining galaxies that can enter the selection need to have a large Balmer break and relatively blue continuum at longer wavelengths. The existence of these peculiar objects have been discussed in Hayes et al. (2012) where an SED of a young burst superimposed on an old stellar population with an extreme 4000Å break can mimic the Lyman break colors. Such objects should be very rare and represent only a minor contamination. Moreover, in the case of a significant contamination by

these interlopers, we expect to detect the blue continuum by stacking the optical images of our candidates, which is not the case.

Alternatively, the Lyman break colors can also be mimicked by extremely strong emission lines in $z = 1 - 3$ star-forming galaxies (e.g., Atek et al. 2011, 2014a; van der Wel et al. 2011). In this class of galaxies the contribution of the emission line flux to the total broadband flux is about 25% on average and can reach 85% (Atek et al. 2011) and can have important implications, not only for the high-redshift galaxy selection, but also on the age and stellar mass derived from SED fitting (Schaerer & de Barros 2009; Wilkins et al. 2013; Schenker et al. 2014a; Pérenin et al. 2015; Huang et al. 2015). In our case, the contamination of one band can lead to an artificial Lyman break, because the faint continuum remains undetected blueward of the break. Atek et al. (2011) estimate that the optical data should be about 1 magnitude deeper than the detection band to be able to rule out such interlopers. Our criteria impose a minimum break of 0.8 mag between the object flux and the limiting magnitude of the optical band. Moreover, such a contamination should be even smaller be-

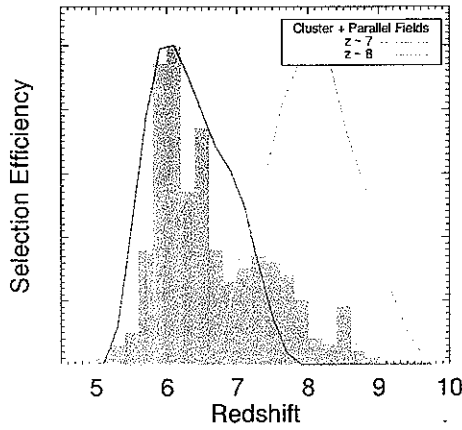


FIG. 3. The redshift selection function for all the fields at $z \sim 7$ and $z \sim 8$. The curves are the result of our completeness simulations in recovering the input sources marginalized over the redshift, while the shaded histograms represent the actual selected galaxies in our sample. The median redshift of our $z \sim 8$ is significantly lower than expected due to the depth of the observations and the redshift evolution of the luminosity function.

cause our stacked candidate images, which are at least 1.3 mag deeper than the candidate’s flux, show no significant detection in the optical bands shortward of the break.

In general, observationally, estimating the contamination rate of such sources proves very difficult at $z > 6$ because it hinges on large spectroscopic follow-up programs that aim at detecting the redshifted Ly α emission line in these galaxies. In addition to the large amount of telescope time needed to reach the required depth, the absence of Ly α emission does not exclude the high-redshift nature of the associated objects. Indeed, the increasing neutral hydrogen fraction at $z > 6$ easily absorbs and diffuses Ly α photons (Atek et al. 2009; Stark et al. 2010; Caruana et al. 2014; Schenker et al. 2014b). Therefore, one must rely on simulated colors based on galaxy spectral templates to estimate the contamination rate of low redshift galaxies, which thus has been found to be less than 10 % in most studies (e.g. Oesch et al. 2010; Bouwens et al. 2015b).

Another source of contamination to consider are low-mass stars. As we can see in Fig. 2 a significant fraction of these stars (shown in magenta) can have similar colors to high redshift galaxies. However, we have excluded any source that has a SExtractor stellarity parameter greater than 0.8 to minimize point-like objects in our sample. Also, the number density of cool stars ranges from 0.02 to 0.05 arcmin² derived from observations with similar or better depths (Bouwens et al. 2015b). This translates to about 1 contaminant in the entire $z \sim 7$ and $z \sim 8$ galaxy samples. We also consider transient events as possible contaminants. Because observations are taken at two different epochs (cf. Table 4), a supernova explosion may appear in IR images and not in the optical ones and could be selected as a dropout candidates. Similar colors can also be obtained in the case where the IR data were taken first and the supernova faded until the optical images. Again, different reasons point toward a negligible contamination level from transient sources.

Such objects would show point-like profiles that would be excluded by the stellarity criterion explained above and our additional visual check for unresolved sources. The detection rate of supernovae is also very small and, so far, none of SN detections in the dedicated search program of the HFF (e.g., Rodney et al. 2015) were inadvertently selected as high- z galaxies in the different HFF studies (Atek et al. 2014b, 2015; Finkelstein et al. 2014; Ishigaki et al. 2015a).

4. HFF CLUSTER MASS MODELS

In order to exploit the full potential of the HFF cluster lenses, we first need a robust model describing the total mass distribution and the lensing properties. In an effort to provide the community with all the required lensing maps to interpret background source observations, several groups have submitted their models prior to the HFF (e.g., Bradač et al. 2005; Merten et al. 2011; Johnson et al. 2014; Richard et al. 2014; Grillo et al. 2015; Coe et al. 2015). With the availability of new deep *HST*, *Spitzer*, and spectroscopic observations, the models were significantly improved thanks in particular to the discovery of a large number of multiple images (Diego et al. 2014; Grillo et al. 2015; Zitrin et al. 2014; Wang et al. 2015; Jauzac et al. 2015, 2014; Ishigaki et al. 2015a). Most notably, using a set of ~ 180 and ~ 200 multiple images in A2744 and MACS0416, respectively, Jauzac et al. (2015, 2014) reconstructed the projected cluster mass down to a precision of $\sim 1\%$, which represents a significant improvement over pre-HFF mass models.

In the present study, we use the most recent cluster models based on full-depth HFF observations and constructed by the CATS (Clusters As Telescopes) team. The models developed for the first two HFF clusters, namely A2744 and MACS0416, are already available on the MAST archive. However, STScI has started a new mass mapping initiative so all the teams provide HFF mass models of both clusters taking advantage of the full depth of the HFF data, but using the same inputs. These ‘unified’ models will be submitted to STScI as part of our answer to the HFF call for the community to provide updated lensing maps¹⁵. The mass modeling of the third cluster MACS0717 will be presented in details in Limousin et al. (2015).

The mass reconstruction of each of the clusters is explained in the publications listed above. Here we briefly describe the key points of the mass modeling procedure.

The CATS HFF mass models are built using the Lenstool¹⁶ software (Kneib 1993; Jullo et al. 2007; Jullo & Kneib 2009). For the strong-lensing analysis, we are using a parametric approach that consists of modeling the cluster mass distribution using both cluster-scale and galaxy-scale halos. Galaxy-scale halos are important in the mass modeling, as multiple image configurations are impacted by their location. Indeed if a multiple image is close to a cluster galaxy, then the distortion created by the lensing effect will be a combination of both the cluster potential itself, as well as a galaxy-galaxy lensing effect due to the potential of the cluster member. The potentials describing these components are modeled using

¹⁵ <https://archive.stsci.edu/prepds/frontier/lensmodels/>

¹⁶ <http://projets.lam.fr/projects/lenstool/wiki>

pseudo-isothermal elliptical mass distribution (PIEMD, Elíasdóttir et al. 2007).

For MACSJ0416, we used 149 of the most secure multiple images (over the 194 identified) to constrain the mass model. Our best fit mass model of the cluster inner core consists of two cluster-scale halos, well-aligned with the light peaks from the two brightest cluster galaxies (BCGs), plus 98 galaxy-scale halos, corresponding to the brightest cluster galaxies. For Abell 2744, we used 154 most secure multiple images (over the 181 identified). Our best-fit mass model consists of two cluster-scale halos to describe the dark matter distribution on large-scale, combined with 733 galaxy-scale halos, describing the distribution of cluster galaxies. Both models delivered really good errors in the predictions of multiple image positions, an rms of $0.68''$ and $0.79''$ for MACSJ0416 and Abell 2744 respectively. For MACS J0717, we used the 140 most secure multiple images (over 163 identified) to constrain a mass model composed of 4 large scale DM haloes, plus 92 galaxy scales haloes, corresponding to the brightest cluster galaxies. More details of the SL analysis will be given in Limousin et al. (2015)

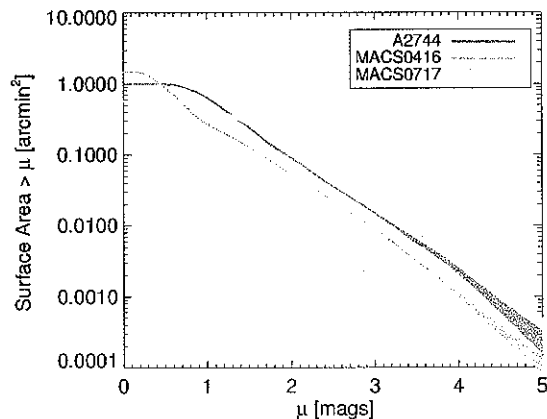


FIG. 4. The cumulative surface area in the source plane at $z \sim 7$ as a function of the amplification factor (in magnitudes) derived from the mass modeling of the three HFF clusters. Uncertainties in the surface area are also shown at the $1-\sigma$ level.

4.1. Multiple Images

In addition to flux magnification, strong lensing produces multiple images of the same background galaxy. Therefore, we need to identify multiple images to be removed from the galaxy number counts before computing the UV LF. For each galaxy, we predict the position of potential counter images by using the mass model and Lenstool to project its position into the source plane before lensing back the source into the image plane, where Lenstool predicts the position of all the multiple images. In the vicinity of these positions, we look for dropout sources that show similar colors and photometric redshifts. In the case of well resolved sources, we also verify that they have similar morphological and geometrical symmetries constrained by the lensing model. In A2744 we find three systems at $z \sim 7$, with a total of nine multiple images, which were discussed in Atek et al. (2014b). In MACS0416, we identify seven systems at $z \sim 7$.

In the case of MACS0717, the multiple image region extends beyond the WFC3 field of view. Therefore, we expect each of the galaxy candidates to have counter-images with similar flux/magnification ratios given the lensing geometry of this cluster. The uncertainty on the image position predicted from the lensing model of MACS0717 is about $rms = 1.9$ arcsec, much larger than in the two other clusters (cf. Limousin et al. 2015). Moreover, the shape of the critical line is not as well constrained, which makes it harder to identify multiple-image systems. From the Lenstool model, we estimate the multiplicity in MACS0717 to be three on average in the WFC3 field of view. Therefore, we divide the number of galaxy candidate by this multiplicity to obtain the final number counts used for the LF calculation.

5. THE GALAXY UV LUMINOSITY FUNCTION

We now turn to computing the galaxy UV luminosity function at $z \sim 7$ and $z \sim 8$ by combining the high- z candidate samples from all the fields. For each field we calculate the completeness function and the effective survey volume to derive the LF. While in the parallel fields we proceed with standard completeness simulations widely used in analyzing blank field surveys, we need to take into account the lensing effects in the cluster fields.

The primarily goal of the HFF program is to extend the current limits of deep galaxy surveys by using strong lensing to magnify intrinsically faint sources behind the galaxy clusters. In parallel, the survey area in the source plane is significantly reduced for high magnifications. Consequently, the efficiency of a given lensing cluster in probing the high-redshift Universe, which can be quantified by the number of magnified sources discovered, is the result of a trade off between the amplification factor μ and the source plane area σ . We show in Fig. 4 the cumulative survey area as a function of the amplification factor.

While the survey area in the image plane corresponds to the WFC3 field of view, i.e. about 4.7 arcmin^2 , we can see the total survey area is reduced to about 0.6 to 1 arcmin^2 in cluster fields. Wong et al. (2012) define the cluster magnification power as the cross section for magnifying a source above minimum threshold of $\mu = 3$ (see also Richard et al. 2014). Although the precise value of the threshold is somewhat arbitrary, the number of high-redshift detections in our galaxy samples peaks around $\mu = 3$, or $\mu \sim 1.2 \text{ mag}$ (cf. the top panel of Fig. 6. For this typical value, the total survey area drops to values of around 0.2 to 0.5 arcmin^2 for the three clusters of this study.

The parallel fields are located about 6 arcmin away from the cluster lenses with a small yet non negligible magnification. It is important to assess the lensing effects on the UV LF, both on the magnitude and on the survey volume for these regions as well. Among the different lensing models available in the HFF project, only Merten et al. (2011) provide a wide-field magnification map for the three clusters that covers their flanked fields. We used the magnification maps at $z \sim 9$ that have a resolution of $25 \text{ arcsec pix}^{-1}$ to estimate the flux amplification of galaxies. Since no deflection map is available, the volume reduction is estimated by dividing the area by the amplification map. The harmonic mean of the amplification factor ranges from 1.11 to 1.23 in the three

fields, with typical errors of about 10%. The inclusion of the lensing effects introduce only small changes in the UV LF, within the error bars, basically shifting M_{UV} to slightly fainter values and ϕ to higher values.

We now describe how we estimate the effective survey volume for each cluster by combining the source plane area with the redshift selection function and the recovery rate of simulated galaxies as a function of different galaxy and lensing parameters.

5.1. Completeness Simulations

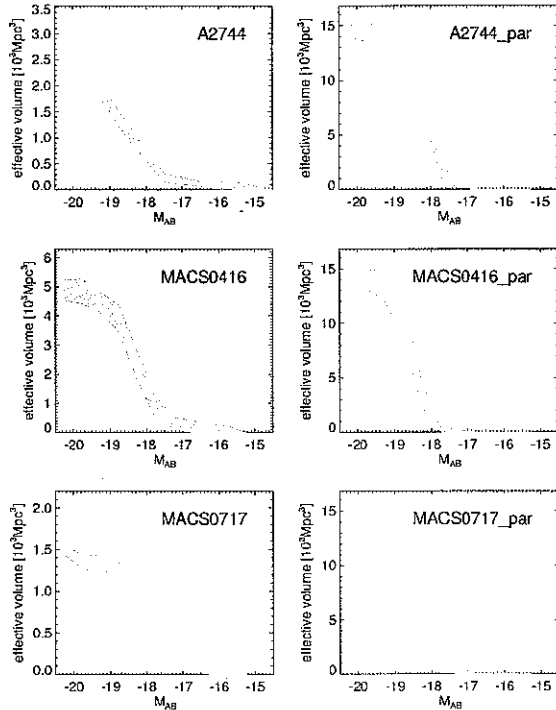


FIG. 5. The effective survey volume as a function of absolute UV magnitude for the redshift $z \sim 7$ sample. Each curve is based on our completeness simulations for each cluster or parallel field. The total area for the cluster fields is corrected for lensing effects whereas the blank field area is based on the full WFC3 field of view (cf. text for a detailed explanation). The blue colored region represents the 95% confidence intervals of the completeness estimate.

Following Atek et al. (2015) we run extensive Monte Carlo simulations to assess the completeness level as a function of the intrinsic magnitude. A total of 10,000 galaxies were simulated for each of the six fields. We take the galaxy profile into account by creating two samples of exponential disks and de Vaucouleur shapes (Ferguson et al. 2004; Hathi et al. 2008). These galaxy profiles are then distorted by gravitational lensing according to our mass model. As for galaxy sizes, we adopt a log-normal distribution with a mean half light radius (hlr) of $0.15''$ and $\sigma=0.07''$. For consistency with previous results (Ferguson et al. 2004; Bouwens et al. 2004; Hathi et al. 2008; Oesch et al. 2010; Grazian et al. 2011, 2012; Huang et al. 2013; Oesch et al. 2014), the distribution is based on the sizes derived from spectroscopically confirmed LBG samples at $z \sim 4$

(Vanzella et al. 2009), while accounting for a redshift evolution of the intrinsic physical size of galaxies with a factor of $(1+z)^{-1}$. *HST* observations of dropout galaxies at $z > 6$ redshifts also reveal small sizes of less than $0.3''$ (Mosleh et al. 2012; Ono et al. 2013) and smaller for galaxies fainter than $M_{UV} - 21$ mag. More recently Kawamata et al. (2015) measured the size of lensed dropout galaxies in the HFF cluster A2744 and report similar results. Small sizes around $0.1''$ have also been observed in lensed galaxies by Coe et al. (2013); Zitrin et al. (2014). In addition, based on the results of (Huang et al. 2013), we adopt a size luminosity relation $r \propto L^\beta$, with $\beta = 0.25$, for our simulations (see also Mosleh et al. 2012; Kawamata et al. 2015).

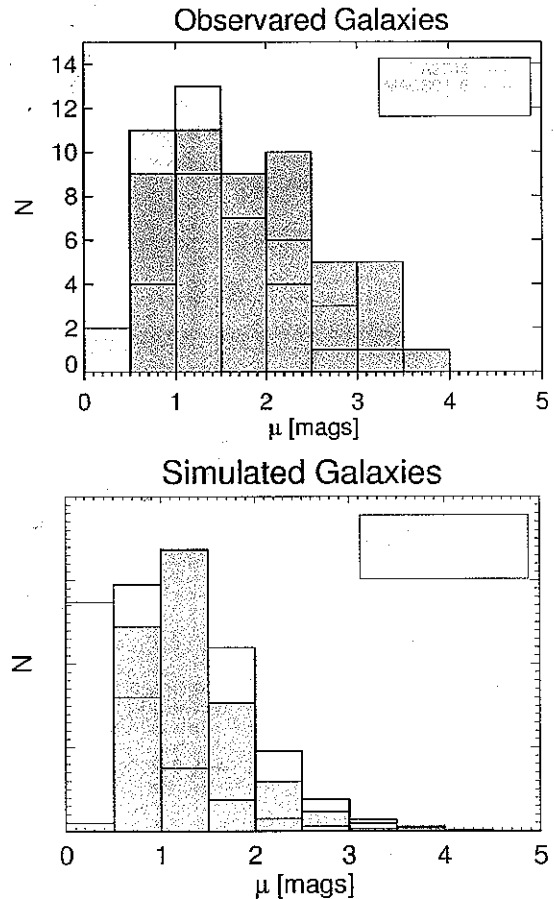


FIG. 6. The distribution of the amplification factor, expressed in magnitudes. Top panel shows the distribution for the $z \sim 7$ candidates in each of the lensing clusters. For comparison, the bottom panel presents the result of the completeness simulations in the three cluster cores, which contain about 80,000 objects each.

After assigning random redshifts in the range $[5.5-7.5]$, we simulate galaxy magnitudes in each *HST* band using star-forming SED templates from (Kinney et al. 1996). In this step, we also assign random intrinsic absolute magnitude (in the rest-frame UV) in the range $M_{UV} = [-14, -24]$ mag. We then include 10 simulated galaxies in the actual images of each band for a total of thousand images. This is where the cluster and parallel fields are treated

TABLE 3
THE BEST FIT $z \sim 7$ SCHECHTER PARAMETERS IN EACH INDIVIDUAL FIELD

Field	M_{UV}^* [AB mag]	α	$\log_{10} \phi^*$ [Mpc $^{-3}$]
A2744	-20.92 ± 0.64	-2.03 ± 0.13	-3.56 ± 0.45
MACS0416	-21.04 ± 1.10	-2.07 ± 0.16	-3.67 ± 0.59
MACS0717	-21.13 ± 1.52	-2.02 ± 0.22	-3.73 ± 0.83
A2744 par	-20.99 ± 1.05	-2.03 ± 0.19	-3.62 ± 0.55
MACS0416 par	-20.94 ± 1.19	-2.02 ± 0.20	-3.59 ± 0.70
MACS0717 par	-20.89 ± 1.10	-2.00 ± 0.21	-3.54 ± 0.62

differently in the simulations. In the cluster fields, galaxies are not included directly in the images but are simulated in the source plane. They are lensed into the image plane using the corresponding mass model. This way we ensure that all the lensing effects including magnification, shape distortion, position relative to the critical line are fully taken into account. Finally, for all the images, we run the same procedure used to select high- z galaxies and determine the completeness function, which represents the fraction of the original galaxies recovered in our selection as a function of the parameters described above. This completeness function is in turn incorporated in the computation of the effective volume in each magnitude bin following the equation:

$$V_{eff} = \int_0^\infty \int_{\mu > \mu_{min}} \frac{dV_{com}}{dz} f(z, m, \mu) d\Omega(\mu, z) dz \quad (3)$$

where μ_{min} is the minimum amplification factor μ_{min} required to detect a galaxy with a given apparent magnitude m and V_{com} is the comoving volume. $f(z, m, \mu)$ is the completeness function that depends on the redshift z , apparent magnitude m , and amplification factor μ , and $d\Omega(\mu)$ is the area element in the source plane, which is function of magnification and redshift.

Figure 5 presents the results of our effective volume estimates in each field, marginalized over the intrinsic absolute UV magnitude. The extent of each filled region represents the 68% confidence intervals. We can clearly see the importance of gravitational lensing in extending the survey depth to fainter galaxies. While the blank fields completeness drop abruptly before $M_{UV} = -18$ mag, it becomes shallower in cluster fields and extends down to $M_{UV} = -15$ mag, although at a level of 10% or less.

5.2. The UV LF at $z = 6 - 7$

Using the derived effective volume as a function of absolute magnitude we compute the UV luminosity function following the equation:

$$\phi(M)dM = \frac{N_i}{V_{eff}(M_i)}, \quad (4)$$

where N_i and M_i are the number of galaxies and the absolute magnitude, respectively, in each magnitude bin. Following most of the studies in the literature, we choose a bin size of 0.5 mag whereas the magnitude varies from one field to another. The individual LF determinations in each field are presented in Figure 7. The top panels show the results in the cluster fields while the bottom panels are for parallel fields. The LF extends to

$M_{UV} = -18.25$ in blank fields, whereas it reaches a magnitude of $M_{UV} = -15.25$ in cluster fields thanks to the lensing magnification. This gain can already be seen in the completeness function results (see Fig. 5) which reaches fainter magnitudes in cluster fields. Figure 6 also shows the distribution of the amplification factor for the dropout samples selected behind the three lensing clusters. The flux amplification ranges essentially from $\mu \sim 1.25$ to 75, with a median value of 5.2, 3.9, and 4.9 in A2744, MACS0416, and MACS0717, respectively.

In addition to the sample contamination discussed in Sect. 3.3, several sources contribute to the uncertainties of the LF data points. The mass model errors, affecting the magnification and the survey volume (or the source plane area Ω), are propagated into the LF determination in the case of cluster fields. We also include Poisson errors, and the cosmic variance estimate based on the recent results of Robertson et al. (2015) for both cluster and parallel fields. Finally, uncertainties from incompleteness simulations, as seen in Fig. 5 are incorporated in the final LF. The bright-end is dominated by cosmic variance errors and small number counts, while the faint-end error bars reflect mostly the large incompleteness uncertainties at these faint magnitudes and small statistics.

We determine the shape of the rest-frame UV LF by fitting a Schechter function (Schechter 1976) to our data, which has been extensively used to describe the galaxy UV LF across a wide redshift range (e.g. Bunker et al. 2004; Beckwith et al. 2006; Reddy & Steidel 2009; Bouwens et al. 2006; McLure et al. 2009; Ouchi et al. 2009; Wilkins et al. 2010; Oesch et al. 2012; Willott et al. 2013). The Schechter function can be expressed in terms of absolute magnitudes as:

$$\phi(M) = \frac{\ln(10)}{2.5} \phi^* 10^{0.4(\alpha+1)(M^*-M)} \exp(-10^{0.4(M^*-M)}) \quad (5)$$

We perform a Schechter fit to the LF in each field and find a faint-end slope α between -2.0 and -2.07 (cf. Eq. 5). We include in the fit the data points from Bouwens et al. (2015b, black squares) to constrain the bright-end of the LF. The faintest bin is reached in A2744 at $M_{UV} = -15.25$. There is a hint of a shallower slope at those magnitudes that could be the result of large uncertainties in the completeness estimate, which is typically less than 10% in this bin. As a matter of fact, a similar decline is observed in MACS0416 around $M_{UV} = -16.25$, whereas the slope is constantly steep at these magnitudes in A2744. It is clear that the constraints on the LF in MACS0717 are not as good as in the other cluster fields. This is likely the result of larger uncertainties on the lensing model due to the complex structure of the galaxy cluster. For instance, the model prediction for image position has an rms error of ~ 1.9 arcsec, while it is about 0.7 to 0.8 arcsec in MACS0416 and A2744. These kind of deviations in the LF data points can also be observed in A2744 when using an old model based on pre-HFF observations that similarly yielded large uncertainties. Nonetheless, the overall shape of the LF remains consistent with the other clusters' results.

We now combine all the LF constraints from the lensed and parallel fields to compute the most robust UV LF

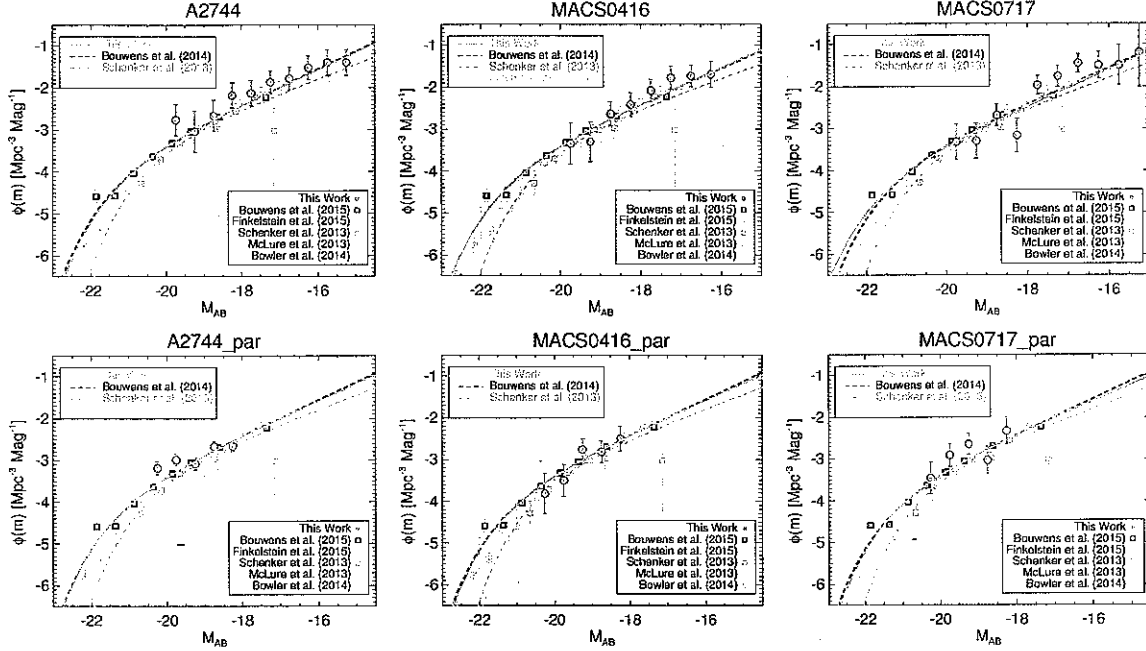


FIG. 7. UV luminosity function at $z \sim 7$ computed individually in each field. The blue circles represent our determination with $1-\sigma$ uncertainties while the blue solid curve is our best Schechter function fit to the LF. We compare our results to previous literature results in blank fields. The black squares and dashed curve are from a compilation of *HST* legacy fields by Bouwens et al. (2015b). We also show the LF determination of Schenker et al. (2013b, blue squares and dashed curve) and McLure et al. (2013, green squares and dashed curve) derived in the UDF12 field. We also include data points on the bright-end of the LF [magenta squares] from a wide area survey by Bowler et al. (2014).

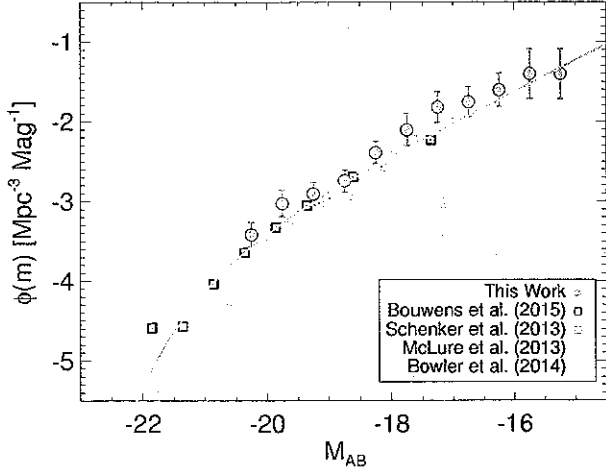


FIG. 8. The combined constraints from the clusters and parallel fields on the UV luminosity function at $z \sim 7$. The color code for the data points and the best fit Schechter function are the same as in Fig. 7.

at $z \sim 7$. The result is shown in Fig. 8. Together with our data points, we show the most recent results from the literature from the blank fields as described in the legend. We ran MCMC simulations with 10^6 realizations to find the best fit to the LF and estimate the uncertainties on the Schechter parameters. We find that the UV LF at $z \sim 7$ has a faint-end slope of $\alpha = -2.04^{+0.13}_{-0.17}$, a characteristic $M^* = -20.89^{+0.60}_{-0.72}$ and $\text{Log}(\phi^*) = -3.54^{+0.48}_{-0.45}$. This is an excellent agreement

with earlier results presented in Atek et al. (2015), where we computed the UV LF in the HFF cluster A2744. In Figure 9, we show the likelihood analysis of the Schechter parameters at $z \sim 7$, marginalized over two parameters at a time. Our results are in good agreement with the most recent results reported from the blank fields (cf. Table 5). In particular, the faint-end slope is close to the determination of Finkelstein et al. (2014) with $\alpha = -2.03^{+0.21}_{-0.20}$ and Bouwens et al. (2015b) with $\alpha = -2.03 \pm 0.14$. Our results show only a slightly steeper faint-end slope than the values reported in the Hubble extreme deep field (Schenker et al. 2013a; McLure et al. 2013) with $\alpha = 1.90^{+0.14}_{-0.15}$ and $\alpha = 1.90^{+0.14}_{-0.15}$, respectively, and agree within the reported uncertainties. Also using the gravitational lensing of the first HFF cluster A2744, Ishigaki et al. (2015a) find a slightly shallower slope of $\alpha = -1.94^{+0.09}_{-0.10}$, which still in agreement with our value within the errors. We also, performed a schechter fit while excluding data points from MACS0717 cluster, which yields very similar parameters: $\alpha = -2.03$, a characteristic $M^* = -20.86$ and $\text{Log}(\phi^*) = -3.52$. Importantly, we note that the uncertainties on the faint-end slope decreases to $\sigma_\alpha \sim 0.1$.

Several theoretical models and cosmological simulations produce predictions for the UV LF at high redshift. For instance, our LF results at $z \sim 7$ are in good agreement with the hydrodynamical simulations of Jaacks et al. (2012) with $M^* = -20.82$ and $\text{Log}(\phi^*) = -3.74$, although they predict a steeper faint-end slope of $\alpha = -2.30$ down to similar lower magnitude limit of $M_{UV} = -15$. Theoretical models by Mason et al. (2015) find a closer slope of $\alpha = -1.95 \pm 0.17$ down to a

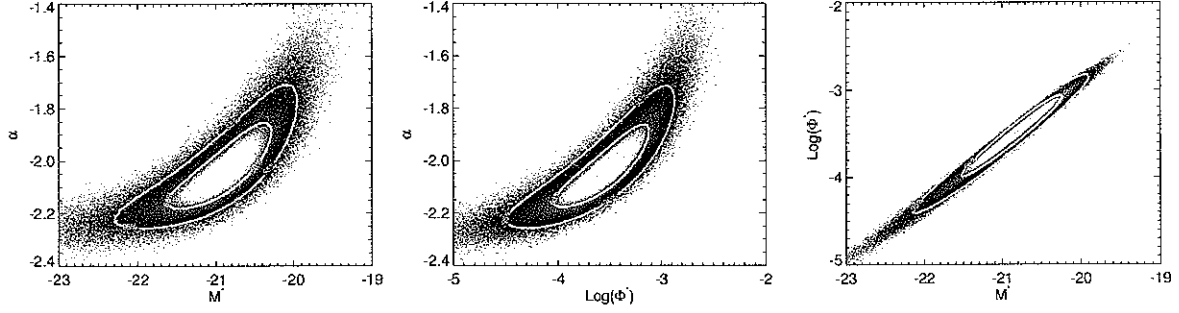


FIG. 9. Likelihood analysis of the Schechter parameters for the UV LF at $z \sim 7$: the faint-end slope α , the characteristic M_{UV}^* and $\text{Log}(\phi^{star})$. The density plots show the result of MCMC simulations marginalized over two parameters in each of the three panels. The orange curves represent the 68% and 95% confidence contours.

TABLE 4
COMBINED CONSTRAINTS ON THE UV LF AT $z \sim 7$

M_{UV}	$\text{Log}(\phi)$	φ_{err}
-20.25	-3.4184	0.1576
-19.75	-3.0263	0.1658
-19.25	-2.9044	0.1431
-18.75	-2.7418	0.1332
-18.25	-2.3896	0.1401
-17.75	-2.1032	0.1990
-17.25	-1.8201	0.1940
-16.75	-1.7548	0.1893
-16.25	-1.6044	0.2117
-15.75	-1.4012	0.3123
-15.25	-1.4012	0.3122

TABLE 5
COMPARISON OF THE BEST FIT $z \sim 7$ SCHECHTER PARAMETERS

Reference	M_{UV}^* [AB mag]	α	$\log_{10} \phi^*$ [Mpc $^{-3}$]
This work	$-20.89^{+0.60}_{-0.72}$	$-2.04^{+0.17}_{-0.13}$	$-3.54^{+0.48}_{-0.45}$
Atek et al. (2015) ^a	$-20.90^{+0.90}_{-0.73}$	$-2.01^{+0.20}_{-0.28}$	$-3.55^{+0.57}_{-0.57}$
Ishigaki et al. (2015a) ^a	$-20.45^{+0.71}_{-0.2}$	$-1.94^{+0.09}_{-0.10}$	$-3.30^{+0.10}_{-0.20}$
Bouwens et al. (2015b)	-21.04 ± 0.26	-2.06 ± 0.12	$-3.65^{+0.27}_{-0.17}$
Finkelstein et al. (2014)	$-21.03^{+0.37}_{-0.50}$	$-2.03^{+0.21}_{-0.20}$	$-3.80^{+0.41}_{-0.26}$
McLure et al. (2013)	$-19.90^{+0.23}_{-0.28}$	$-1.90^{+0.14}_{-0.15}$	$-3.35^{+0.28}_{-0.45}$

^a Using the first HFF cluster A2744

magnitude limit of $M_{UV} = -12$, corresponding to a halo mass of $10^9 M_{\odot}$. Similarly, Kimm & Cen (2014) find a theoretical faint magnitude limit of $M_{UV} = -13$ with a faint-end slope of $\alpha = -1.9$. Semi-analytical models of Dayal et al. (2014) also find a steep faint-end slope of $\alpha \sim -2.02$ at $z \sim 7$ in good agreement with our observations.

5.3. The UV LF at $z \sim 8$

Regarding the redshift $z \sim 8$ LF, we followed the same procedure used for the $z \sim 7$ LF. However, in lensing clusters, the survey volume at $z \sim 8$ is much more significantly reduced than at $z \sim 7$. In the extreme case of MACS0717, the total survey volume is about 460 Mpc^3 , whereas it reaches 3600 Mpc^3 in A2744. Therefore, we expect lower galaxy number counts at $z \sim 8$ for high magnification values. In total, we detect only two galaxies with intrinsic magnitudes fainter than $M_{UV} = -18$, where the uncertainties on the LF estimate are very large (see Fig. 10). Unlike the $z \sim 7$ LF, we do not have strong

constraints on the faint-end part of the LF at $z \sim 8$. At brighter magnitudes the LF is better constrained thanks to the addition of the parallel fields and appears in agreement with previous results in the literature. As discussed in Atek et al. (2015), the redshift 8 galaxy selection in A2744 clearly shows an overdensity (see also Zheng et al. 2014; Ishigaki et al. 2015a), which translates into a large excess in the UV LF. Here we decided not to exclude the entire A2744 field in the combined LF to avoid introducing a well known bias due to cosmic variance (see also Ishigaki et al. 2015b).

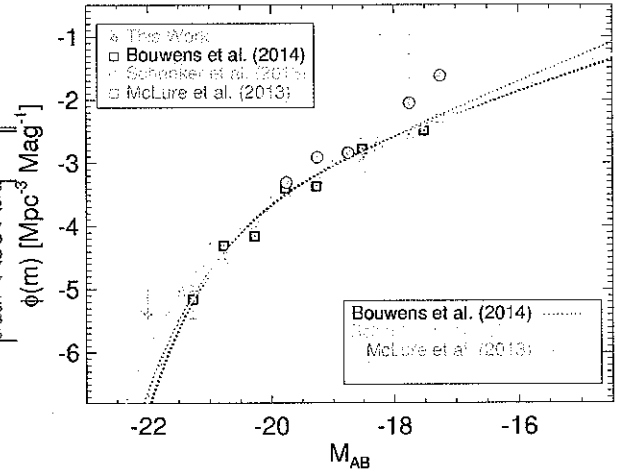


FIG. 10. The combined constraints from the clusters and parallel fields on the UV luminosity function at $z \sim 8$. The red circles represent our LF determination while the red squares are taken from Schenker et al. (2013a), the black squares from Bouwens et al. (2015b), and the blue squares from McLure et al. (2013). The green curve is the unbinned UV LF of Schmidt et al. (2014). The best Schechter fits from the same literature results are also shown with dotted lines (cf. legend in the inset).

5.4. Implications for cosmic reionization

In this study, we reliably extend the UV LF to unprecedented depth, and put strong constraints on the faint-end slope down to $M_{UV} = -15.25$. The most important result is that the faint-end slope remains very steep down to a luminosity of $0.005 L^*$, the characteristic UV luminosity at $z \sim 7$. Also, thanks to the combination of six fields, we

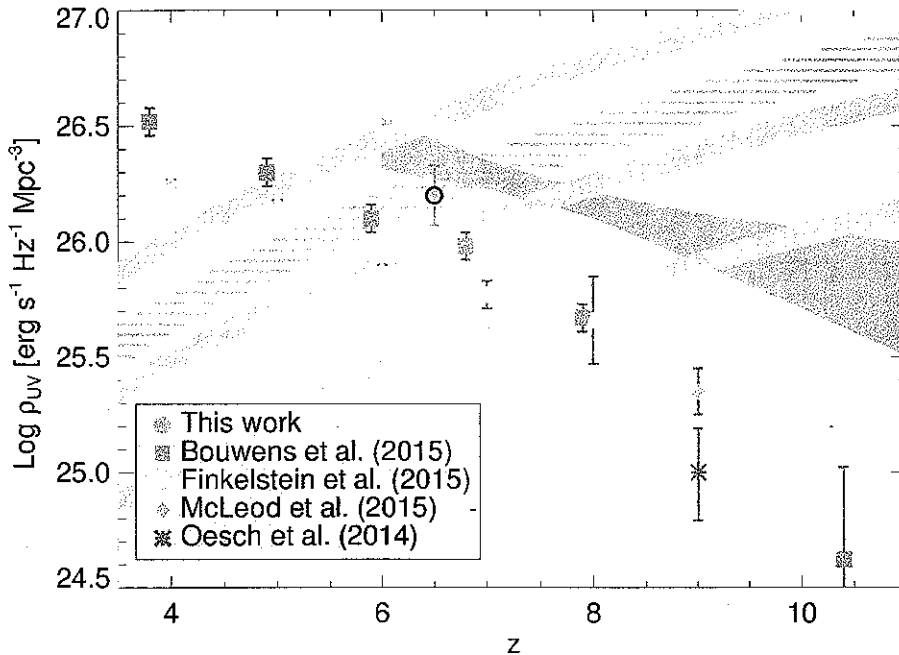


FIG. 11. The evolution of the UV luminosity density of galaxies with redshift. Our determination (blue circle) is based on the integration of the UV LF down to $M_{UV} = -15$. The brown squares represent the most recent results of Bouwens et al. (2015b) who integrate the UV LF down to $M_{UV} = -17$. The orange triangles are the result of (Finkelstein et al. 2014) who integrate to conservative limit of $M_{UV} = -18$. The red diamond and the green asterisk are the UV luminosity densities derived at $z \sim 9$ by McLeod et al. (2015) and Oesch et al. (2014), respectively. Both studies integrate the UV LF down to an observational limit of $M_{UV} = -17.7$, which correspond to a star formation rate of $SFR_{lim} = 0.7 M_{\odot} \text{ yr}^{-1}$. All the error bars indicate $1-\sigma$ uncertainties. The shaded grey region show denote the UV luminosity density based on the ionizing emissivity (Madau et al. 1999) required to maintain the IGM ionized at a certain redshift for and escape fraction of ionizing radiation of $f_{esc} \sim 20\%$ and three different values for the clumping factor $C_{HII} = 1, 3$, and 10 . The red shaded region shows the 68% confidence interval for the evolution of the galaxy UV luminosity density based on constraints on the ionizing emissivity of Bouwens et al. (2015a). The green hatched region corresponds to the UV luminosity density required to maintain the IGM ionized for a clumping factor of $C_{HII} = 3$ and escape fraction in the range $f_{esc} = 10 - 15\%$, which encompasses our value at $z \sim 7$.

have significantly reduced the uncertainties on the faint-end slope to about 5%. With these strong constraints in hand, we can now integrate the UV LF to derive the UV luminosity density at $z \sim 7$. The main advantage of these observations is the ability to set the integration lower limit to $M_{UV} = -15$. Unlike previous results that extrapolate the UV LF to lower magnitude, we use an observational constraint to estimate the ultraviolet photon budget from galaxies. For instance, based on the UV LF results in blank fields, Bouwens et al. (2015b) report a total UV luminosity density of $\text{Log}(\rho_{UV}) = 25.98 \pm 0.06 \text{ erg s}^{-1} \text{ Hz}^{-1} \text{ Mpc}^{-3}$ and Finkelstein et al. (2014) a value of $\text{Log}(\rho_{UV}) = 25.77 \pm 0.06 \text{ erg s}^{-1} \text{ Hz}^{-1} \text{ Mpc}^{-3}$ at $z \sim 7$. Both values are computed down to $M_{UV} = -17$ AB, which is the limiting magnitude of their observations. Here we compute the UV luminosity density down to $M_{UV} = -15$ and find $\text{Log}(\rho_{UV}) = 26.2 \pm 0.13 \text{ erg s}^{-1} \text{ Hz}^{-1} \text{ Mpc}^{-3}$.

In order to determine whether the UV luminosity produced by galaxies is sufficient to reionize the IGM, one needs to estimate additional parameters, for which observational constraints remain challenging. First, the conversion factor from the UV luminosity to ionizing radiation ξ_{ion} ($\text{erg s}^{-1} \text{ Hz}$) depends on the star-formation history of galaxies. The value of ξ_{ion} is generally constrained using stellar population models of early galax-

ies and the observed UV slope β of $z > 6$ galaxies (Bolton & Haehnelt 2007; Kuhlen & Faucher-Giguère 2012; Robertson et al. 2015). Once the ionizing photon production is determined we need to estimate its escape fraction f_{esc} from galaxies to ionize the IGM. Direct observational constraints of f_{esc} are difficult, especially at $z > 6$ because of the high opacity of the intervening hydrogen residuals on the line of sight. Many studies and deep surveys were dedicated to the search of ionizing continuum (Lyman continuum, LyC) escape from $z < 4$ galaxies. Very few detections were reported however and show very low escape fractions relative to the UV radiation, of the order of few percent (Shapley et al. 2006; Iwata et al. 2009; Siana et al. 2010; Nestor et al. 2013). More recently, de Barros et al. (2015) reported a spectroscopic detection of Lyman continuum emission in a $z \sim 3.2$ galaxy with a relative escape fraction¹⁷ of $f_{esc} \sim 65\%$.

In Figure 11, we show our determination of the UV luminosity density together with the most recent results in the literature as a function of redshift including Bouwens et al. (2015b) and Finkelstein et al. (2014) at $z = 4 - 8$, Bouwens et al. (2015b) at $z = 10$,

¹⁷ the relative escape fraction the ratio between the fraction of escaping Lyman continuum photons and the fraction of escaping photons at 1500 Å (Steidel et al. 2001)

McLeod et al. (2015) and Oesch et al. (2014) at $z \sim 9$. While their values are based on the integration of the UV LF down to a magnitude limit between $M_{UV} = -18$ and $M_{UV} = -17$, our integration limit is two magnitudes fainter at $M_{UV} = -15$. There is a significant difference between the results of Bouwens et al. (2015b) and Finkelstein et al. (2014) due to the fact that the UV LF in the former study extends one magnitude deeper than the latter. Finkelstein et al. (2014) did not make use of the full IR data available in the HUDF, which also explains the larger uncertainties in their faint-end slope constraints. The UV luminosity density of galaxies at $z \sim 7$ determined in the present work is clearly larger than previous determinations, owing to a steep faint-end slope and a very faint integration limit. We can now assess whether this UV production is sufficient to ionize the IGM. We use the photon emission rate per unit cosmological comoving volume required to maintain reionization at a given redshift determined in (Madau et al. 1999):

$$\dot{N}(z) = (10^{51.2} \text{ s}^{-1} \text{ Mpc}^{-3}) C_{30} \left(\frac{1+z}{6} \right)^3 \left(\frac{\Omega_b h^2}{0.02} \right)^2 \quad (6)$$

where C_{30} is the clumping factor $C_{HII} = \langle n_{HII}^2 \rangle / \langle n_{HII} \rangle^2$ normalized to $C_{HII} = 30$, and n_{HII} is the mean comoving hydrogen density in the Universe. This is the minimum value for which the ionizing emission balances the recombination rate. We show in Fig. 11 this limit, converted to UV luminosity density (cf. Madau et al. 1999; Bolton & Haehnelt 2007) assuming an escape fraction of the ionizing radiation of $f_{esc} = 20\%$ and three different values for the clumping factor. Assuming a standard value of $C_{HII} = 3$ for the clumping factor of the IGM (Pawlik et al. 2009; Finlator et al. 2012) our determination of galaxy UV luminosity density at $z \sim 7$ is sufficient to maintain reionization. We also show that an escape fraction as low as $f_{esc} = 10 - 15\%$ (green hatched region) is already sufficient to ionize the IGM at $z \sim 7$. At the same redshift, the ionizing emissivity constraints of Bouwens et al. (2015a) are close to our UV density constraints.

With current observations we are not able to put better constraints on the faint-end of the UV LF at $z \sim 8$, hence on the ionizing emissivity of galaxies at this redshift. The conclusions from the deep blank fields are still affected by large uncertainties that prevent any strong claims regarding the contribution of galaxies to the ionizing budget at $z \sim 8$ (Bouwens et al. 2015b; Finkelstein et al. 2014). Future *HST* observations of the remaining HFF clusters might add better constraints on the UV luminosity density at $z \sim 8$ with additional highly magnified faint galaxies.

6. SUMMARY

Combining *HST* observations of lensing clusters and parallel fields of the Hubble Frontier Fields program, we computed the galaxy UV luminosity function between $z = 6$ and $z = 8$. We assembled a large sample of about 250 galaxy candidates in this redshift range using the Lyman break photometric selection. In the cluster cores, we corrected the deep *HST* images for intra-cluster light and bright cluster galaxies light using a median filtering before object detection, while performing the photometry

in the original images. This technique helps the detection of faint galaxies contaminated by cluster light but do not increase significantly the number of galaxies in the cluster center because of the very small volume probed in those regions that have very high magnification.

- Using the latest lensing models produced by the CATS team for the three clusters A2744 (Jauzac et al. 2015), MACS0416 (Jauzac et al. 2014), and MACS0717 (Limousin et al. 2015), we have performed completeness simulations in the source plane that take into account all lensing effects. Thanks to the lensing magnification, the completeness function extends down to fainter magnitudes in the cluster fields than in the parallels. On the other hand, the total survey volume goes from $\sim 4 \text{ arcmin}^2$ in the parallel fields down to $0.6-1 \text{ arcmin}^2$ in lensing fields.
- We computed the UV luminosity function for all individual fields and for the combined sample at $z \sim 7$. The lensing magnification allow us to extend the LF down to an absolute UV magnitude limit of $M_{UV} = -15.25$, which is more than two magnitudes deeper than any study the deep blank fields. Most importantly, we show that the faint-end slope remains very steep at $\alpha = -2.04^{+0.13}_{-0.17}$ at such faint intrinsic luminosity. When excluding MACS0717 cluster, the uncertainties on the faint-end slope decrease to $\sigma_\alpha \sim 0.1$. The best Schechter fit yields a characteristic magnitude of $M^* = -20.89^{+0.60}_{-0.72}$ and $\text{Log}(\phi^*) = -3.54^{+0.48}_{-0.45}$. This is in good agreement with most of the recent results in blank fields (Bouwens et al. 2015b; Finkelstein et al. 2014), which were limited to $M_{UV} = -17.5$ and early results of Atek et al. (2015).
- Our determination of the UV LF at $z \sim 8$ does not reach beyond $M_{UV} = -18$ because we detect only two galaxies in the two fainter magnitude bins. Albeit with large uncertainties due to small number statistics and incompleteness uncertainties, the currently determined LF points confirm the results of previous studies.
- On observational grounds, we integrate the UV LF down to a magnitude limit of $M_{UV} = -15$ and find $\text{Log}(\rho_{UV}) = 26.2 \pm 0.13 \text{ erg s}^{-1} \text{ Hz}^{-1} \text{ Mpc}^{-3}$. Assuming standard values for the ionizing conversion factor ξ_{ion} and the clumping factor $C_{HII} = 3$, the ionizing budget of galaxies would be sufficient to maintain the IGM ionized by $z \sim 7$ provided the ionizing escape fraction from galaxies is greater than 10%. Future observations of the HFF program will allow us to decrease the uncertainties on the UV luminosity density at $z \sim 7$ and perhaps improve the constraints on the UV luminosity density at $z \sim 8$, which are mostly based on blank field observations.

With the help of gravitational lensing we have produced the best constraints currently available on the UV luminosity function and the contribution of galaxies to the IGM reionization at $z \sim 7$. We clearly demonstrate

here the great potential and the feasibility of peering into the early universe through these cosmic telescopes. The remaining HFF clusters will certainly help improve even more our constraints on the UV LF at $z \sim 7$ and in particular at $z \sim 8$ to better assess the role of galaxies in the cosmic reionization process. Our results also show the great promise of future programs targeting lensing fields and pave the way to observing programs with the *James Webb Space Telescope* or the *Wide-Field Infrared Survey Telescope (WFIRST)*, which are scheduled for launch in the near future.

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